Physics 224
The Interstellar Medium

Lecture #8

- Part I: Few last points on the "curve of growth"
- Part II: Ionization Processes
- Part III: Recombination Processes

$$\tau_\nu = \frac{\pi e^2}{m_e c} f_{lu} \; N_l \; \phi_\nu$$

"oscillator strength" related to Einstein coeff

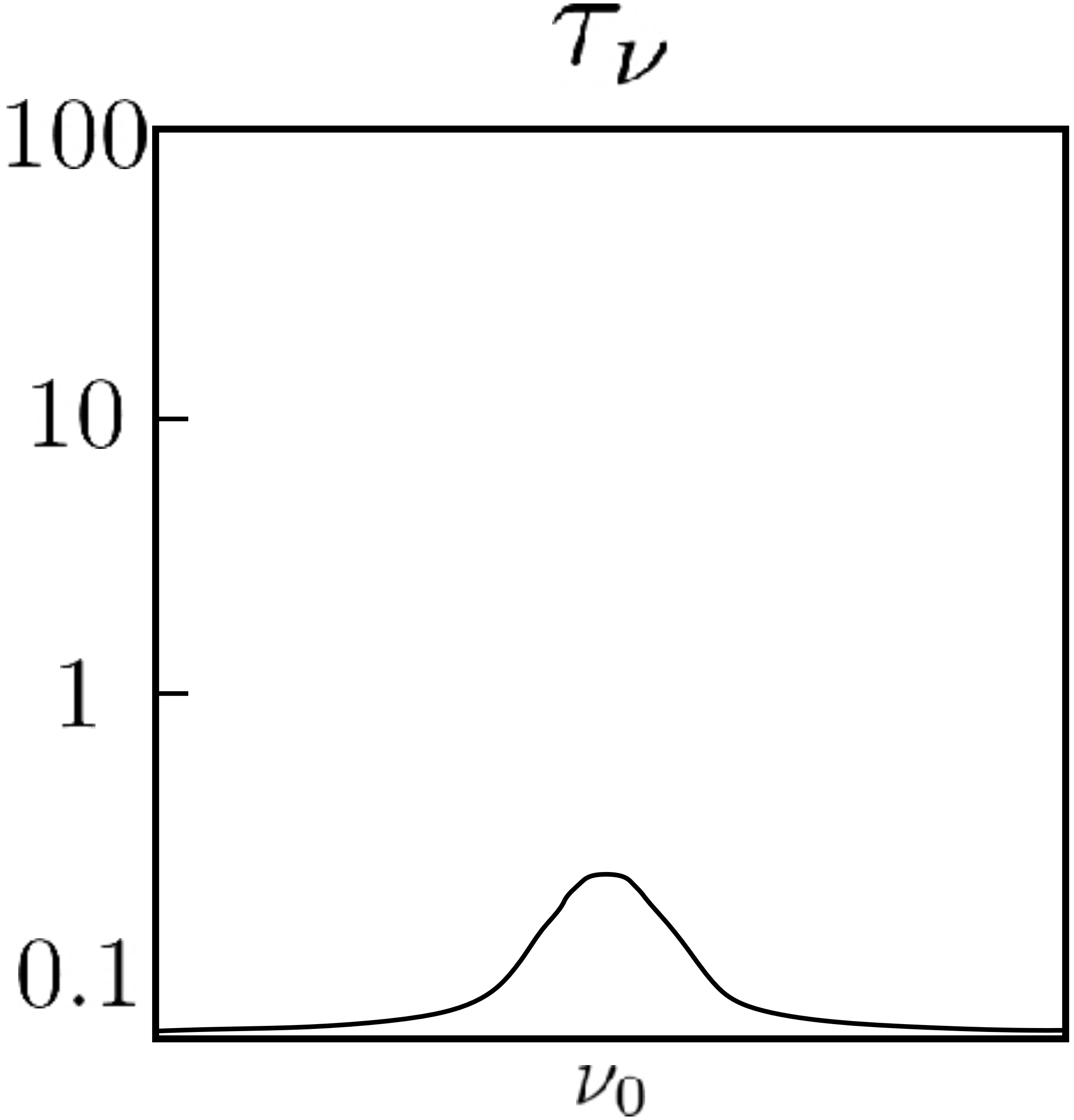
column density of absorbers

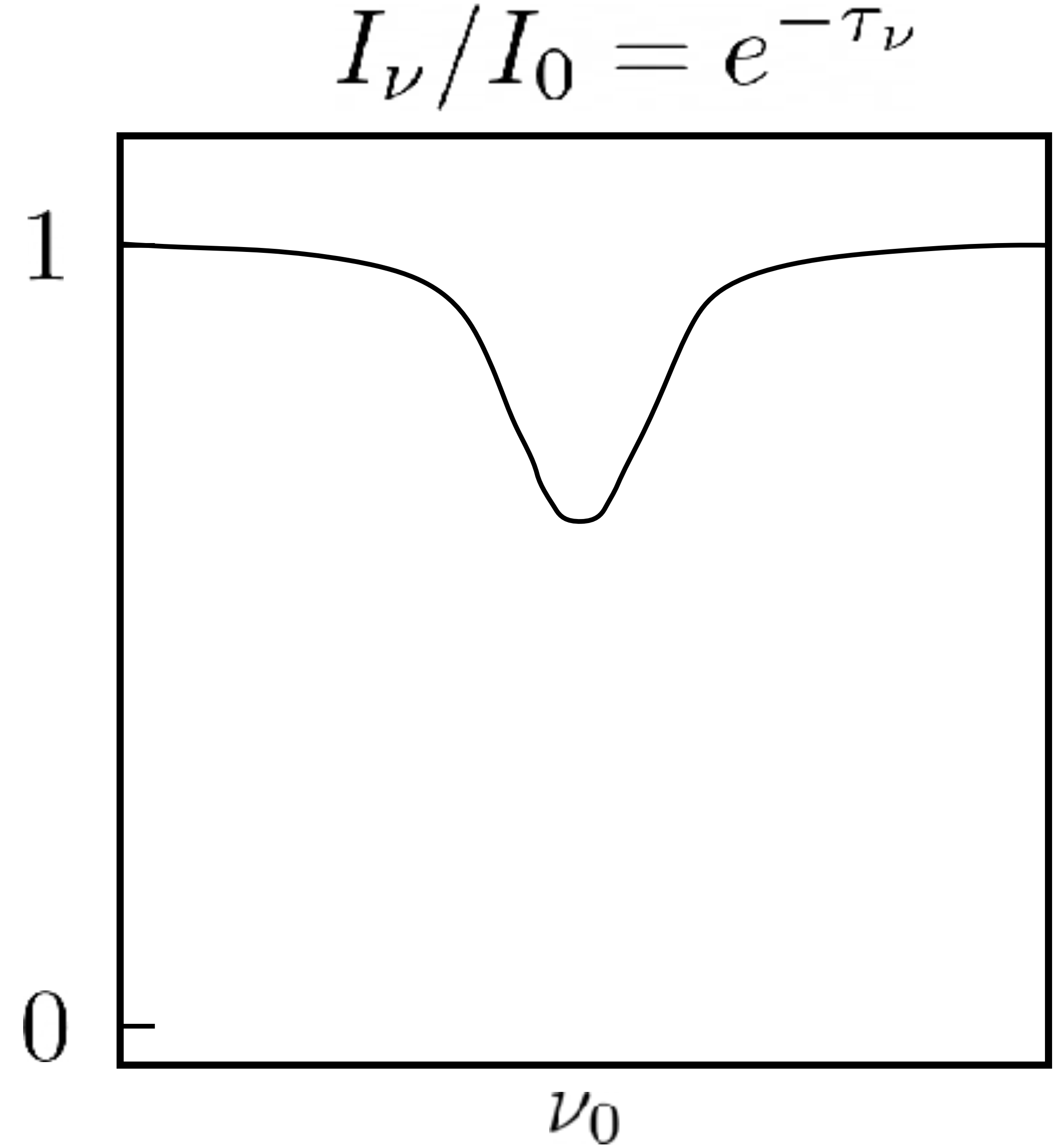
where line profile is:

$$\phi_\nu = \frac{1}{\sqrt{2\pi\sigma_v^2}} \int_{-\infty}^{\infty} e^{-v^2/2\sigma_v^2} \frac{4\gamma_{u\ell}}{16\pi^2(\nu - (1 - v/c)\nu_{u\ell})^2 + \gamma_{u\ell}^2} dv$$

*note no analytic formula for solution of integral

The Curve of Growth





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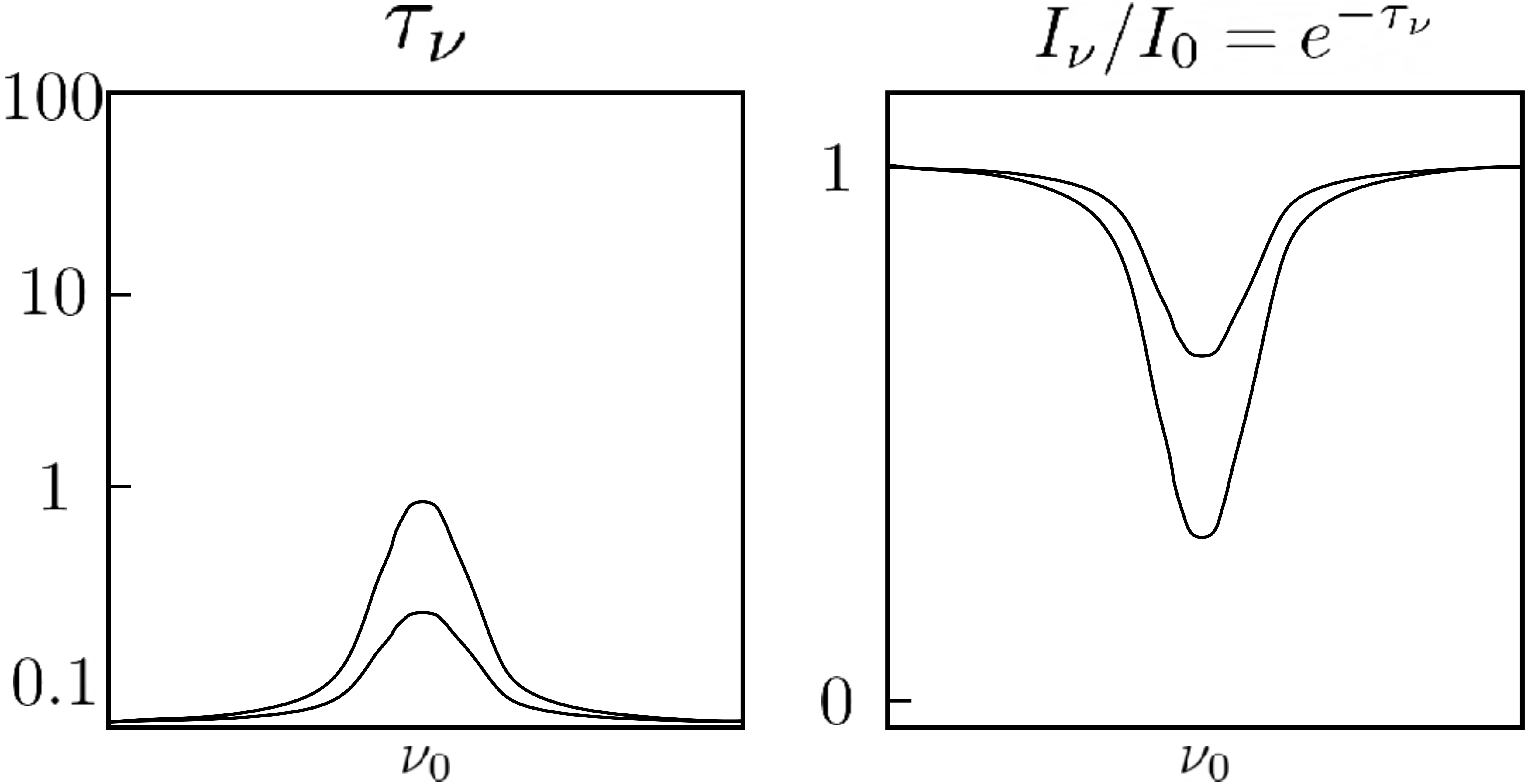

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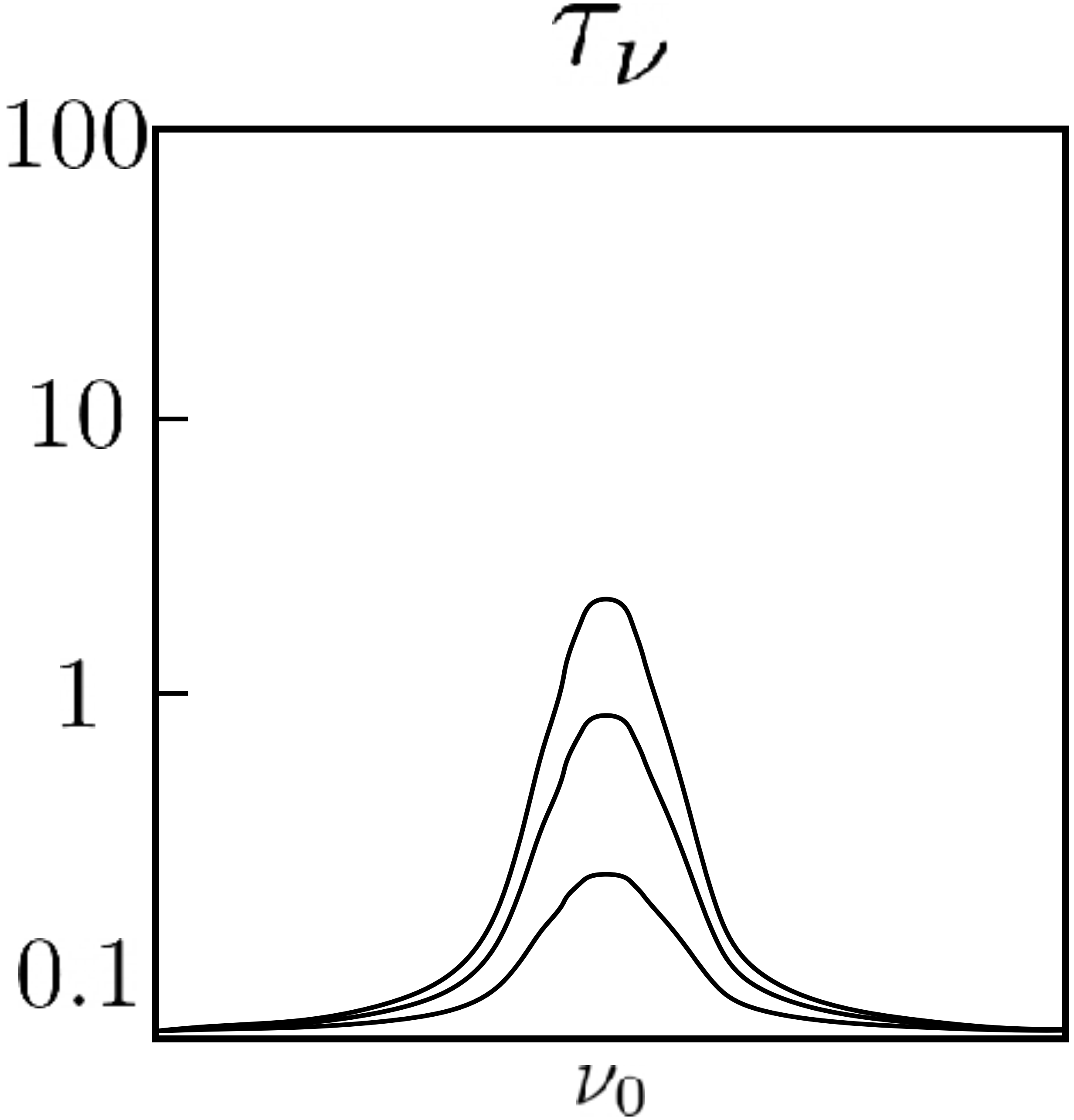

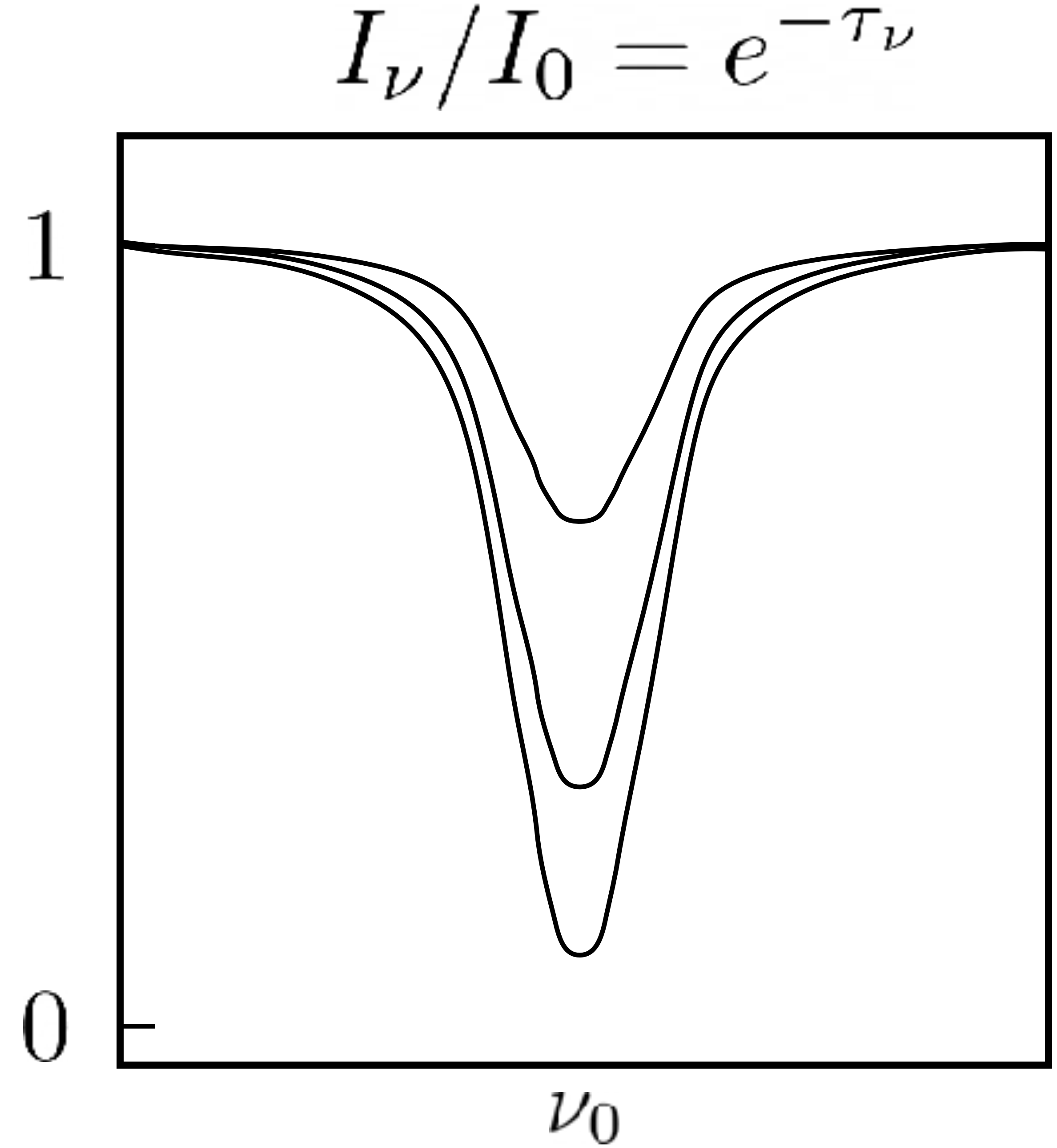

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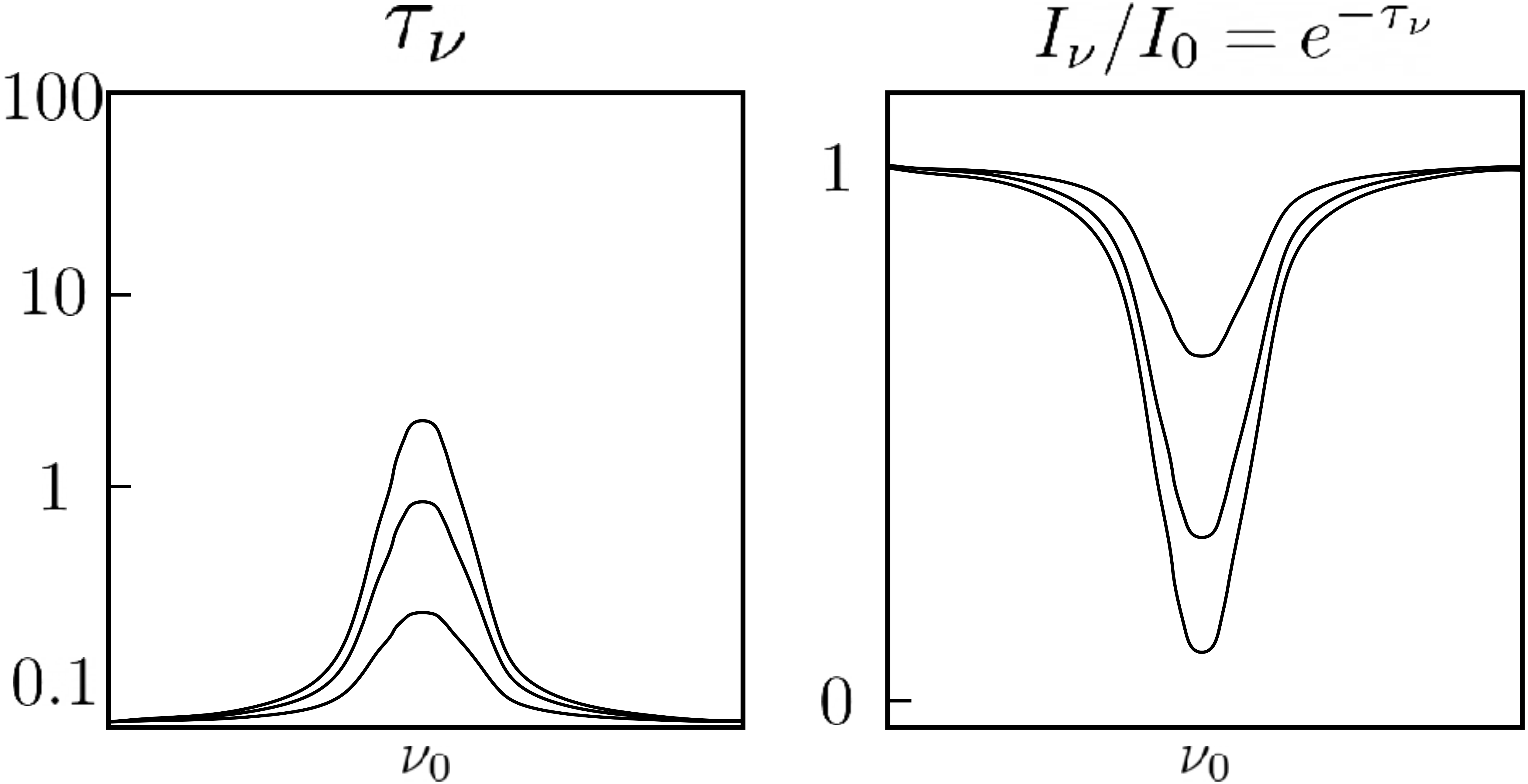
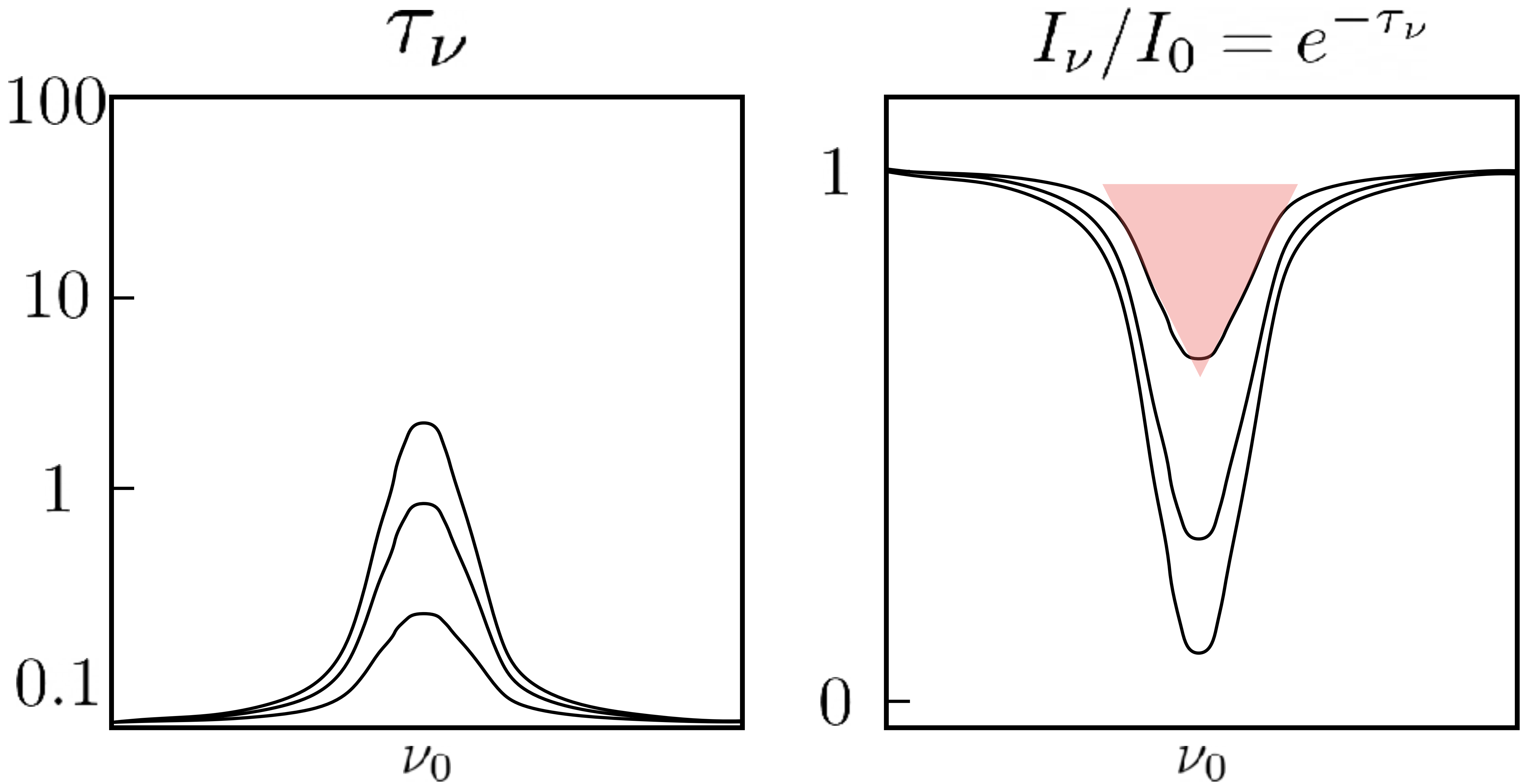


as optical depth at line center becomes >1, $I_{\nu,0}/I_0 \rightarrow 0$

The Curve of Growth



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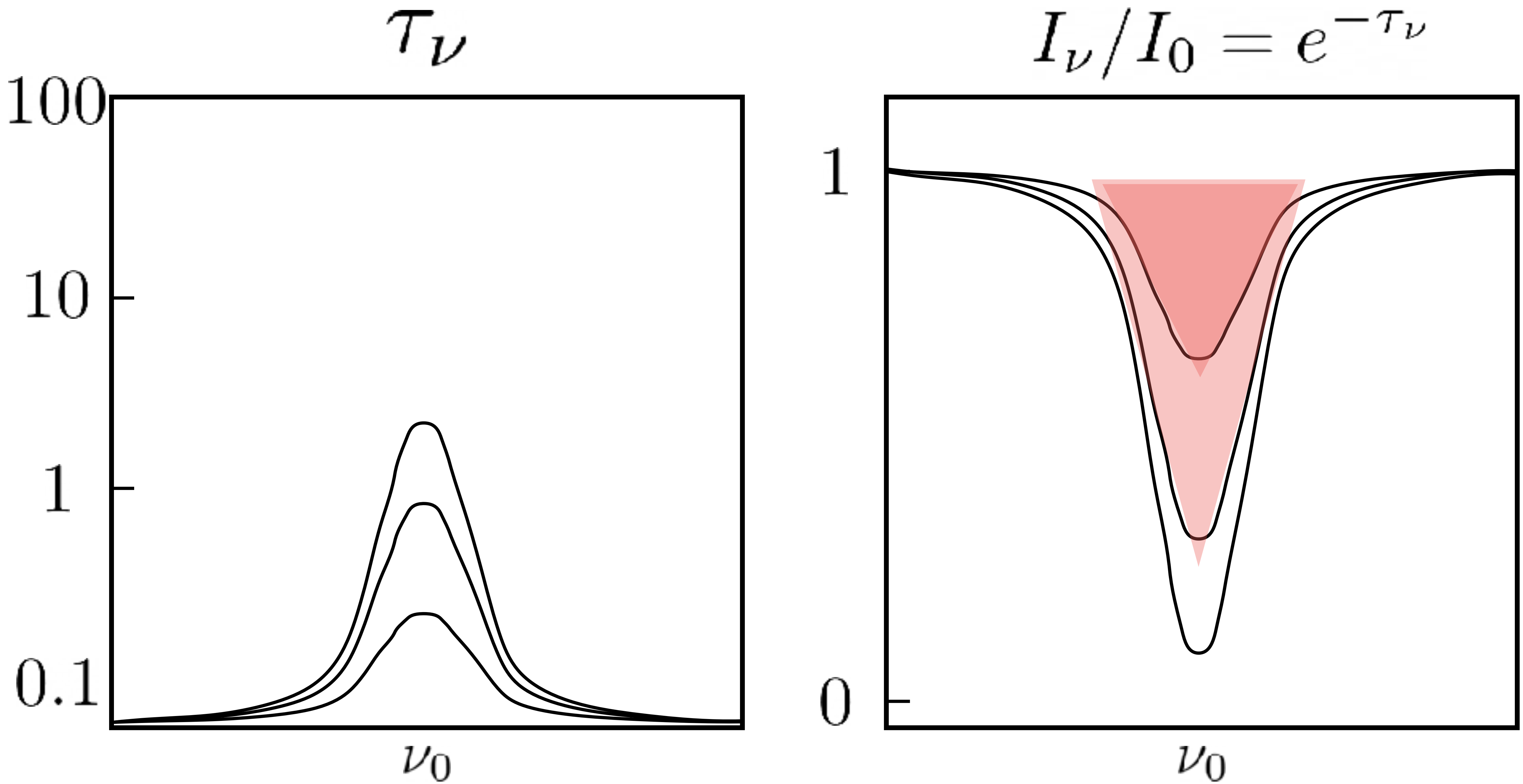

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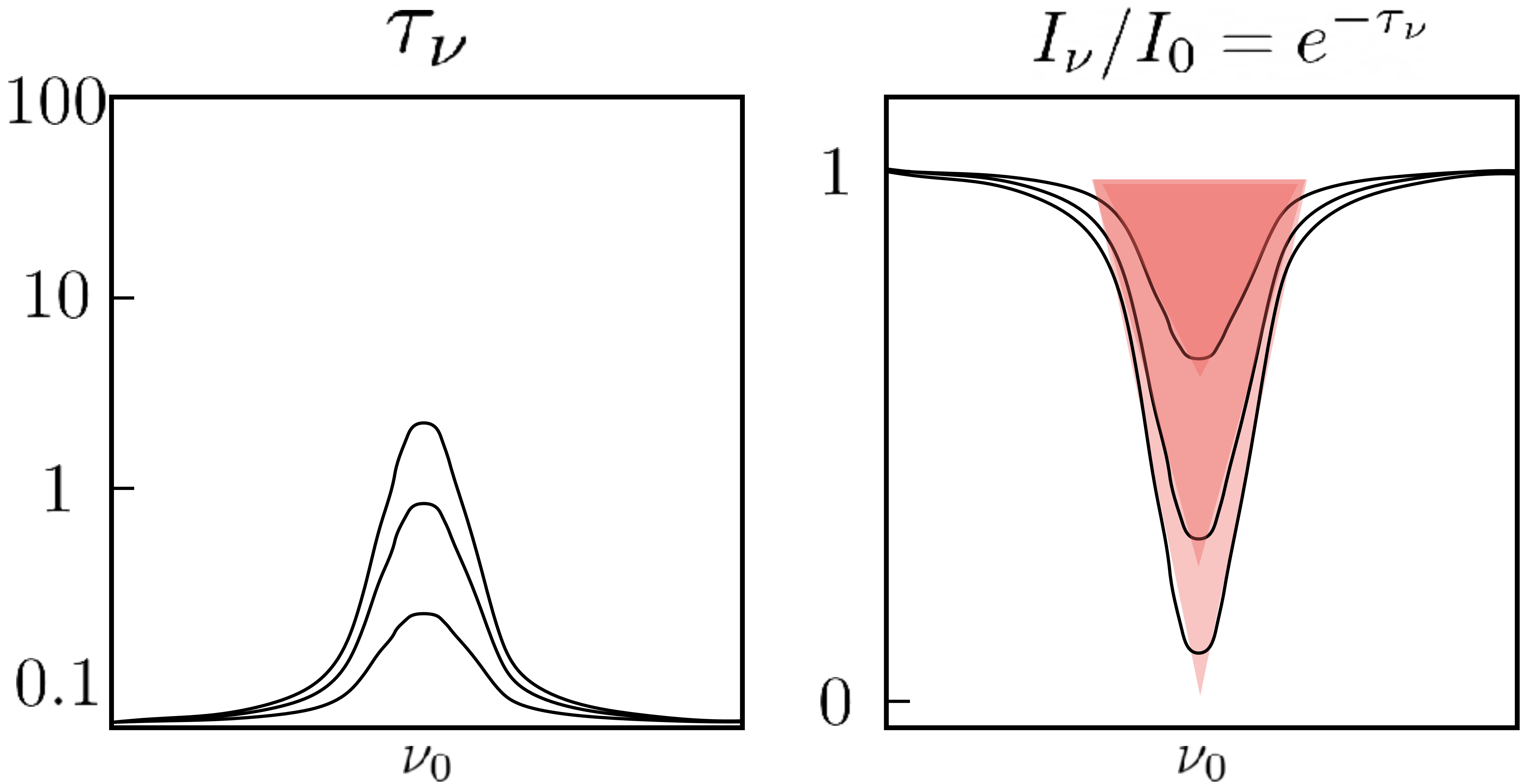

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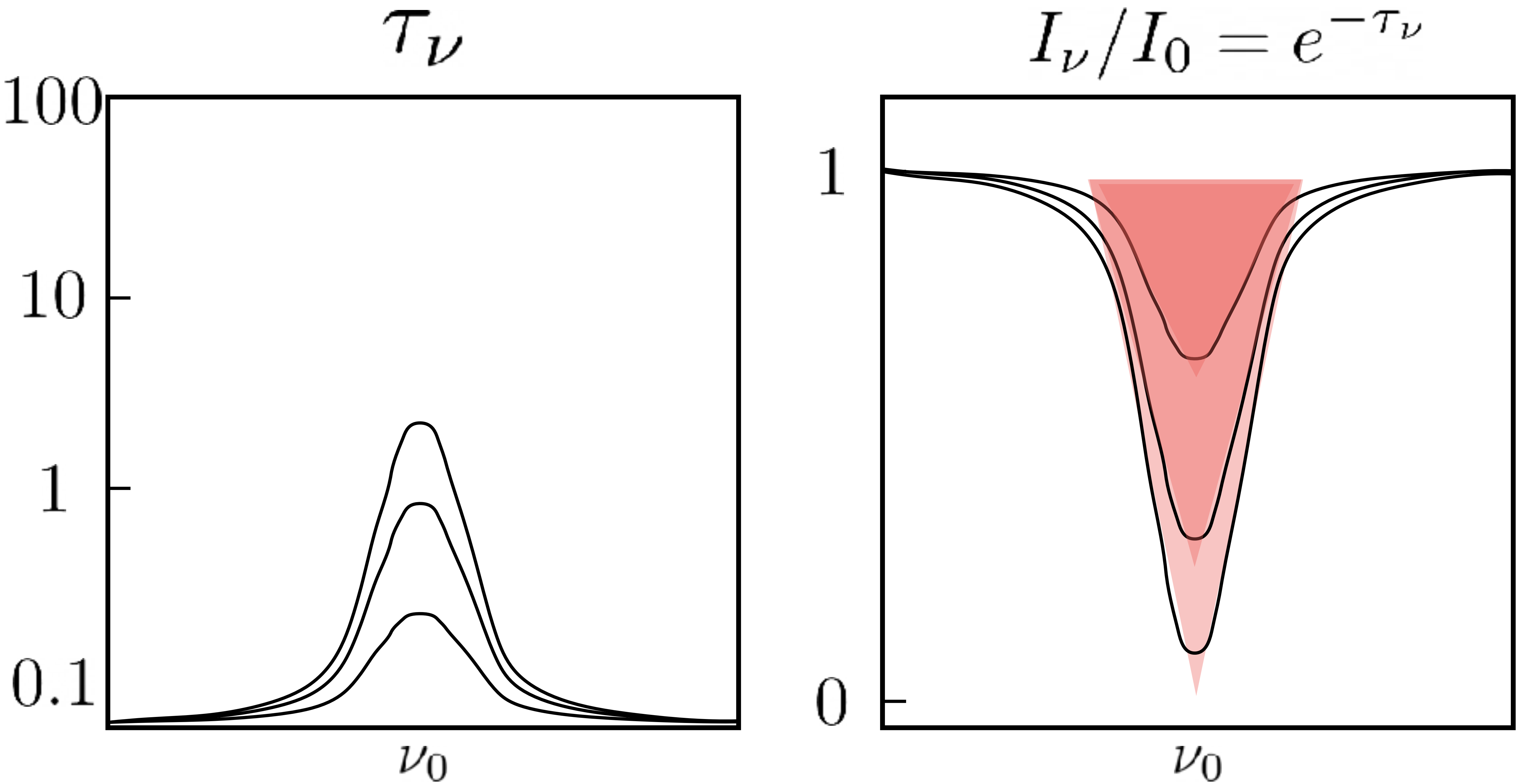
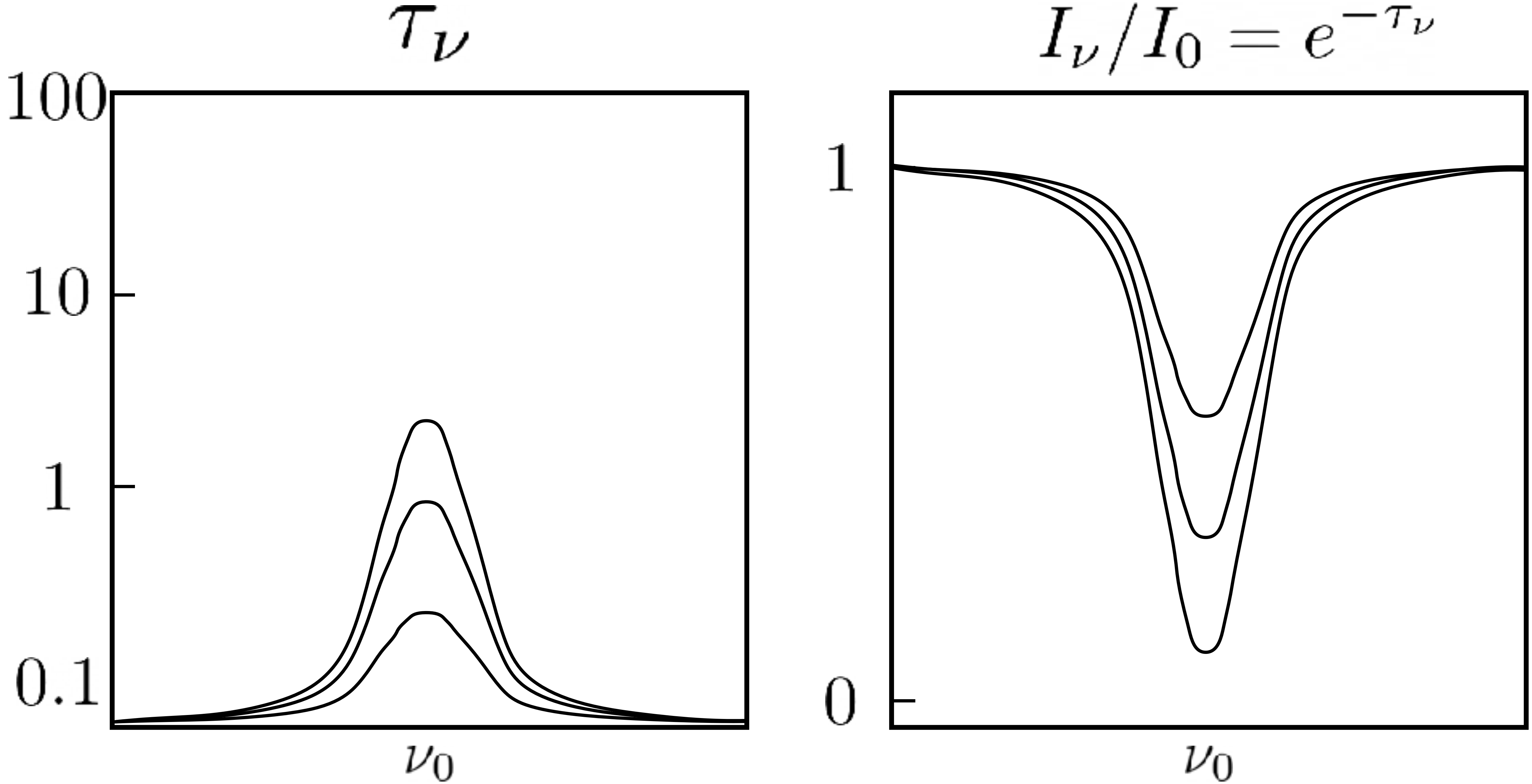


W~1/2 b τ_0 ∝ N

as optical depth at line center becomes >1, $I_{\nu,0}/I_0 \rightarrow 0$

The Curve of Growth



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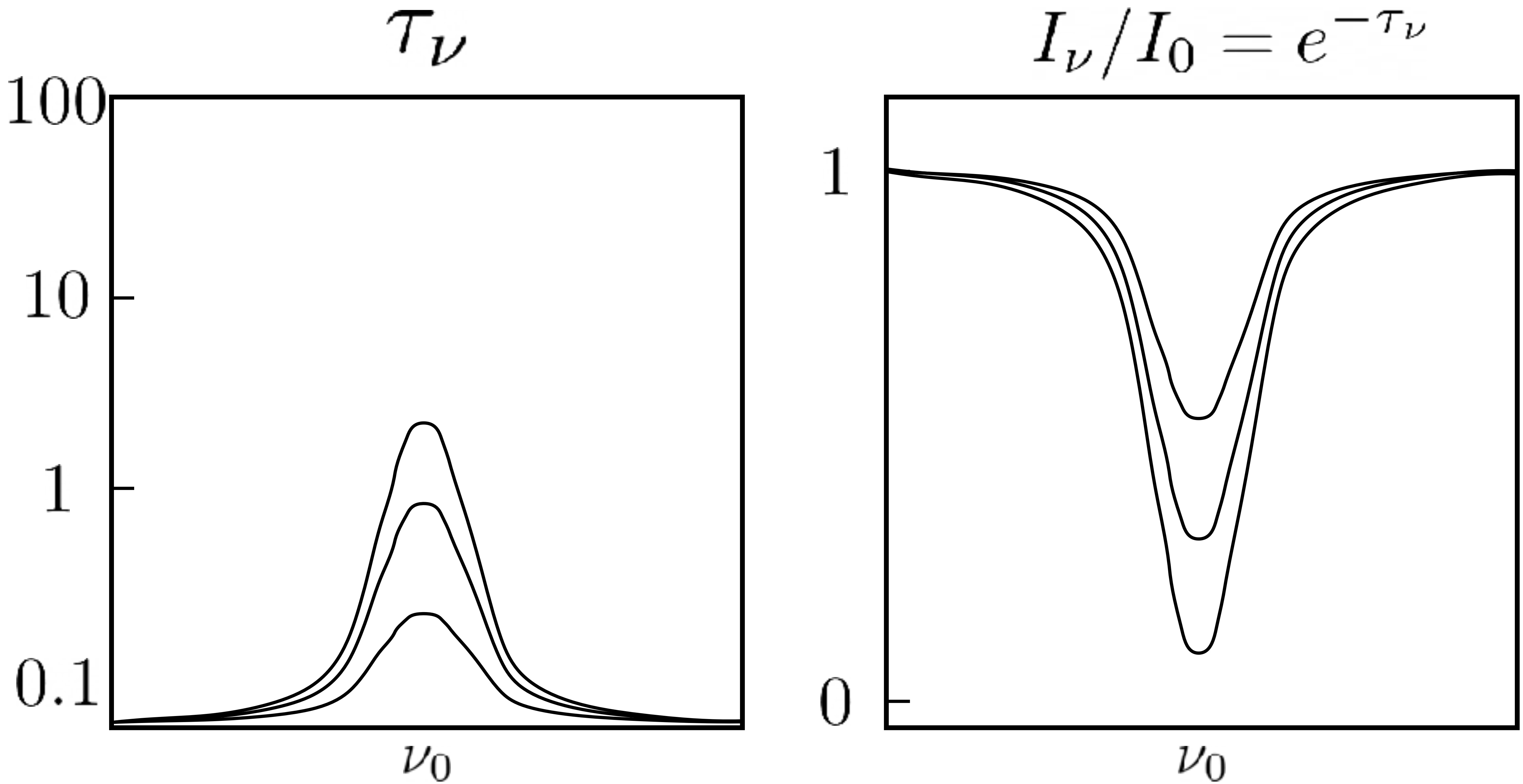
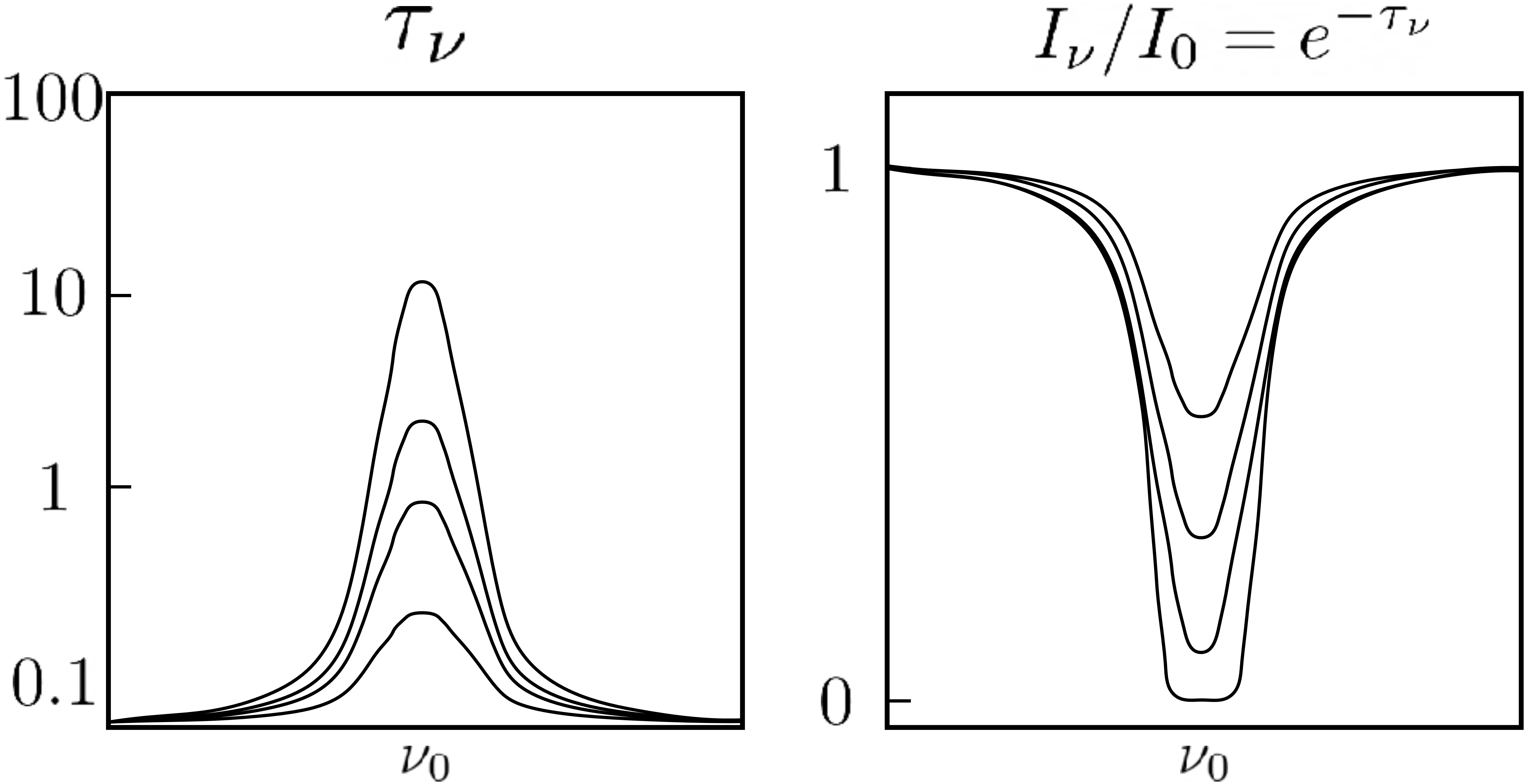


next, optical depth exceeds 1 across the line core

The Curve of Growth



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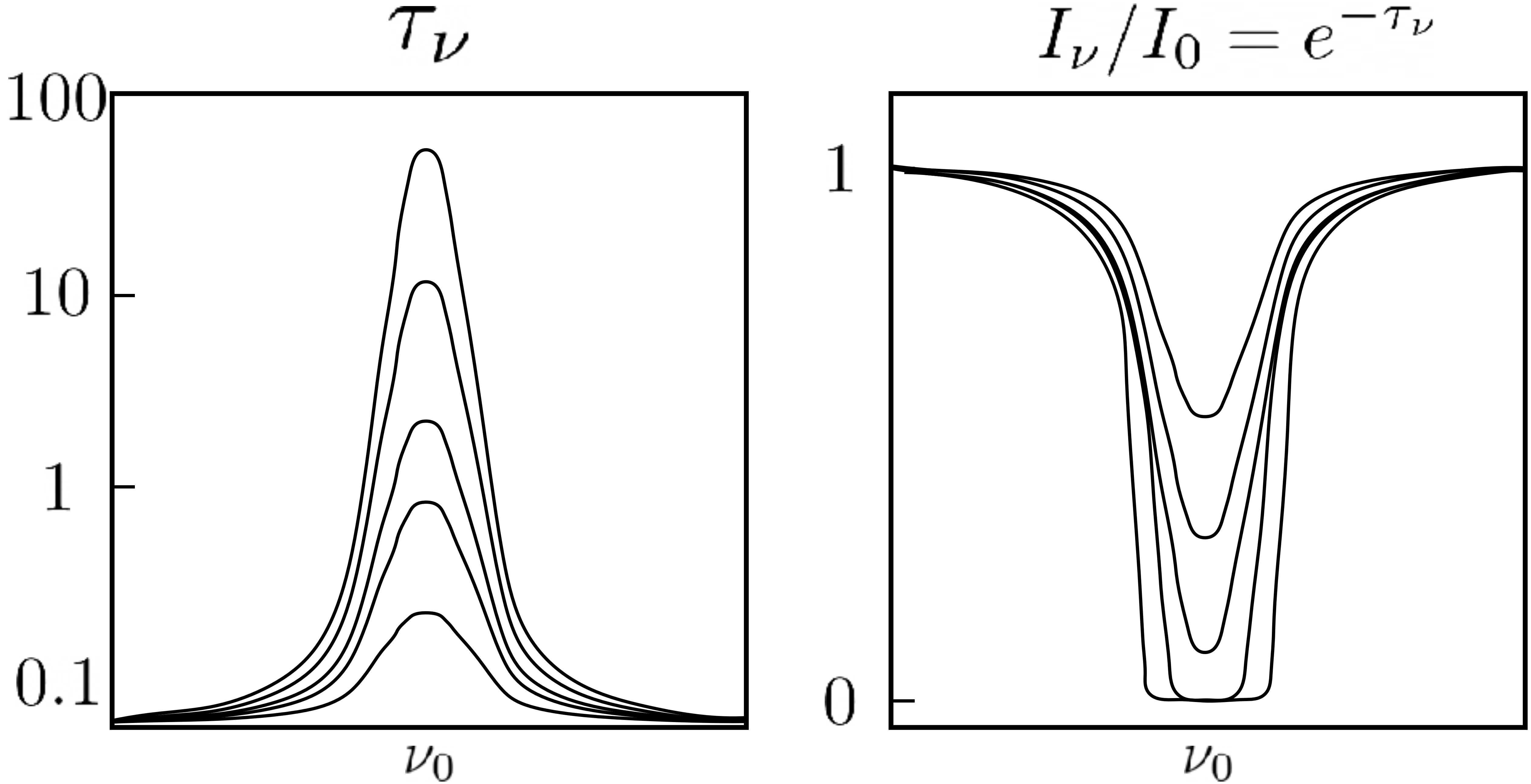

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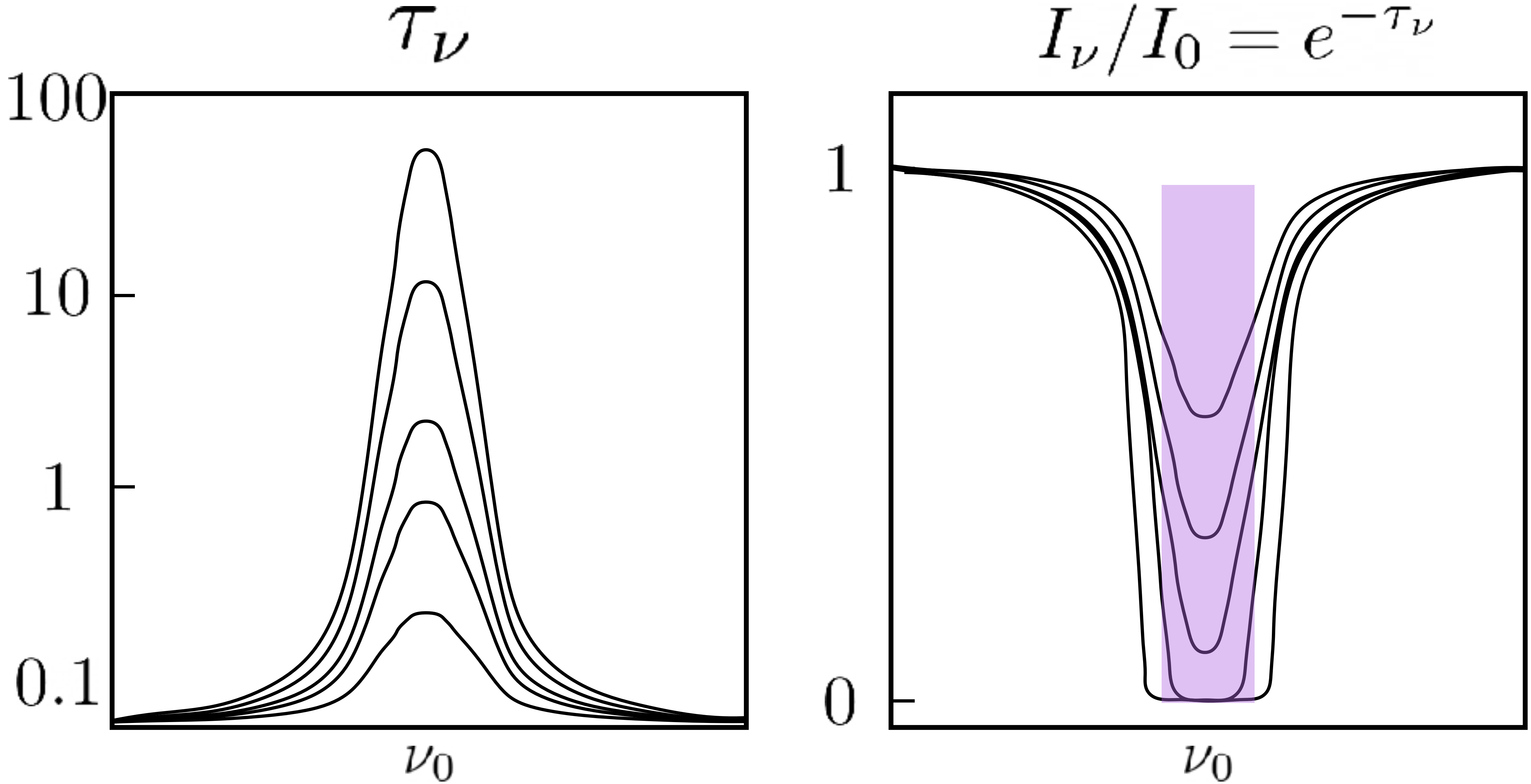

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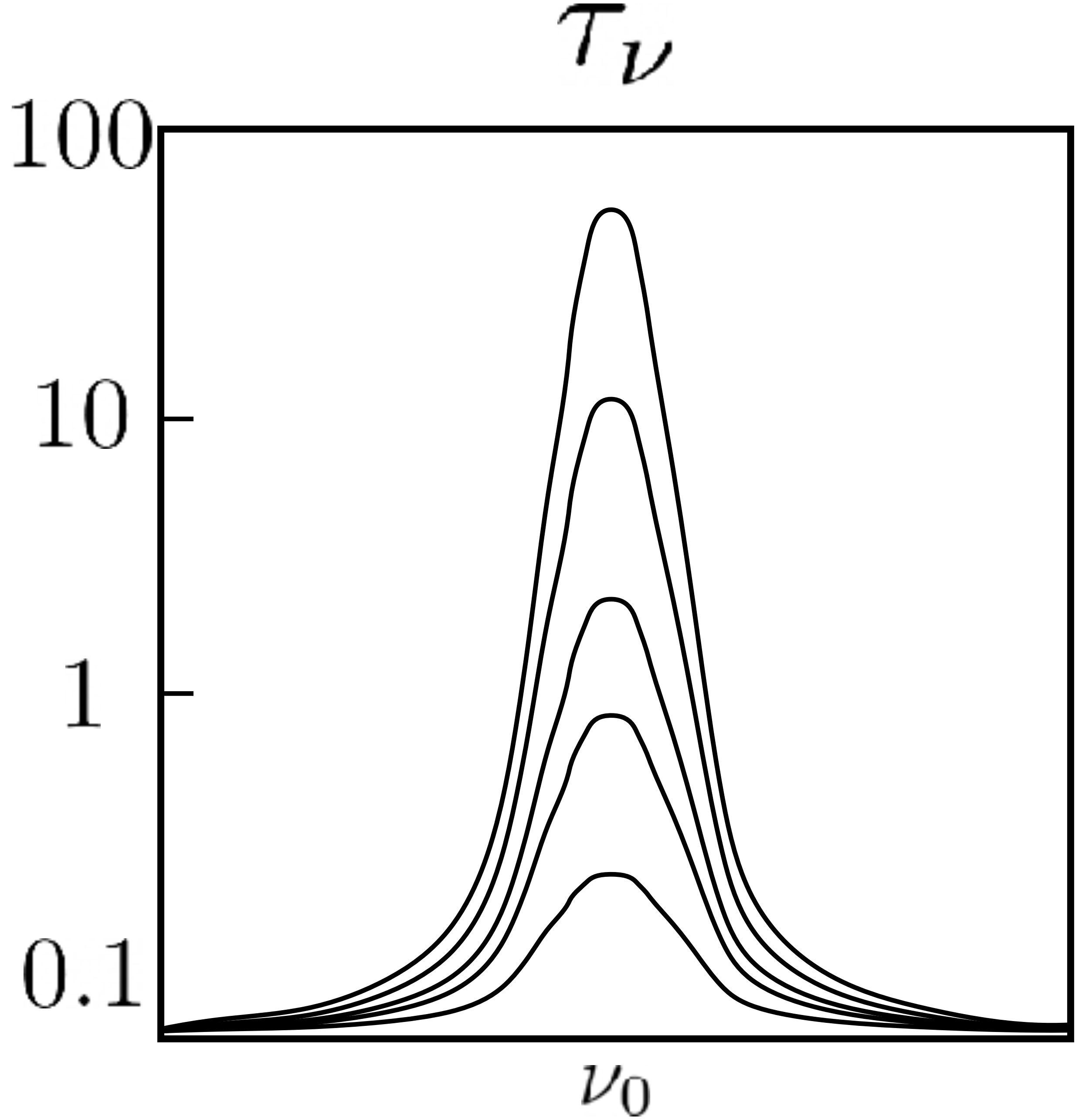

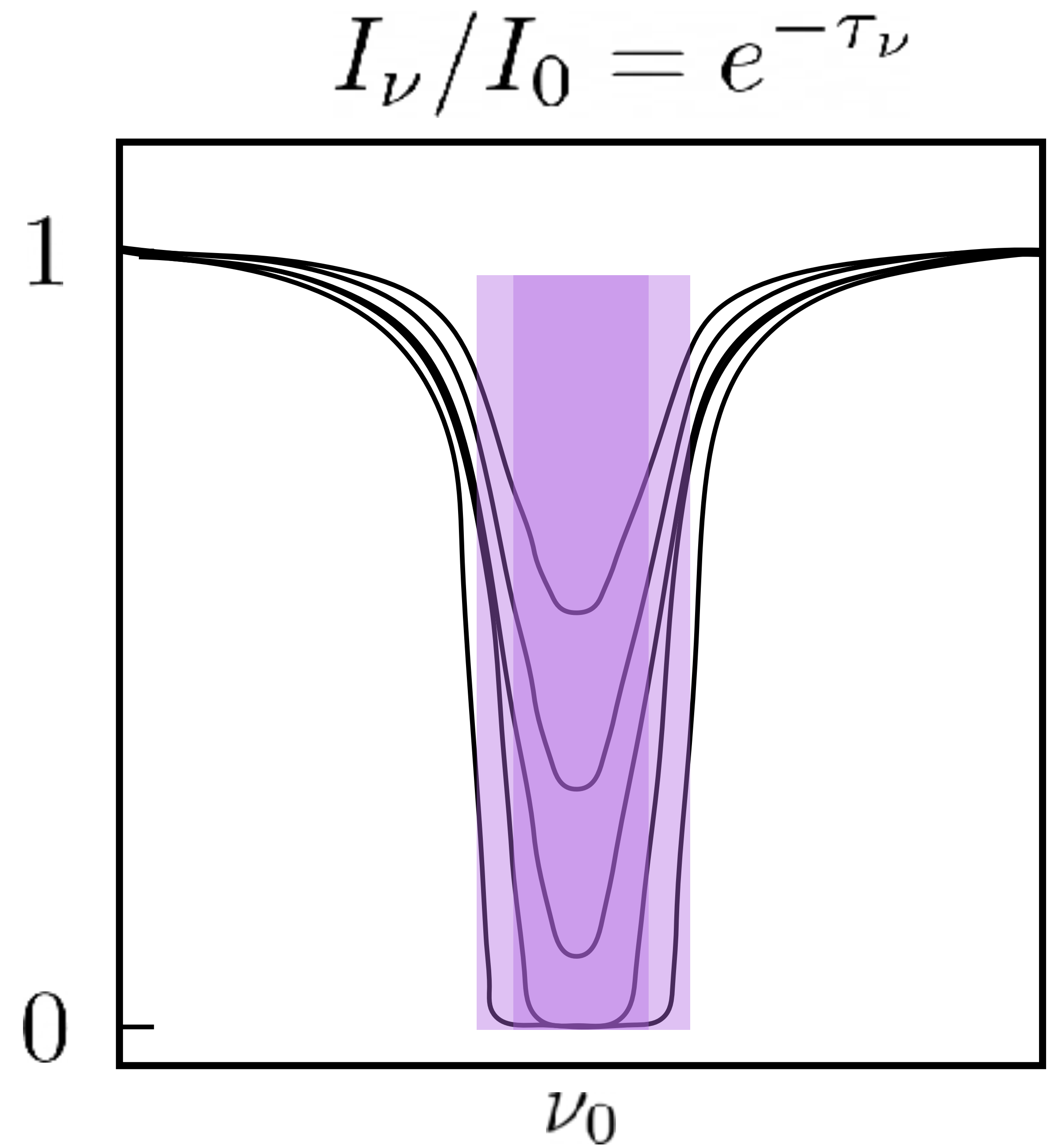

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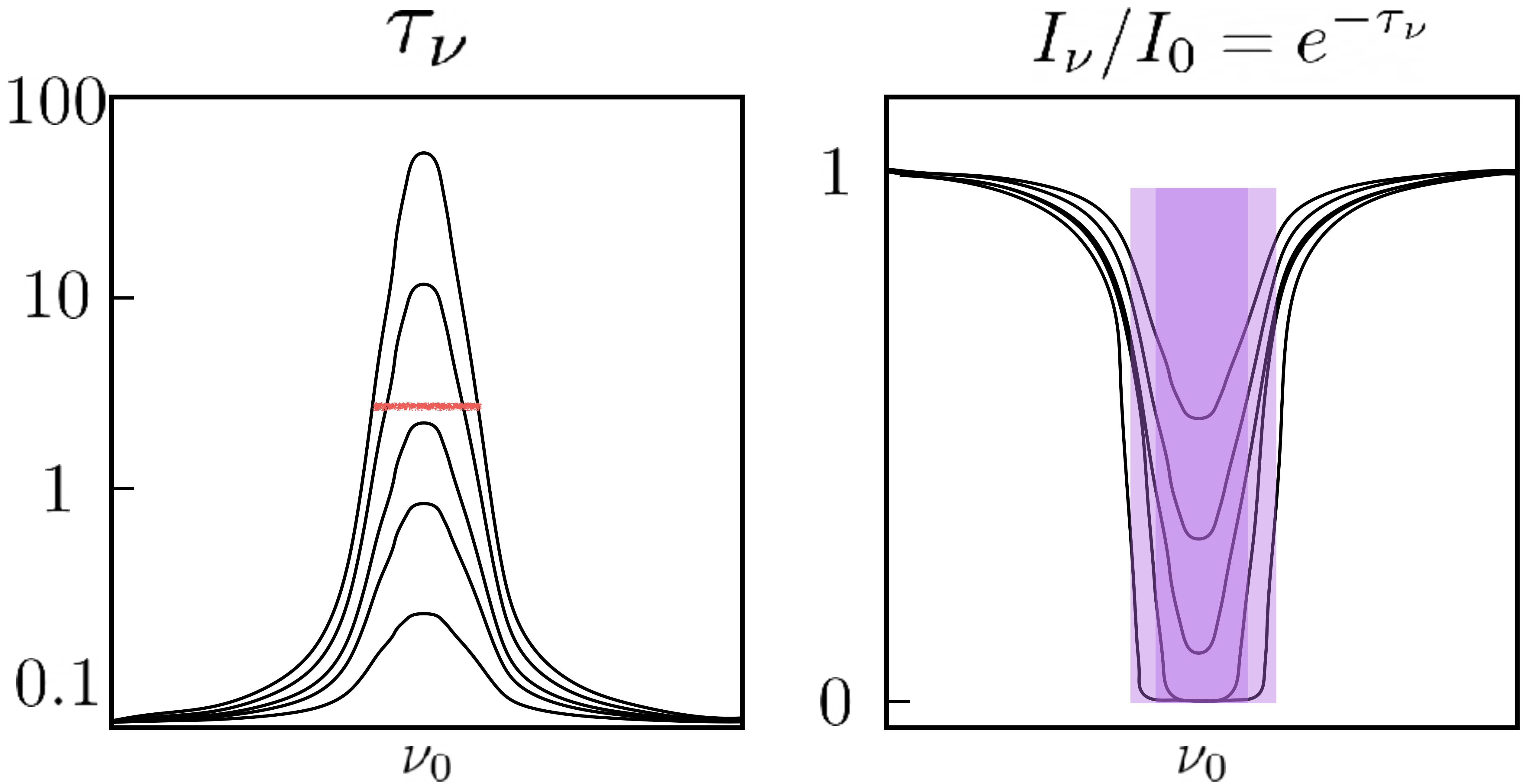

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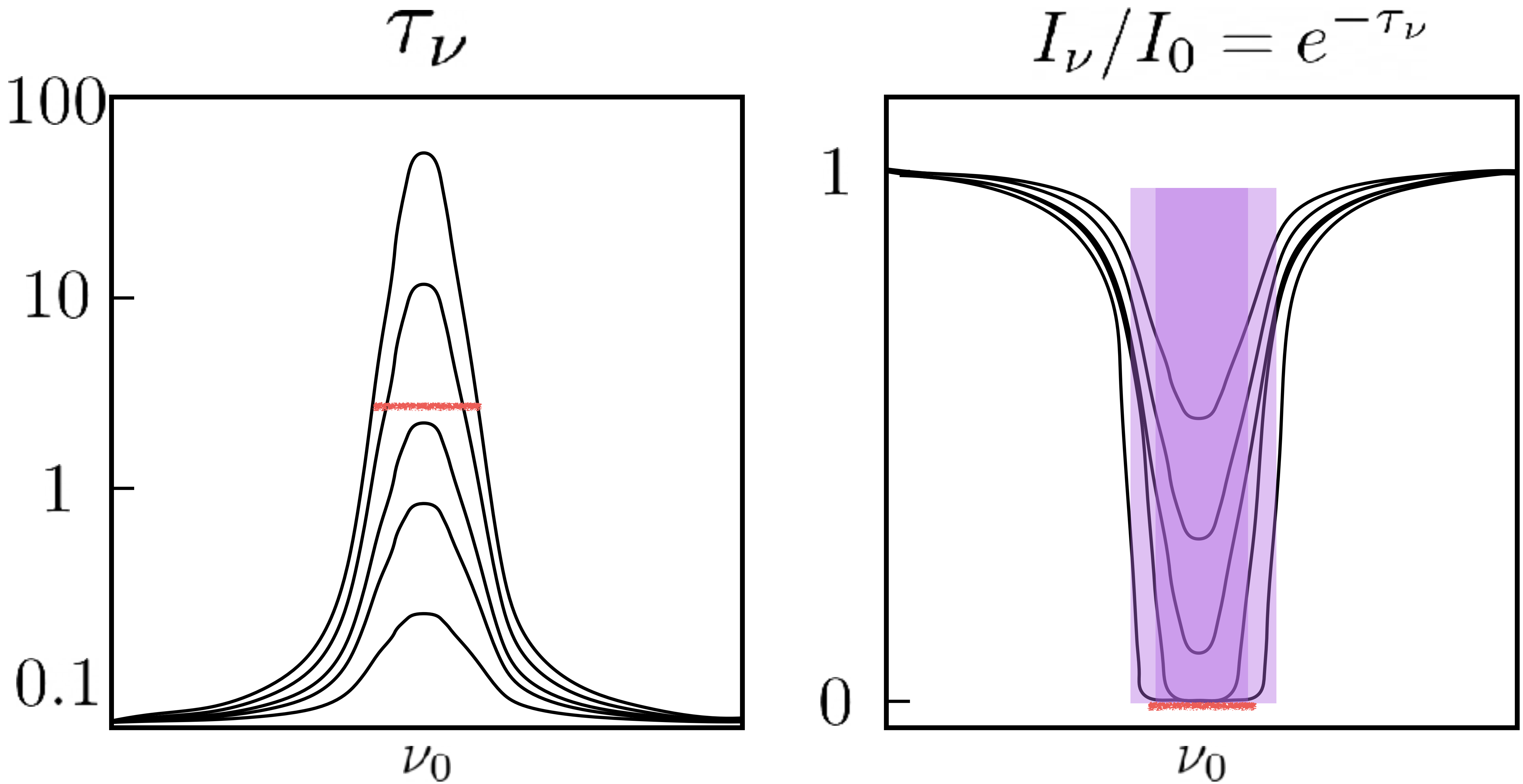

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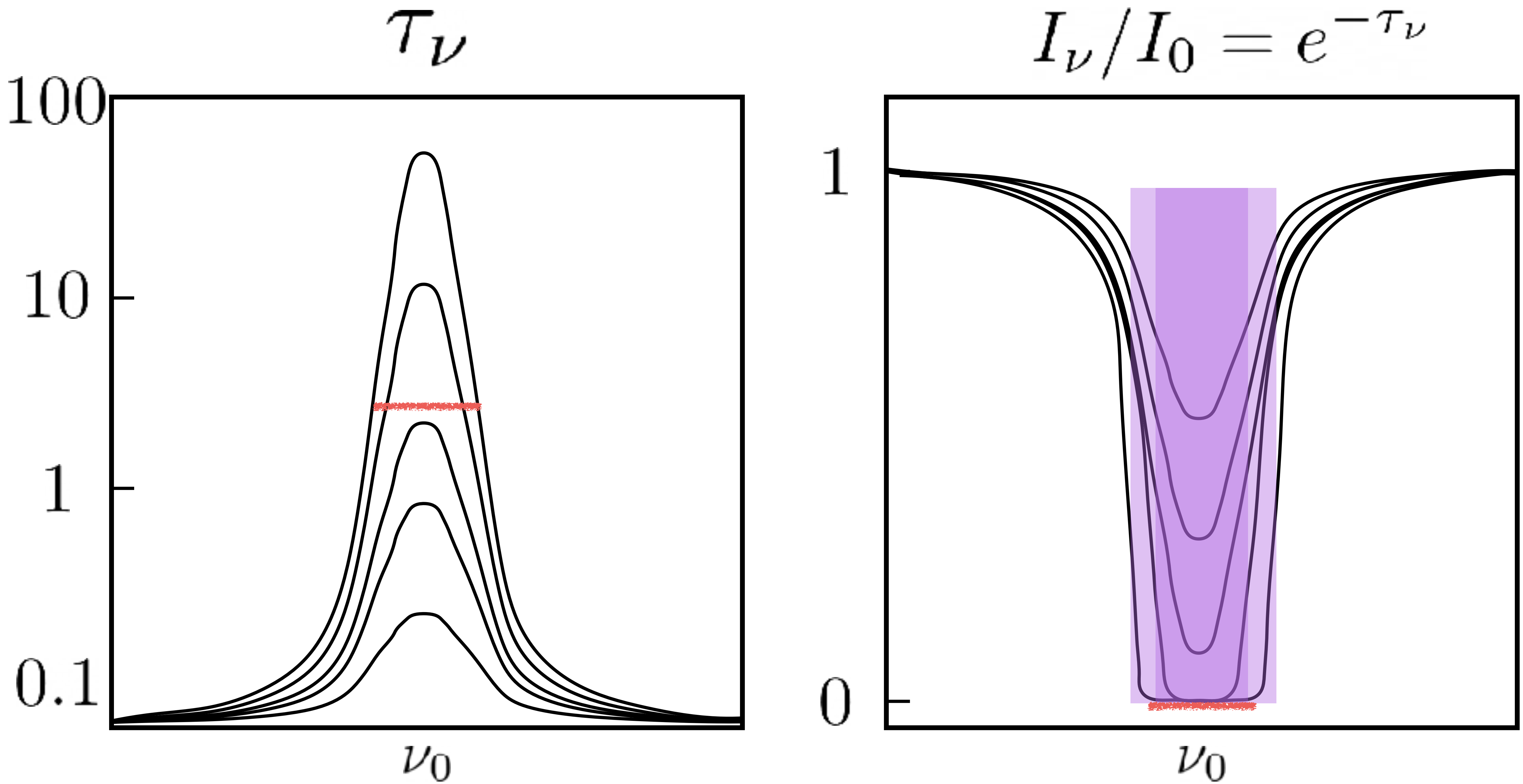
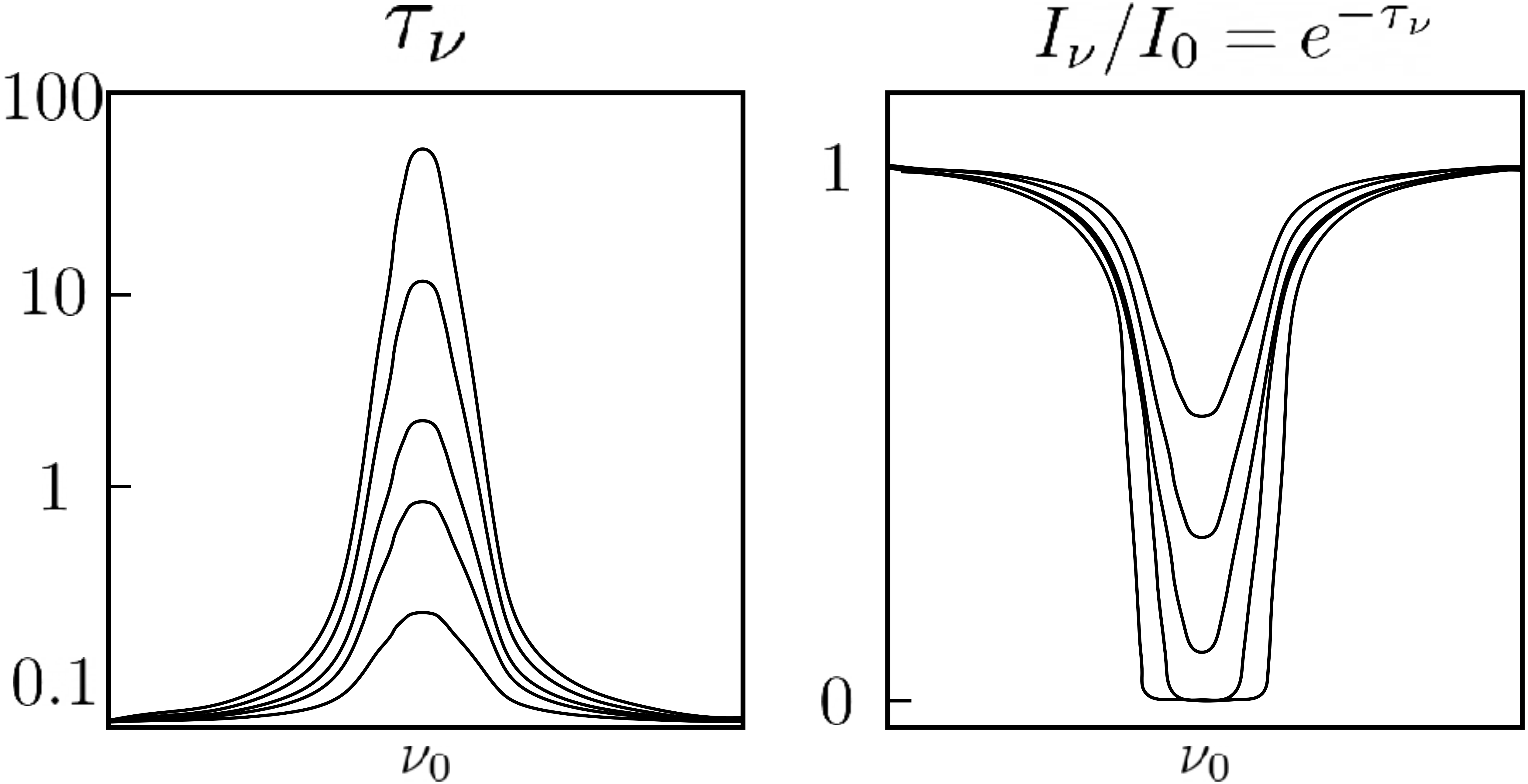


W~b ln(N/b)$^{0.5}$

next, optical depth exceeds 1 across the line core

The Curve of Growth



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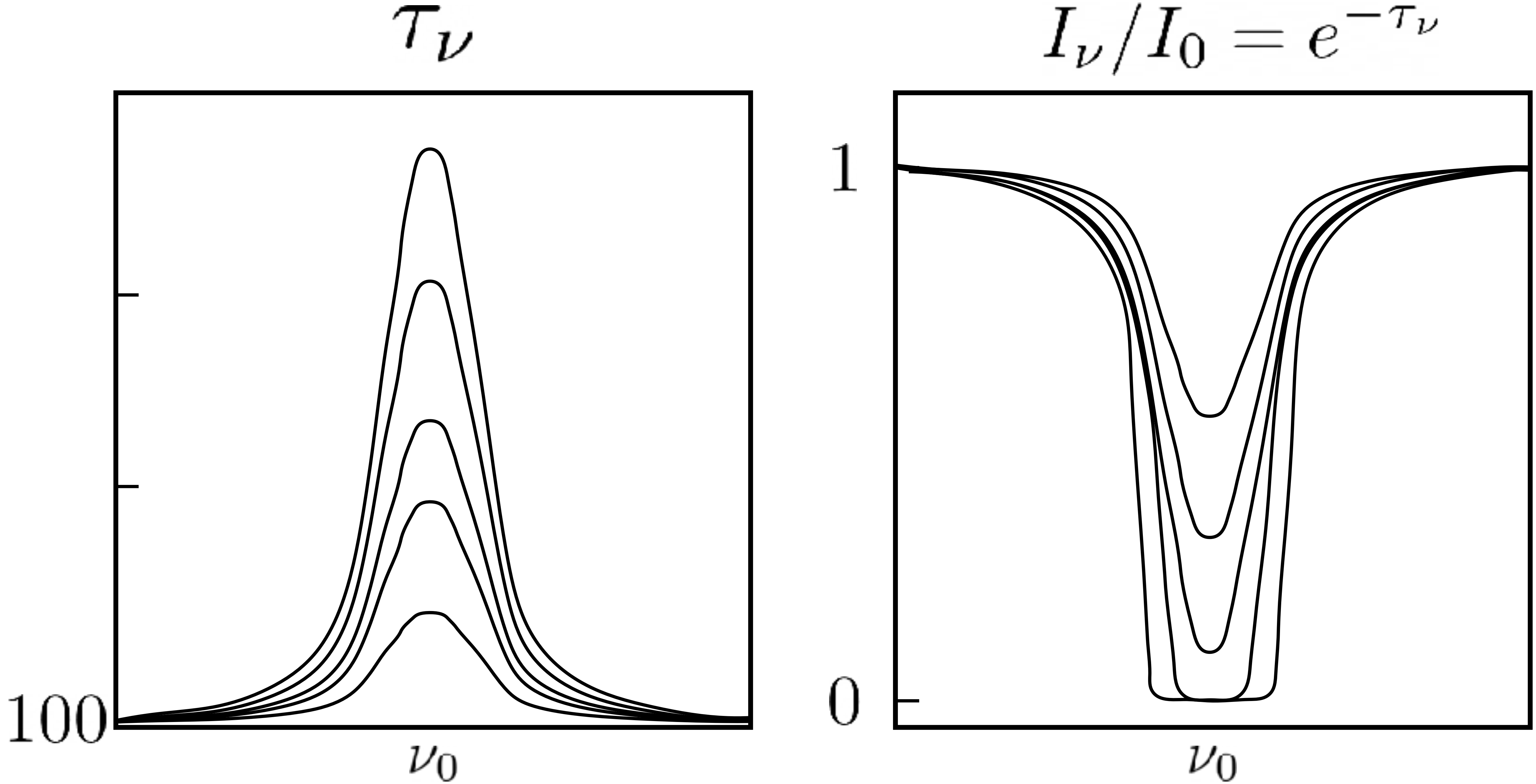

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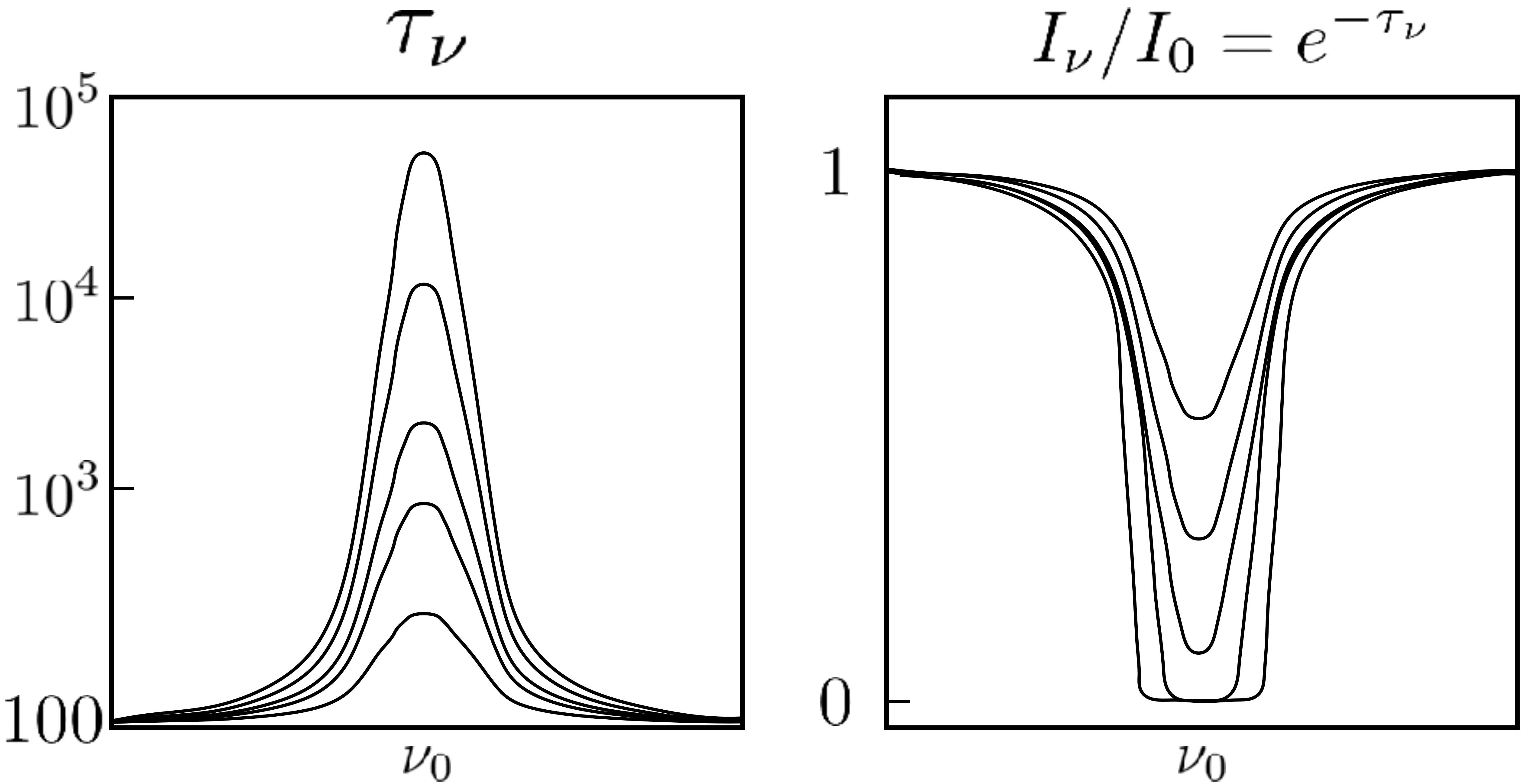
The Curve of Growth



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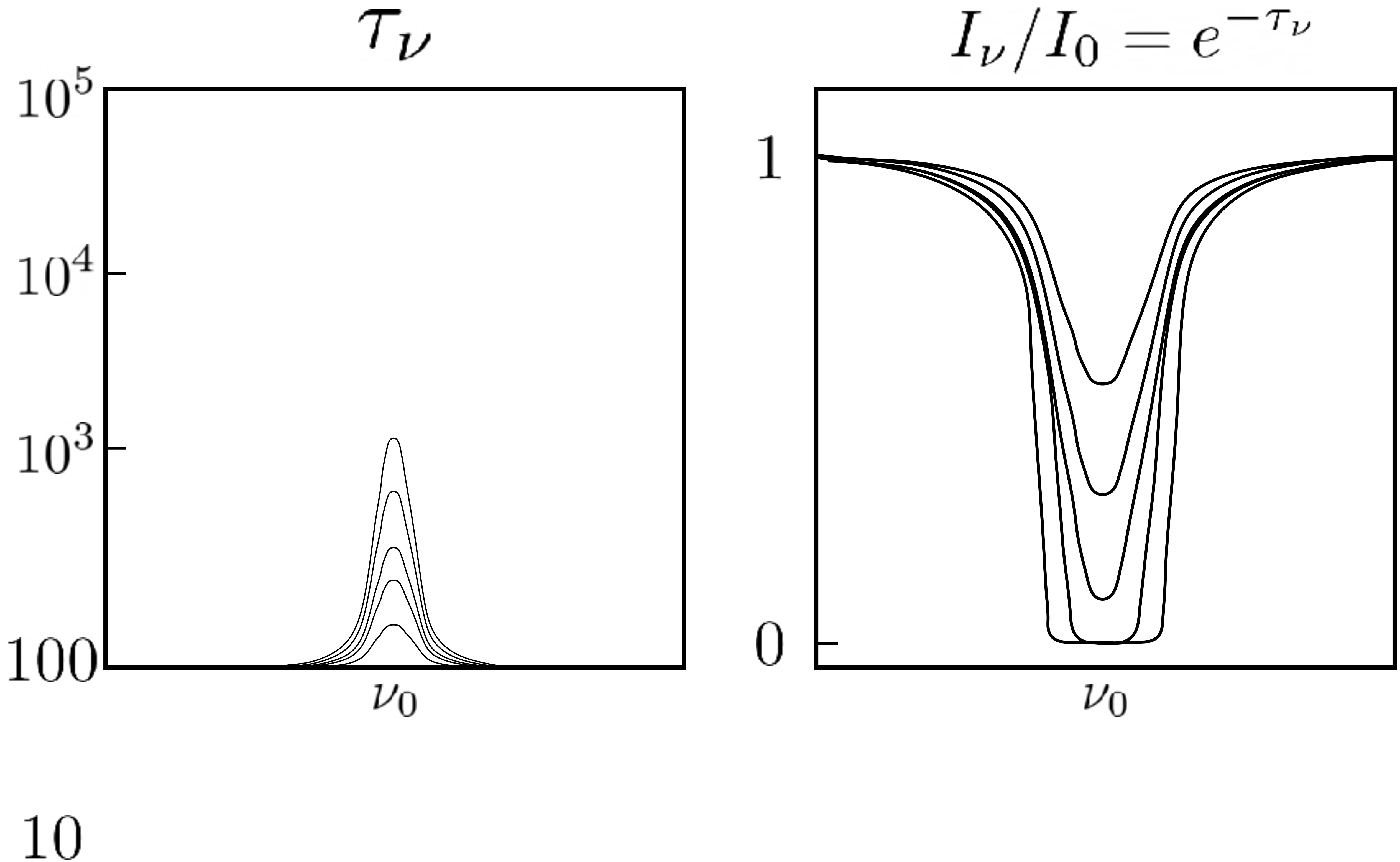
The Curve of Growth



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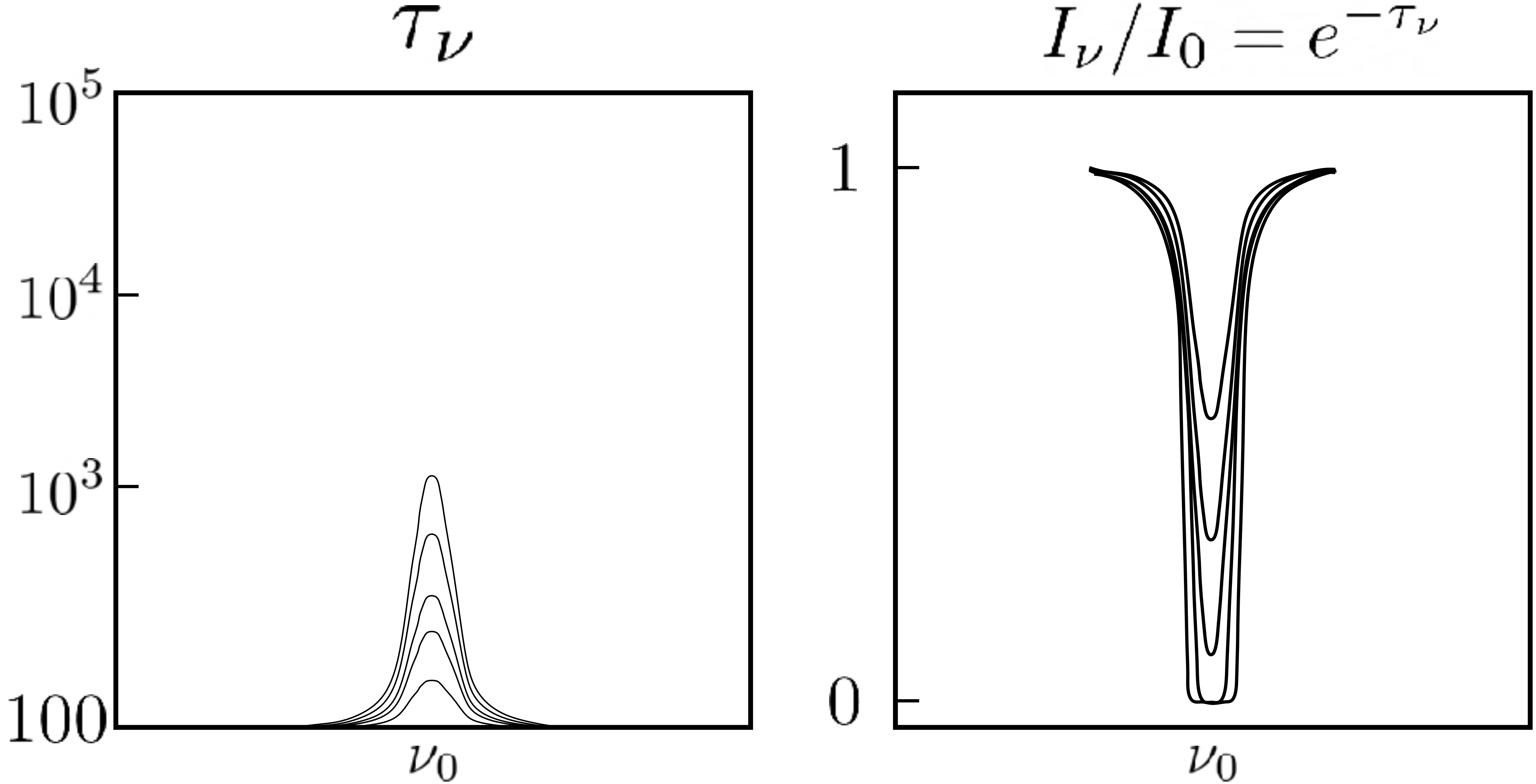

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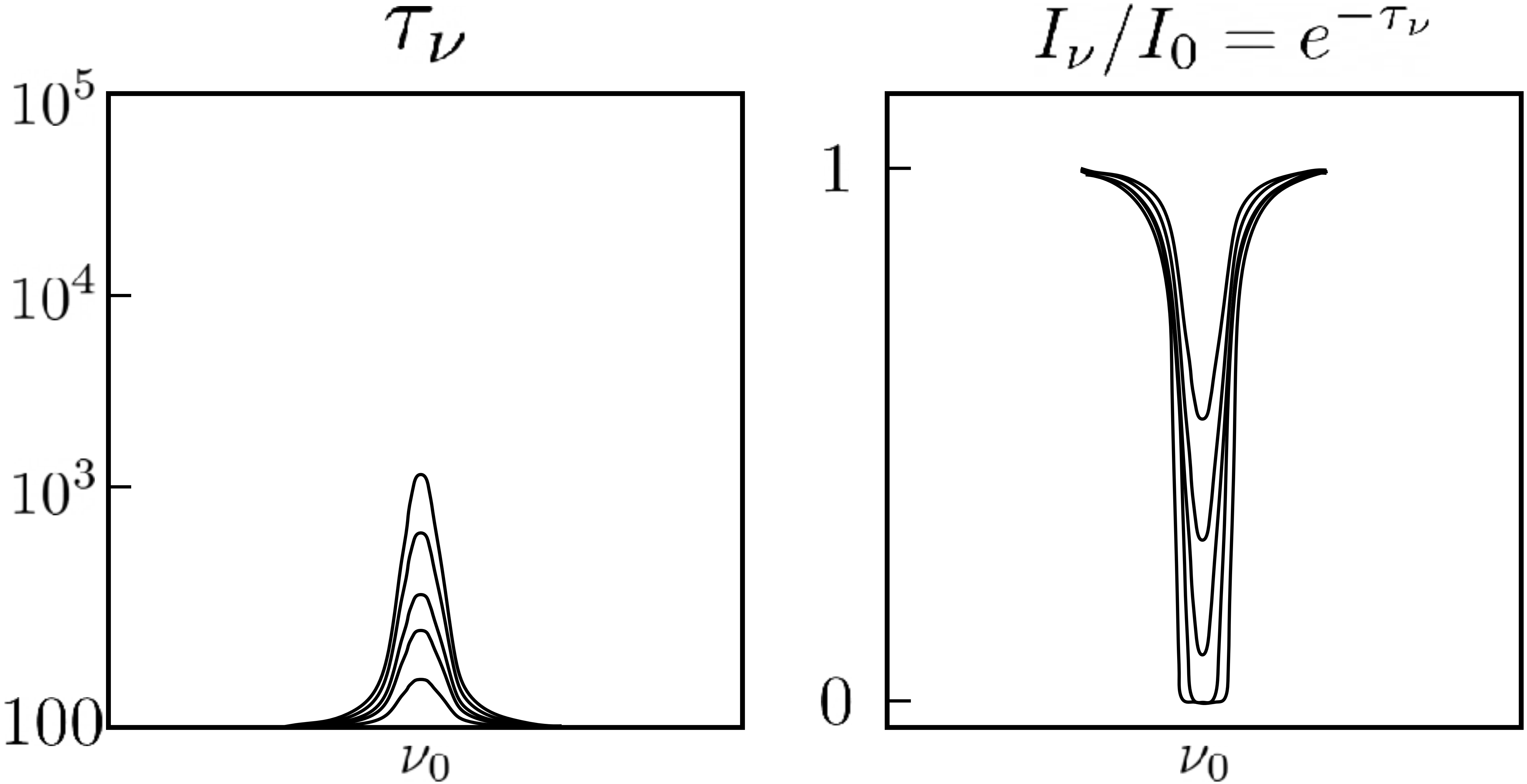
The Curve of Growth



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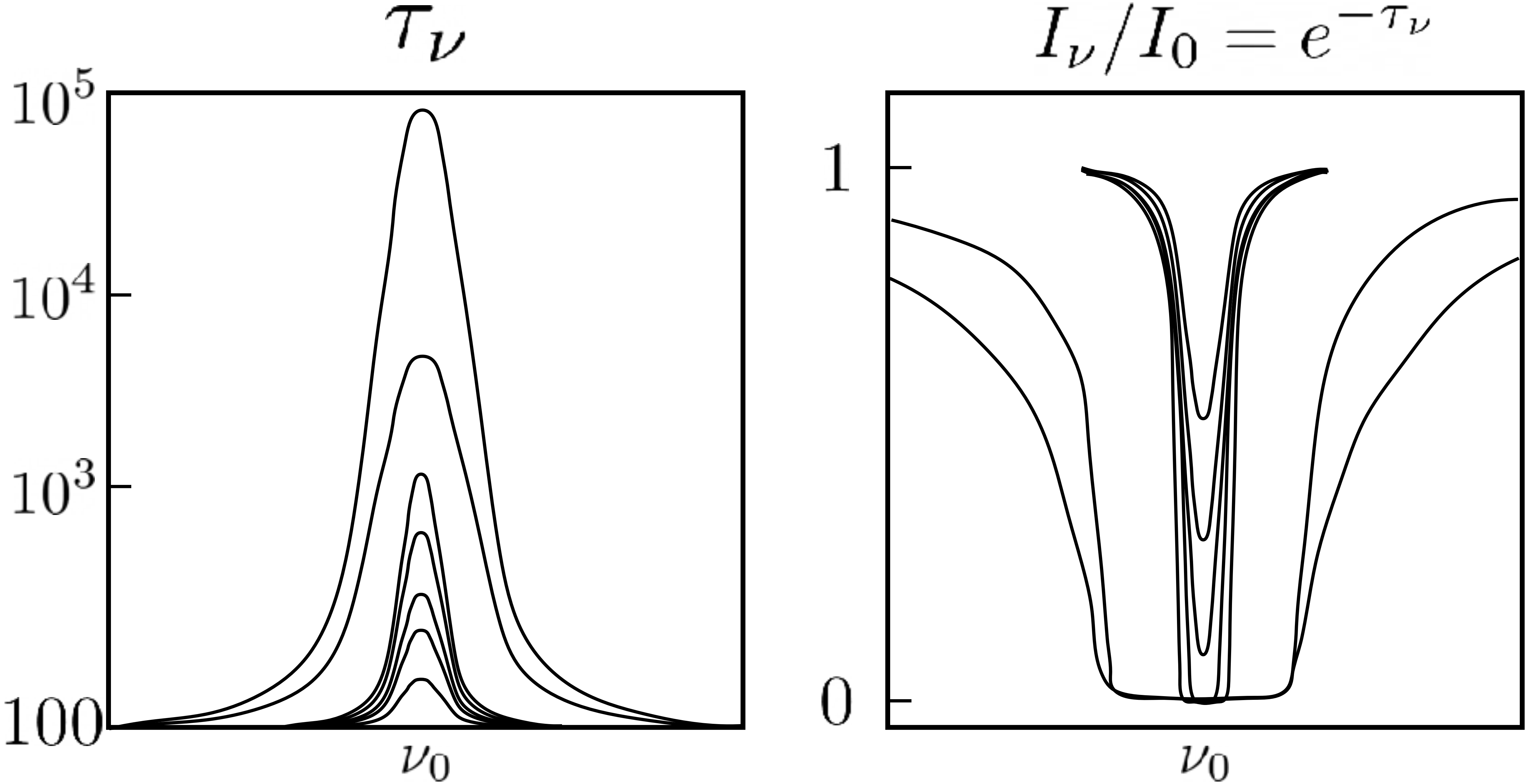

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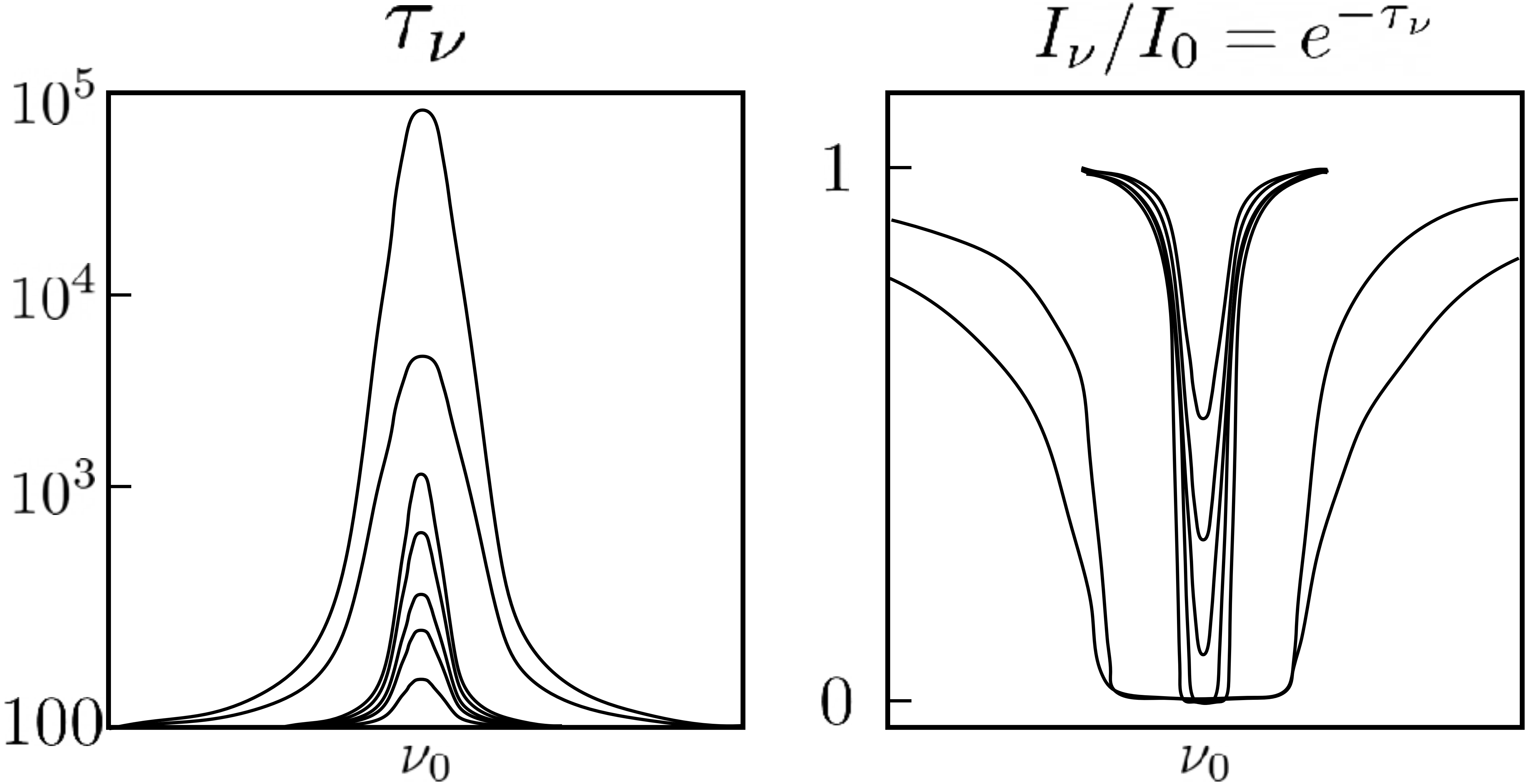
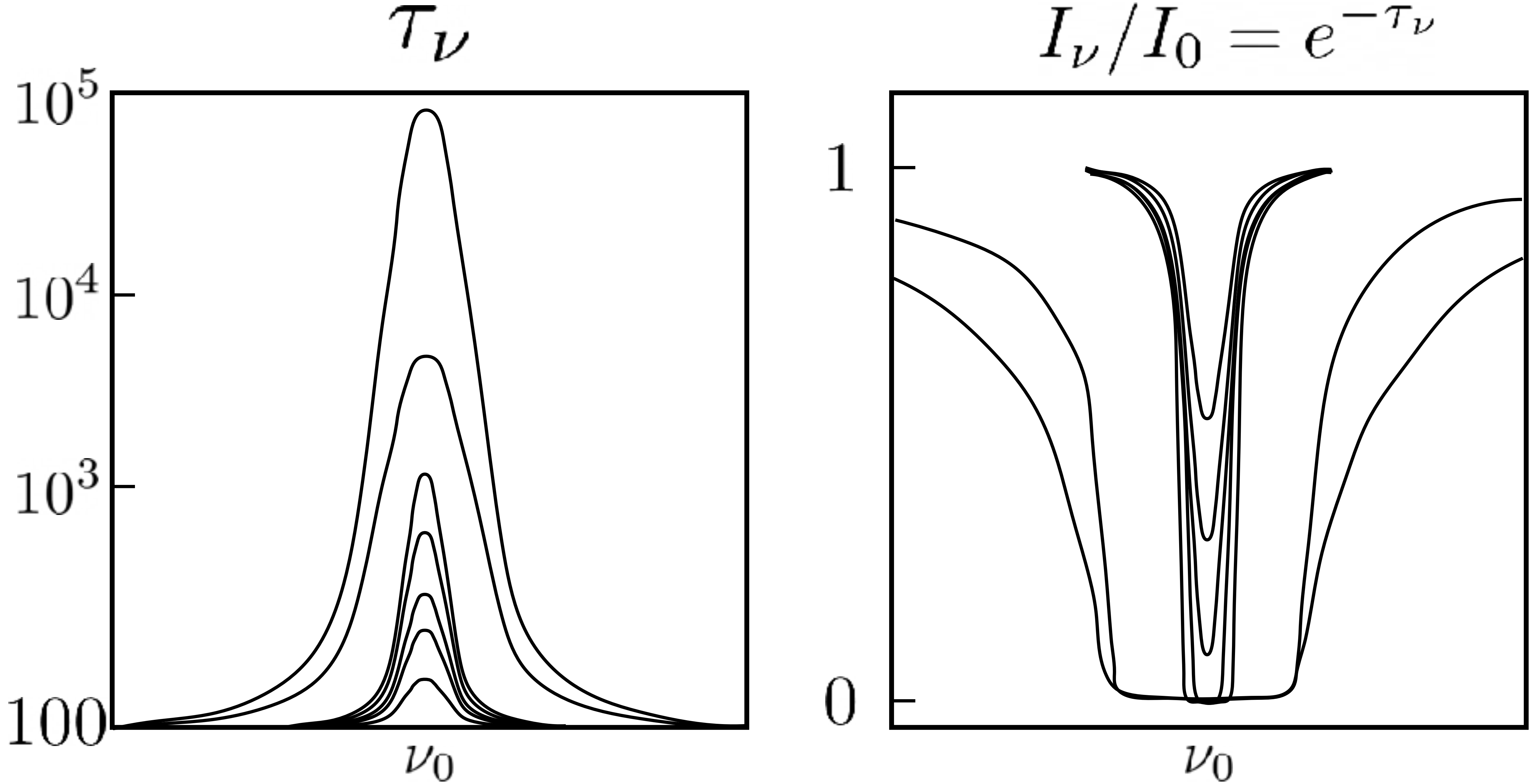


eventually “damping wings” also get close to optical depth ~1

The Curve of Growth



$W \sim N^{0.5}$

eventually "damping wings" also get close to optical depth ~1

Absorption Lines

"Linear"	$W \propto N$	$\tau_o \ll 1$
"Flat" or "Logarithmic"	$W \propto b\sqrt{\ln(N/b)}$	$10 \leq \tau_o \leq 10^3$
"Damped"	$W \propto \sqrt{N}$	$\tau_o \geq 10^4$

$$\tau_o = \frac{\pi^{1/2} e^2}{mc} \frac{\lambda}{b} N f$$

optical depth at line center

Absorption Lines

"Curve of Growth"

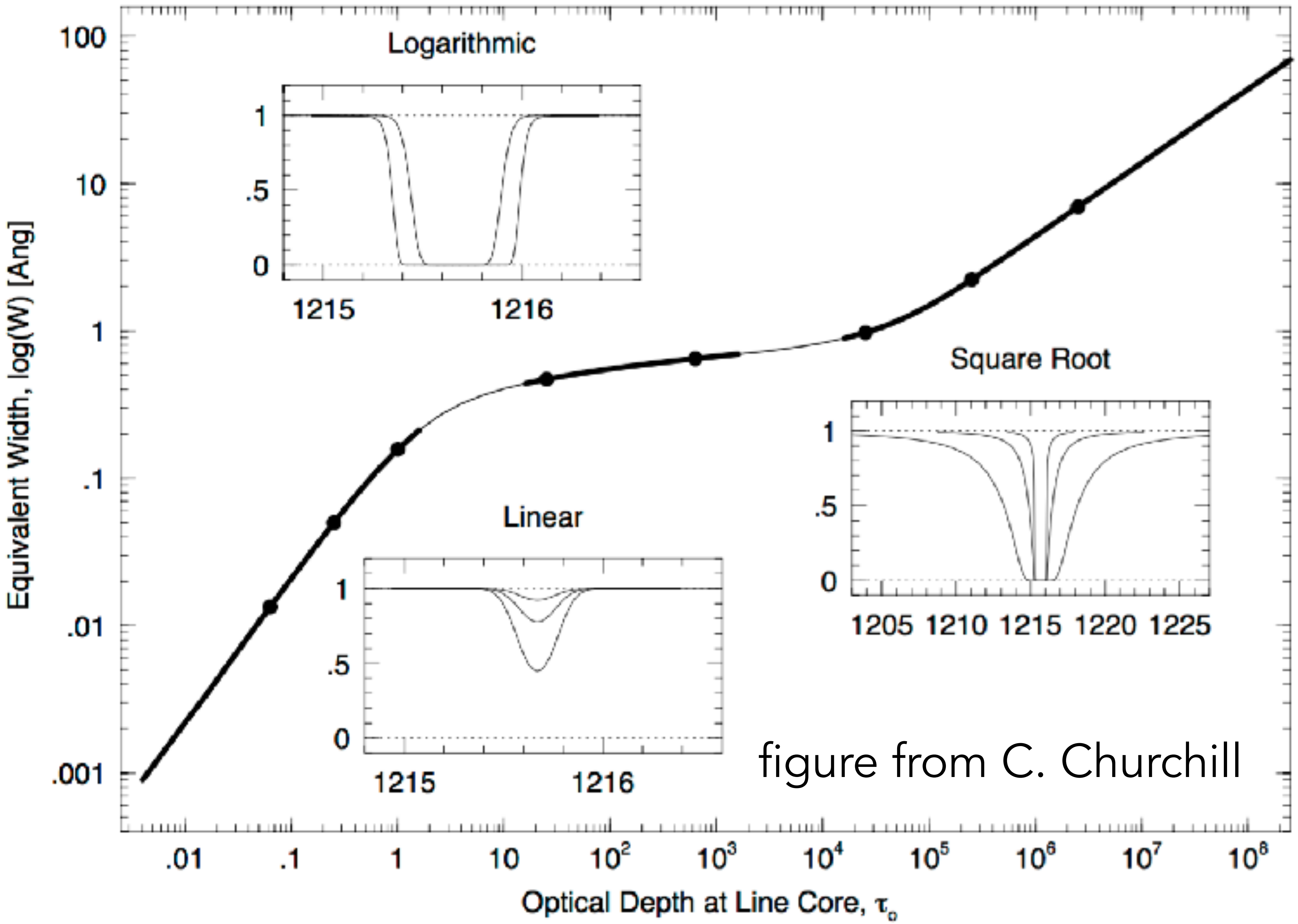


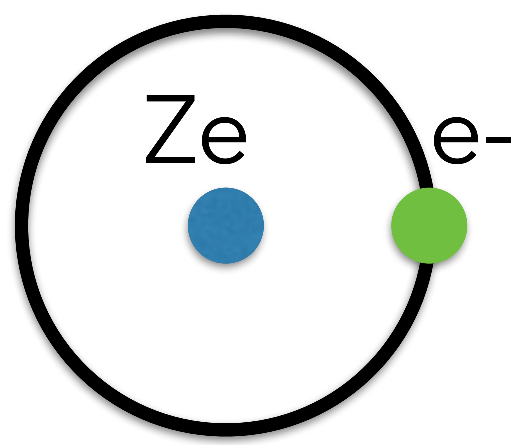
figure from C. Churchill

Summary of Opt/UV Absorption Lines

- E_{ul} is big, don't worry about stimulated emission, most in ground state.
- Line profile is Voigt (convolution of Gaussian with natural broadening)
- In low optical depth limit, only the Gaussian part matters, EW (equivalent width) is proportional to N (column density).
- Once line center saturates, EW has a "flat" dependence on N (i.e. sqrt(log(N))). Bad regime for measuring N!
- At very high optical depth, Lorentzian wings are important and EW depends on sqrt(N), can measure N from EW again.

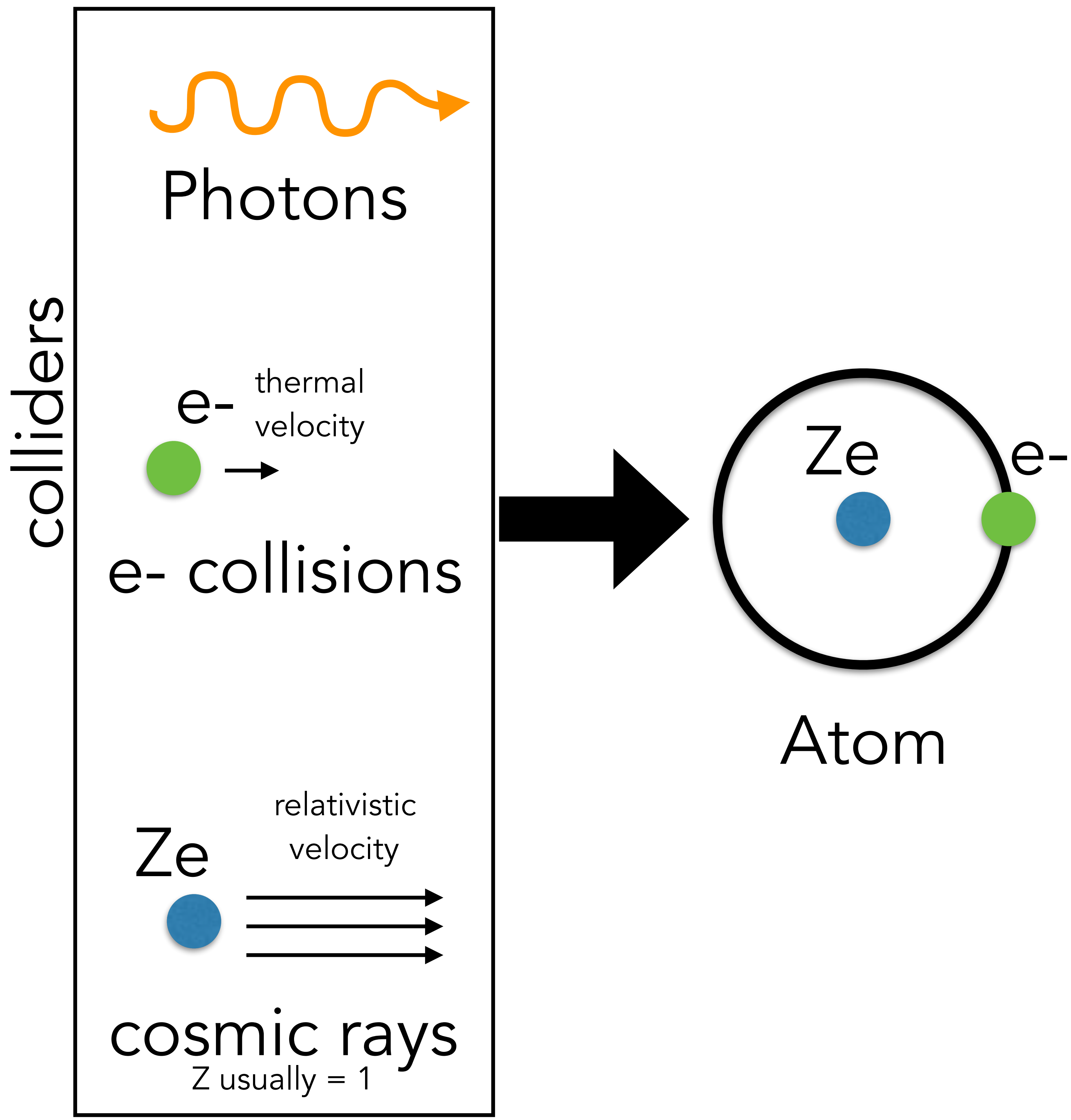
Ionization Processes

Ionization Processes



Atom

Ionization Processes



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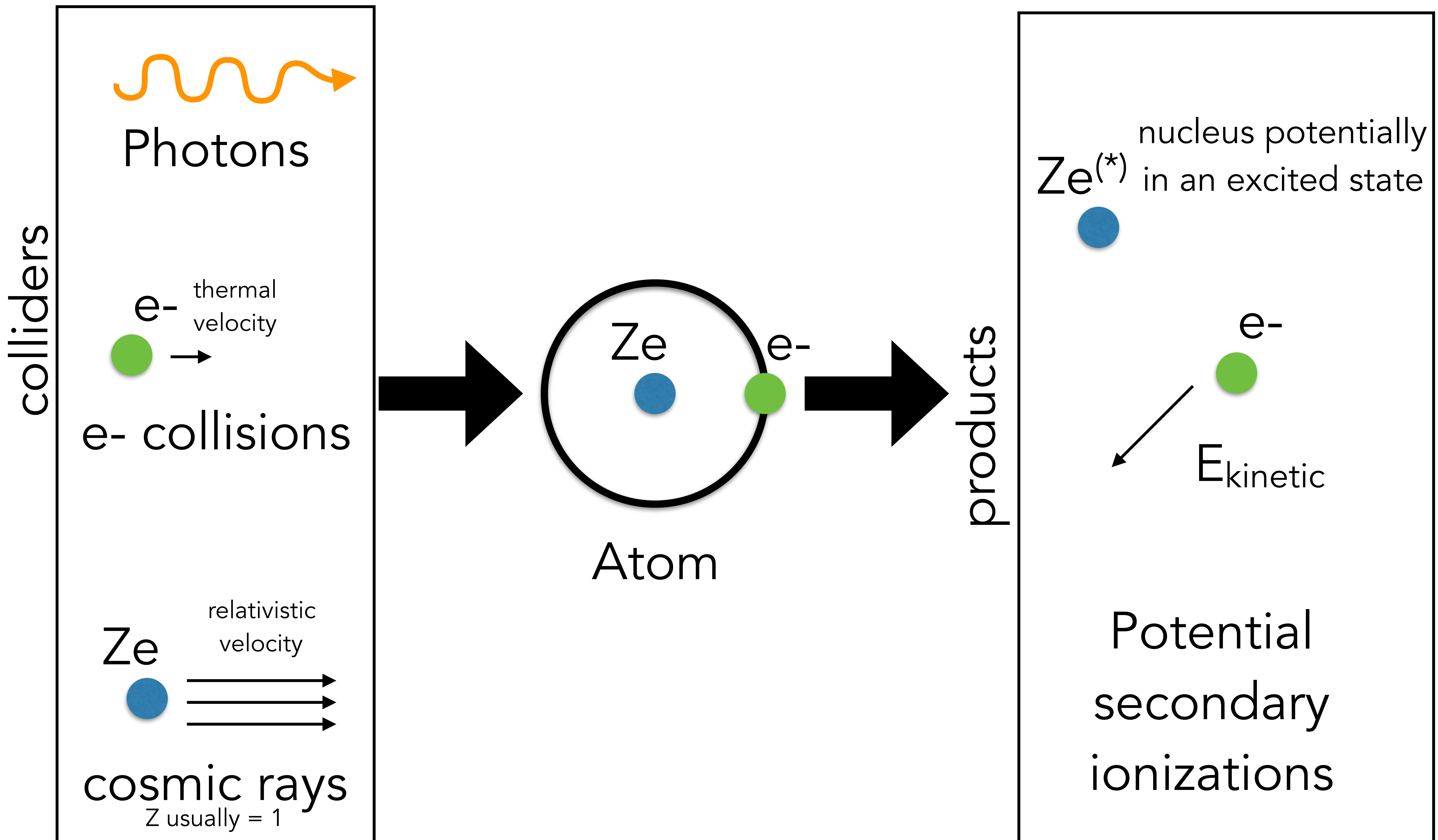

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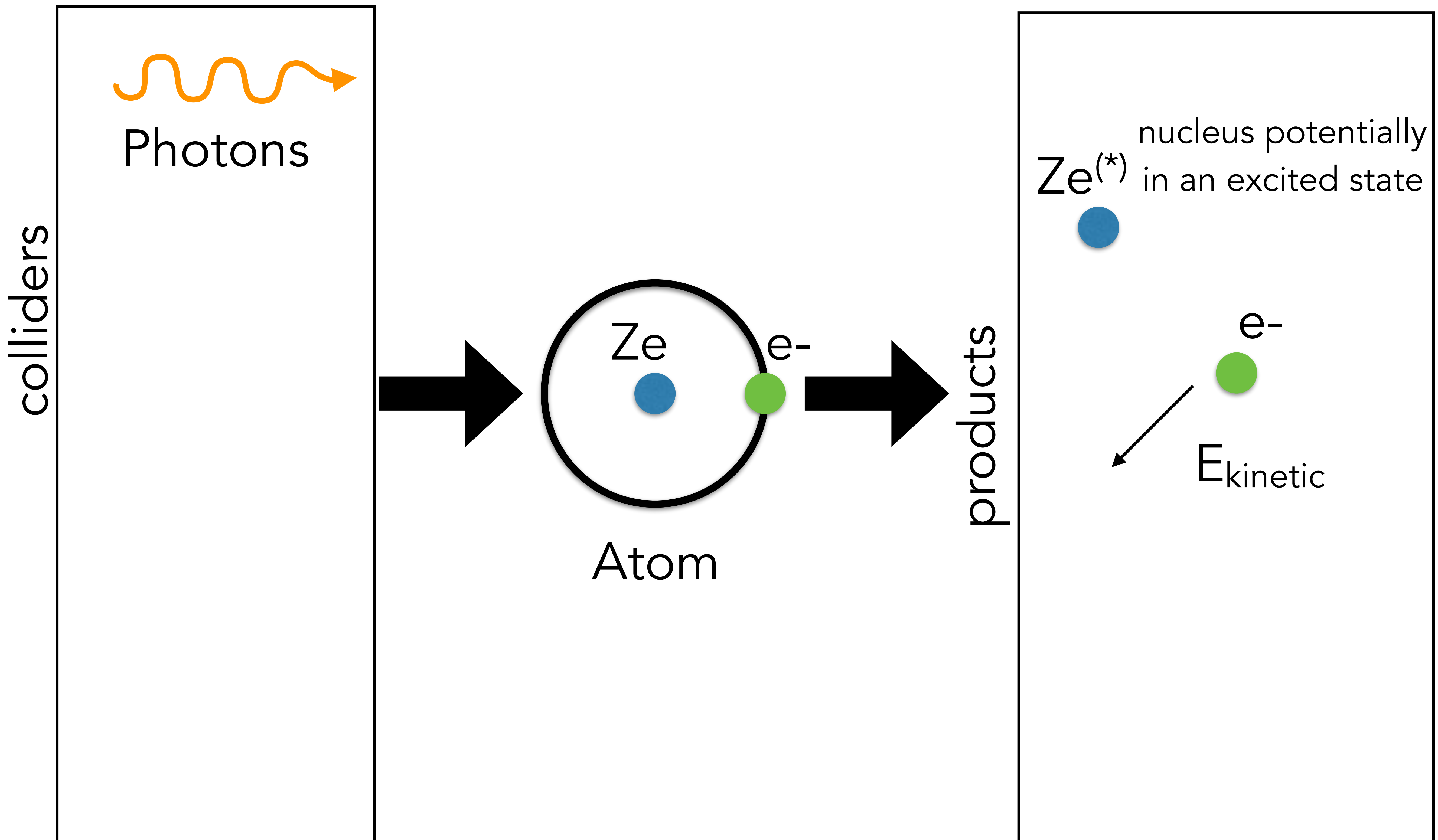


Photoionization

Cross section can be determined analytically for Hydrogen
(and "hydrogenic" ions - those with 1 e- remaining)

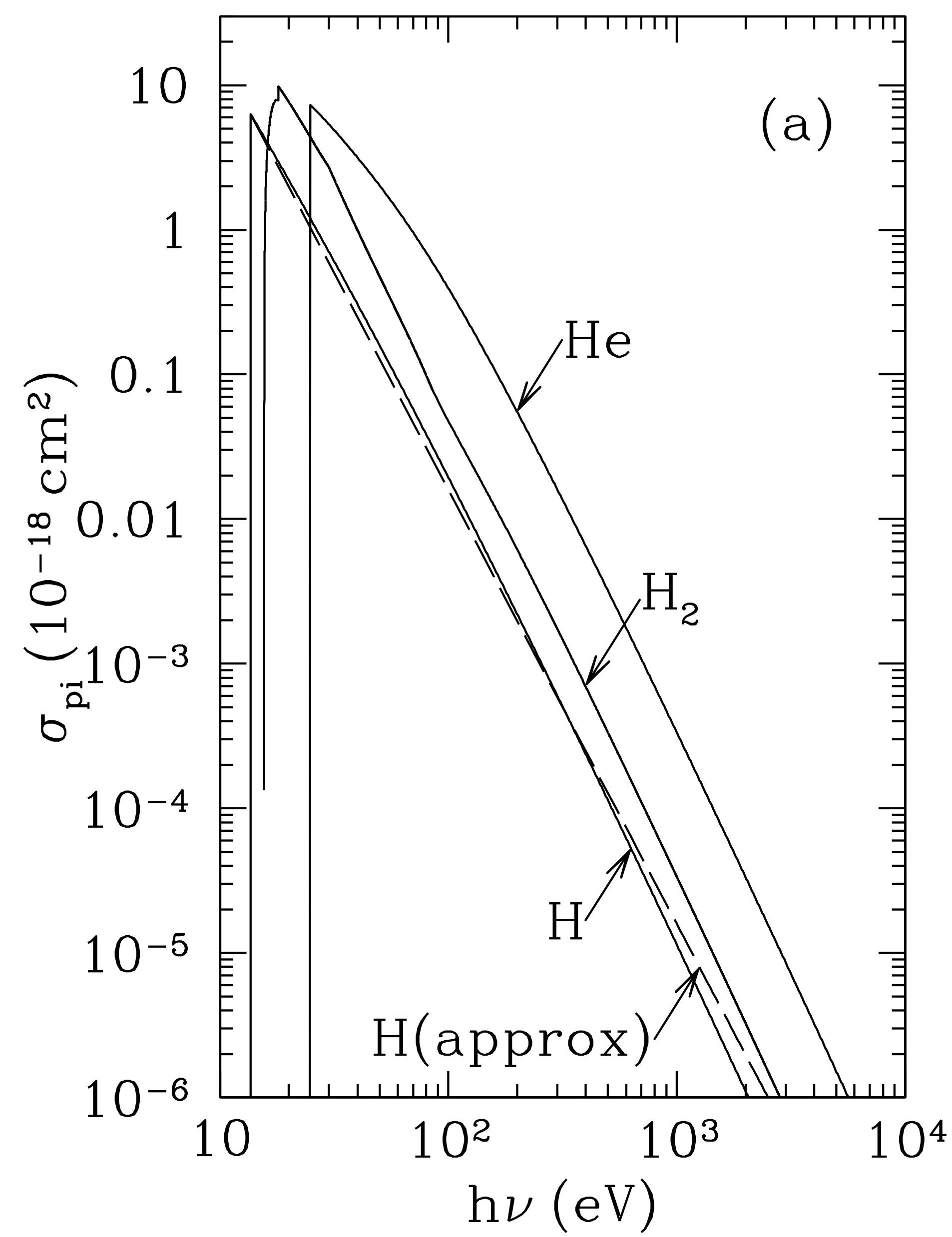
when hν > 13.6 Z² eV

$$\sigma_{\rm pi}(\nu) = \sigma_0 \left(\frac{Z^2 I_{\rm H}}{h\nu}\right)^4 \frac{e^{4-(4\tan^{-1}x)/x}}{1-e^{-2\pi/x}}$$

where: $x = \sqrt{\frac{h\nu}{Z^2 I_{\rm H}} - 1}$

and "cross section at threshold" is

$$\sigma_0 = \frac{2^9\pi}{3e^4} Z^{-2}\alpha\pi a_0^2 = 6.304 \times 10^{-18} Z^{-2}\ \mathrm{cm}^{-2}$$



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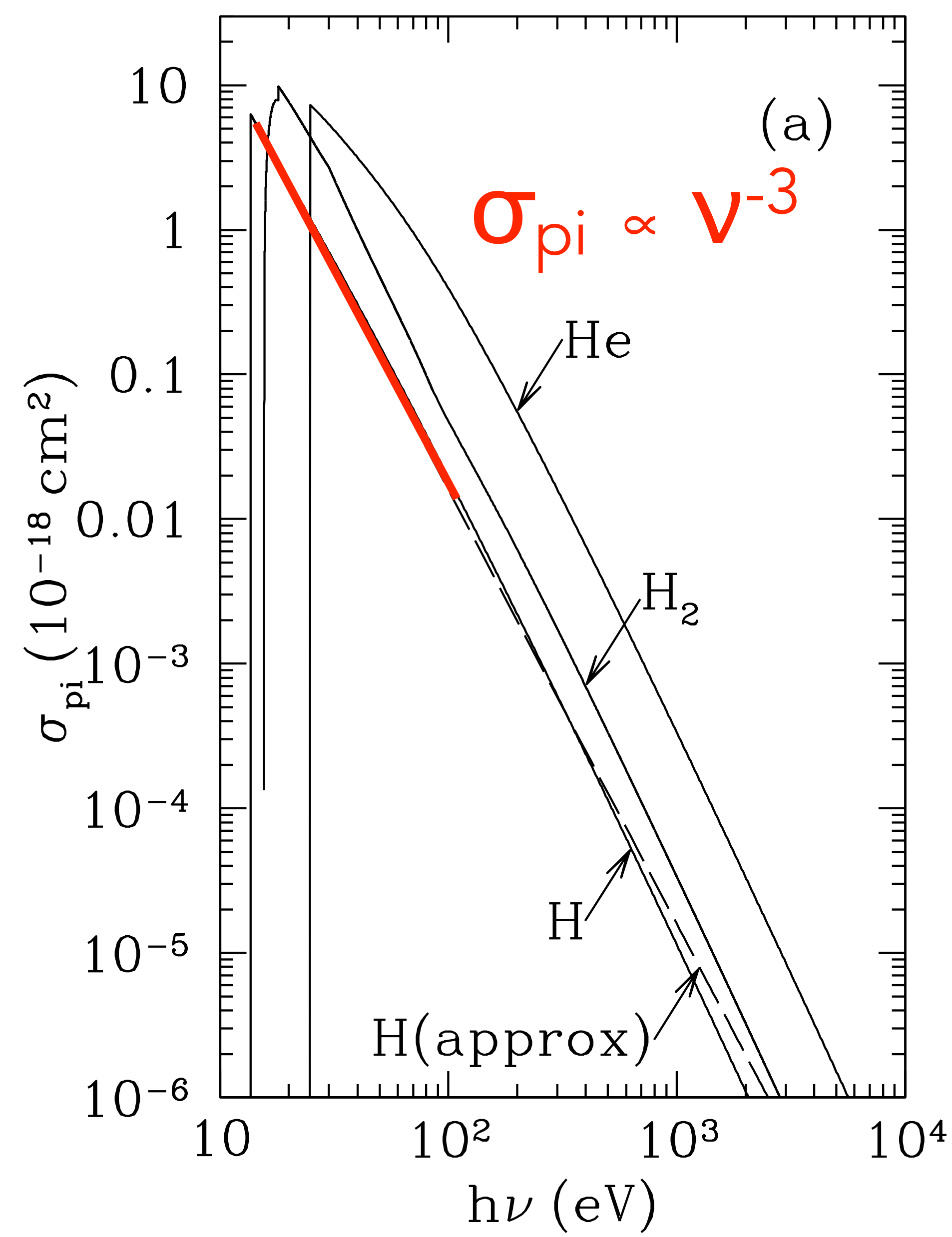

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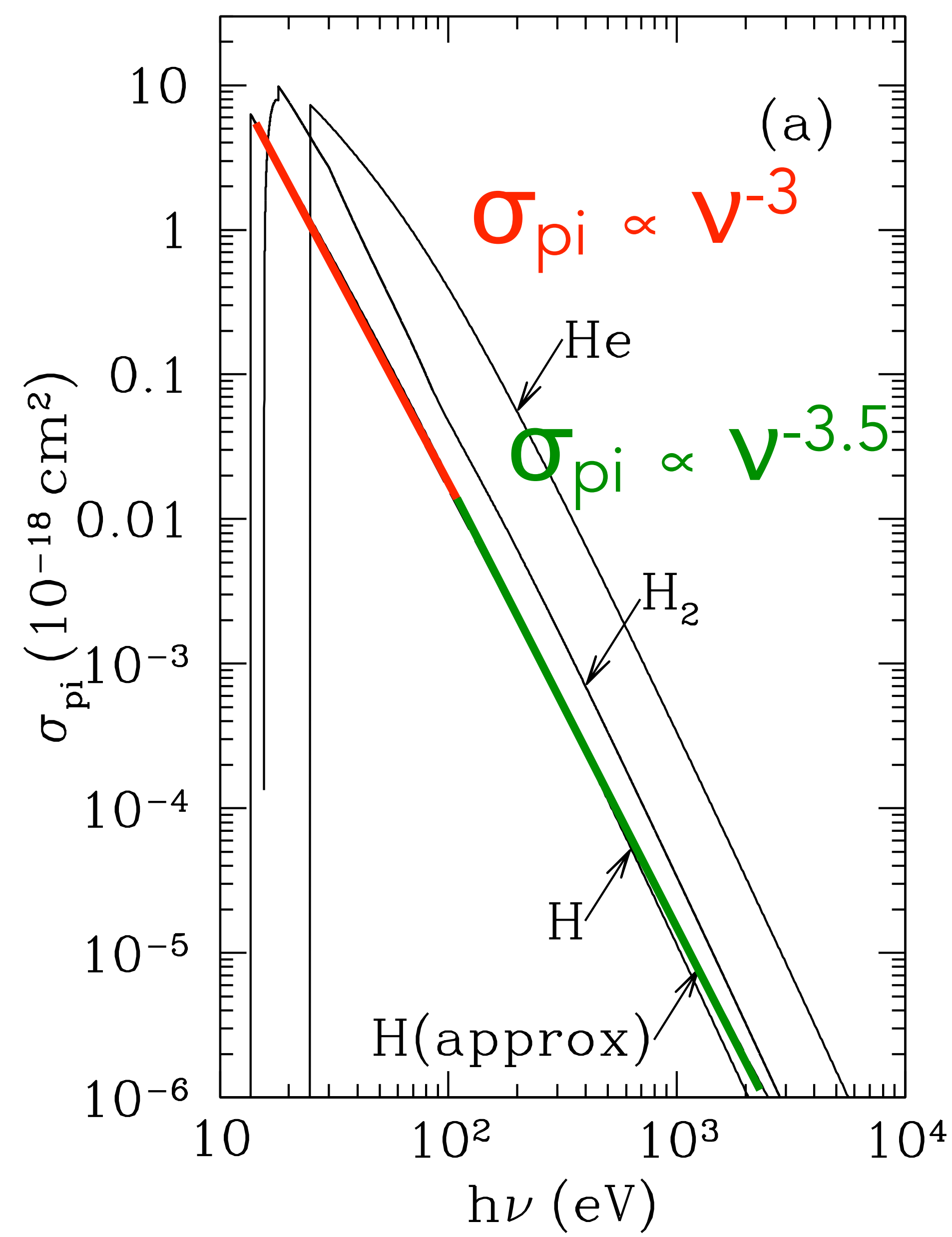
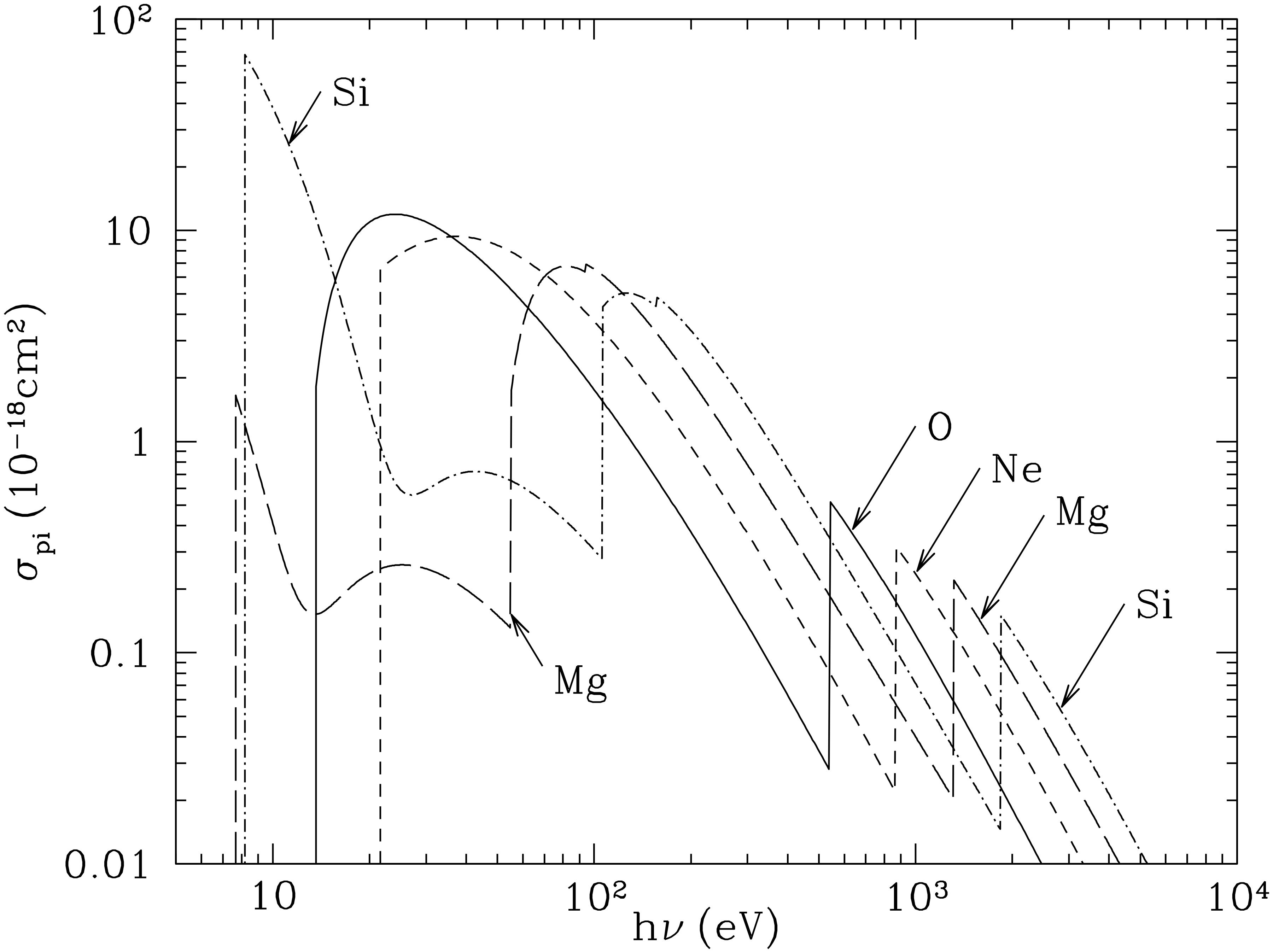


Photoionization

Cross section complexity increases with multiple electrons.



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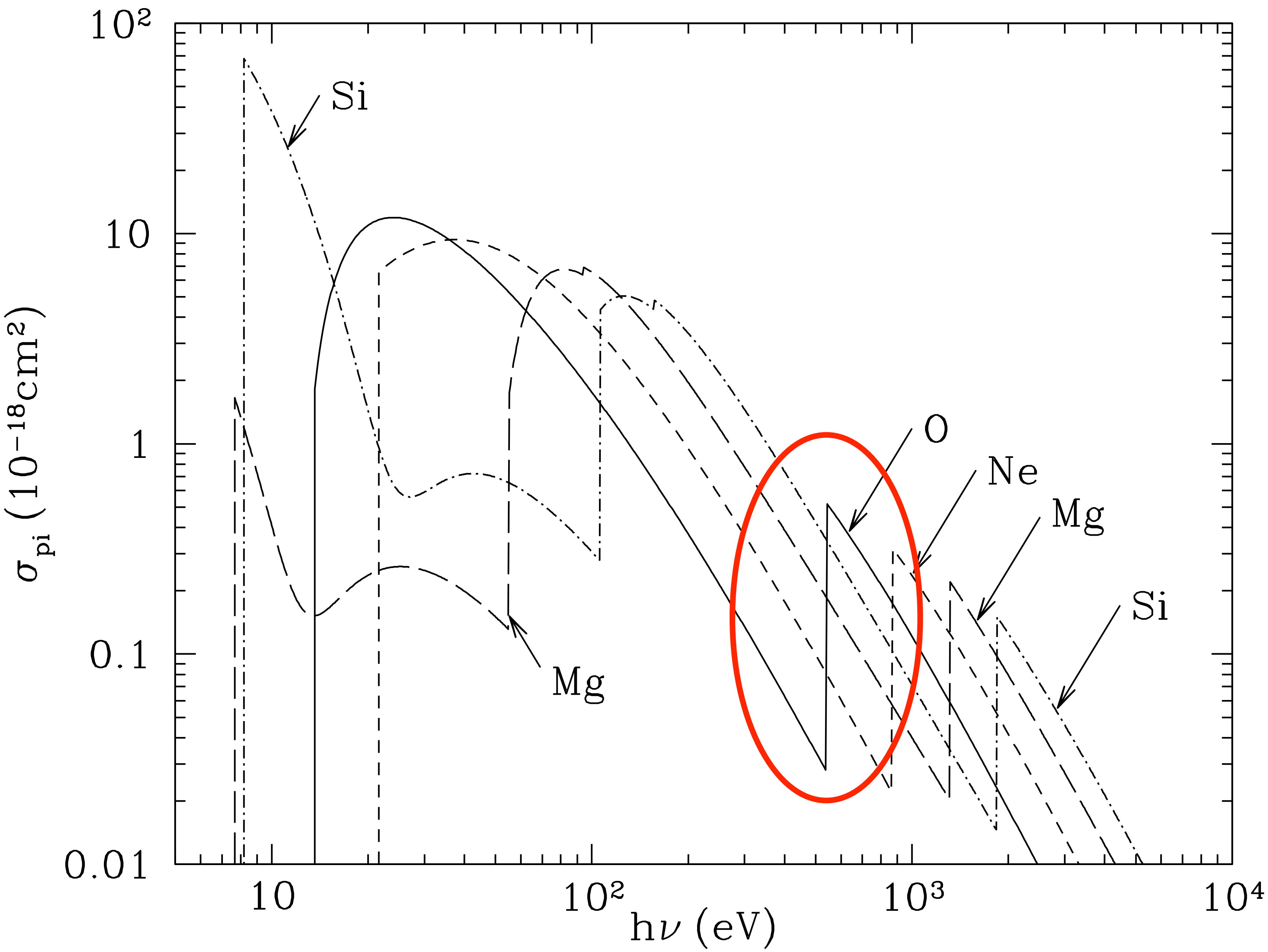
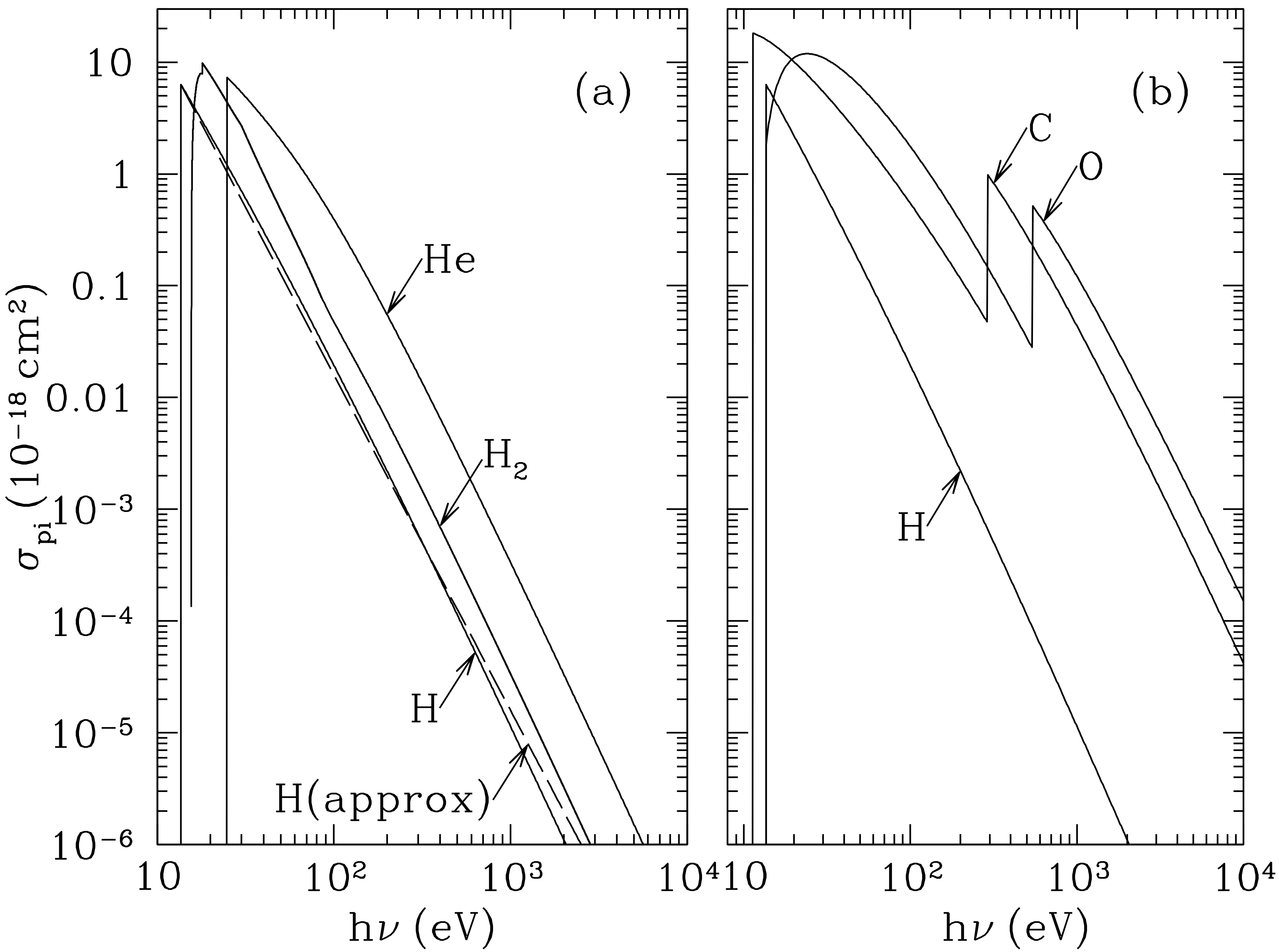


"absorption edge" due to K shell (the 1s shell)

at binding energy of K shell cross section increases sharply

Photoionization

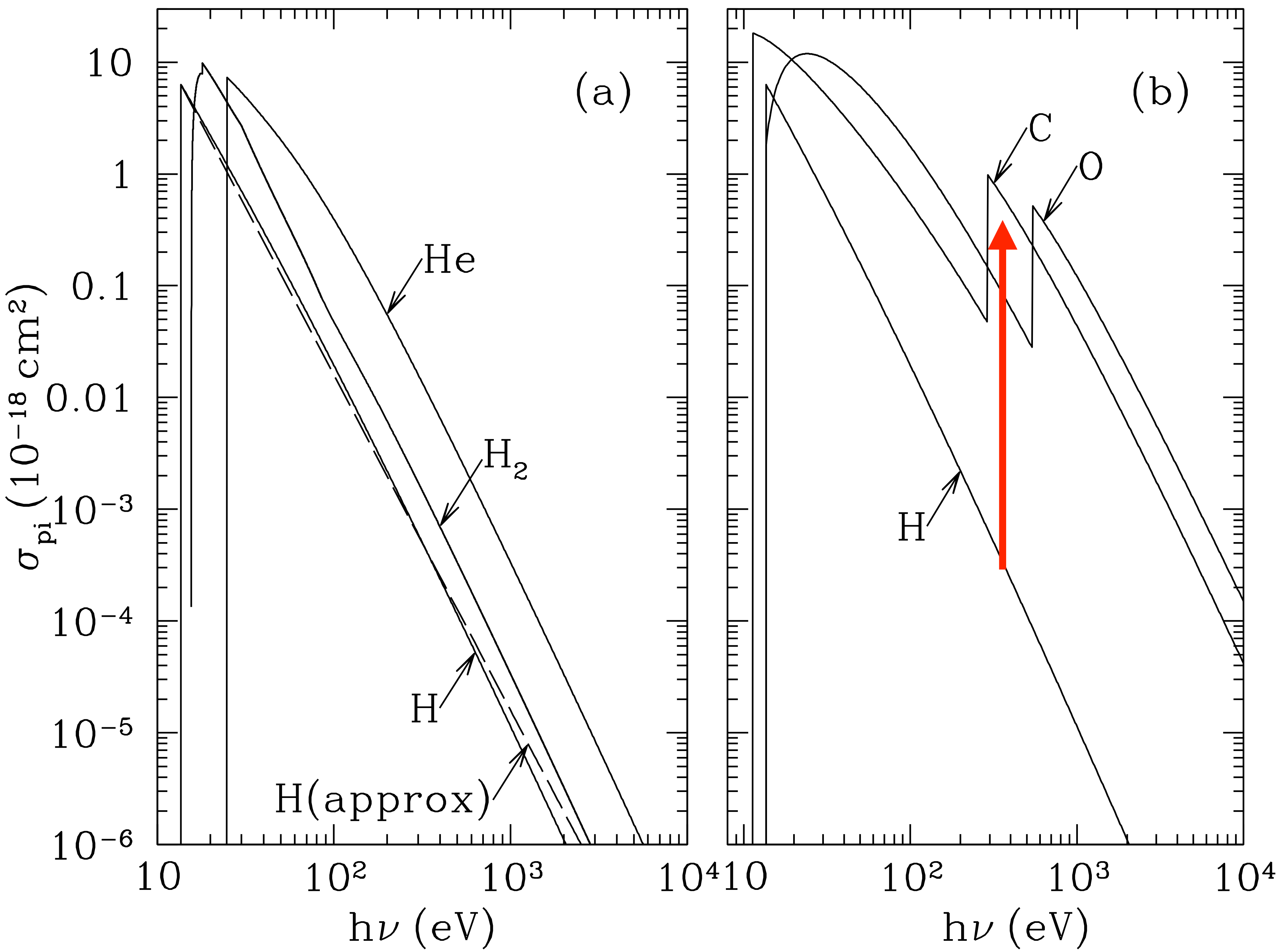


Note:

cross section of C and O and other metals far exceeds H at high energy

Even though they are less abundant, metals dominate PI rate of gas at high energies.

Photoionization



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Photoionization

rate per volume ~ $n_{atom}\ n_{collider}\ \sigma\ c$

Photoionization

rate per volume ~ n_{atom} $n_{collider}$ σ c

ζpi =
photoionization rate

Photoionization

rate per volume ~ n_{atom} $n_{collider}$ σ c

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$$\zeta_{pi} = \int_{\nu_1}^{\infty} \sigma_{pi}(\nu)\ c\ \frac{u_\nu}{h\nu}\ d\nu$$

Photoionization

rate per volume ~ n_{atom} $n_{collider}$ σ c

ζ_{pi} =
photoionization rate

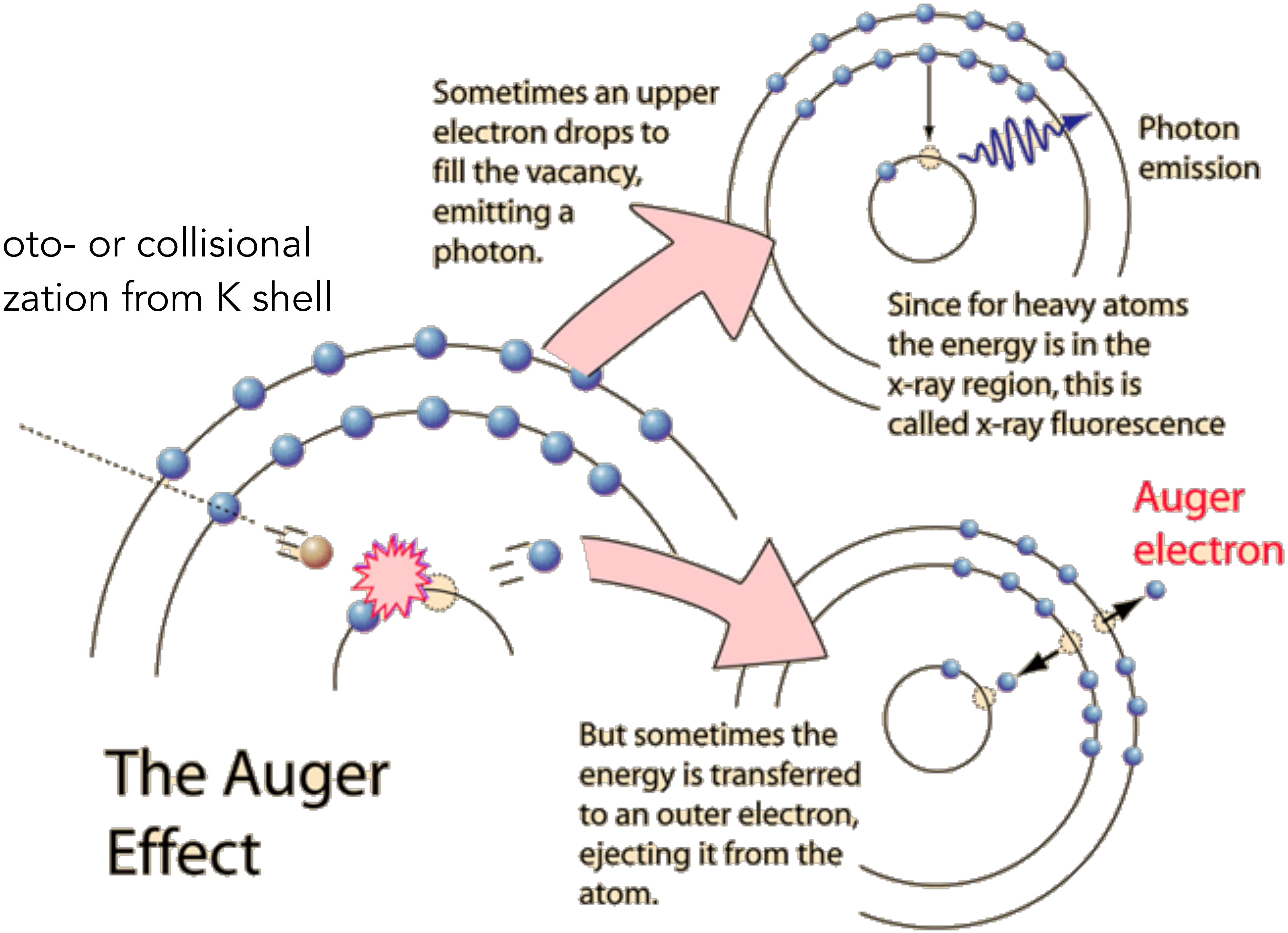
$$\zeta_{pi} = \int_{\nu_1}^{\infty} \sigma_{pi}(\nu)\ c\ \frac{u_\nu}{h\nu}\, d\nu$$

minimum energy for ionization

number density of photons

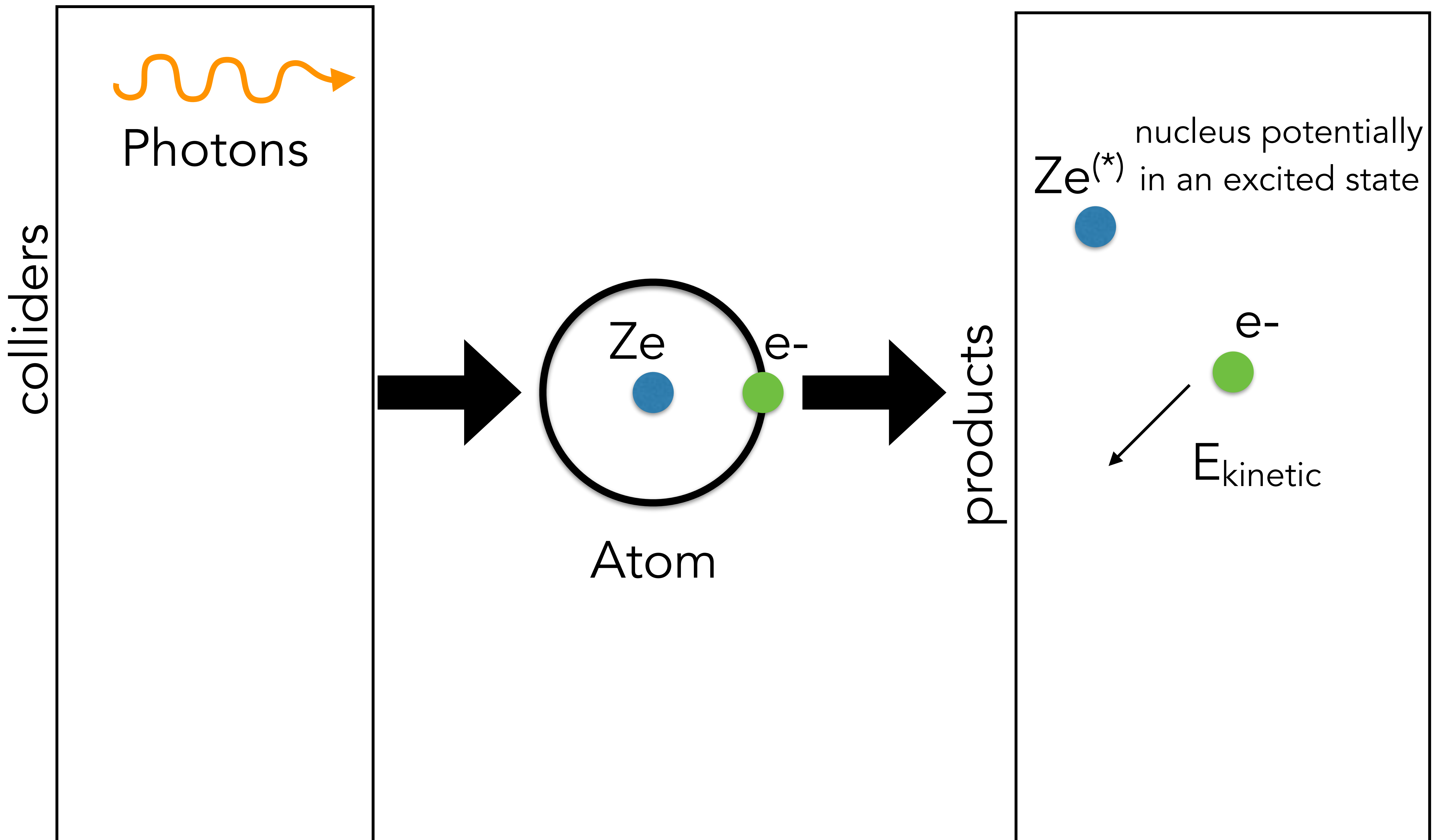
Photoionization

Photo- or collisional ionization from K shell



http://hyperphysics.phy-astr.gsu.edu/hbase/atomic/auger.html

Ionization Processes



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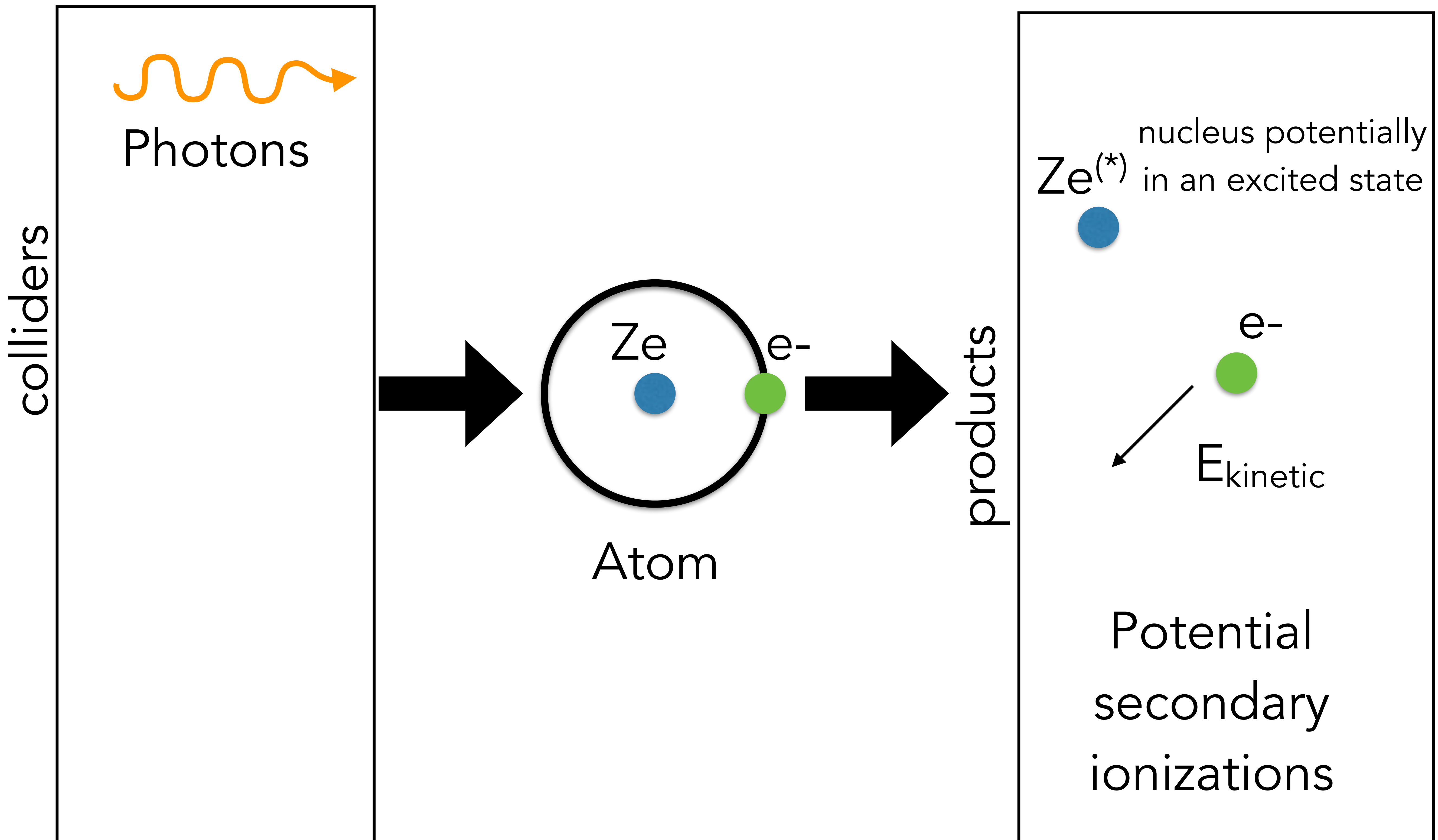


Secondary Ionizations

$$E_{pe} = h\nu - I_s$$

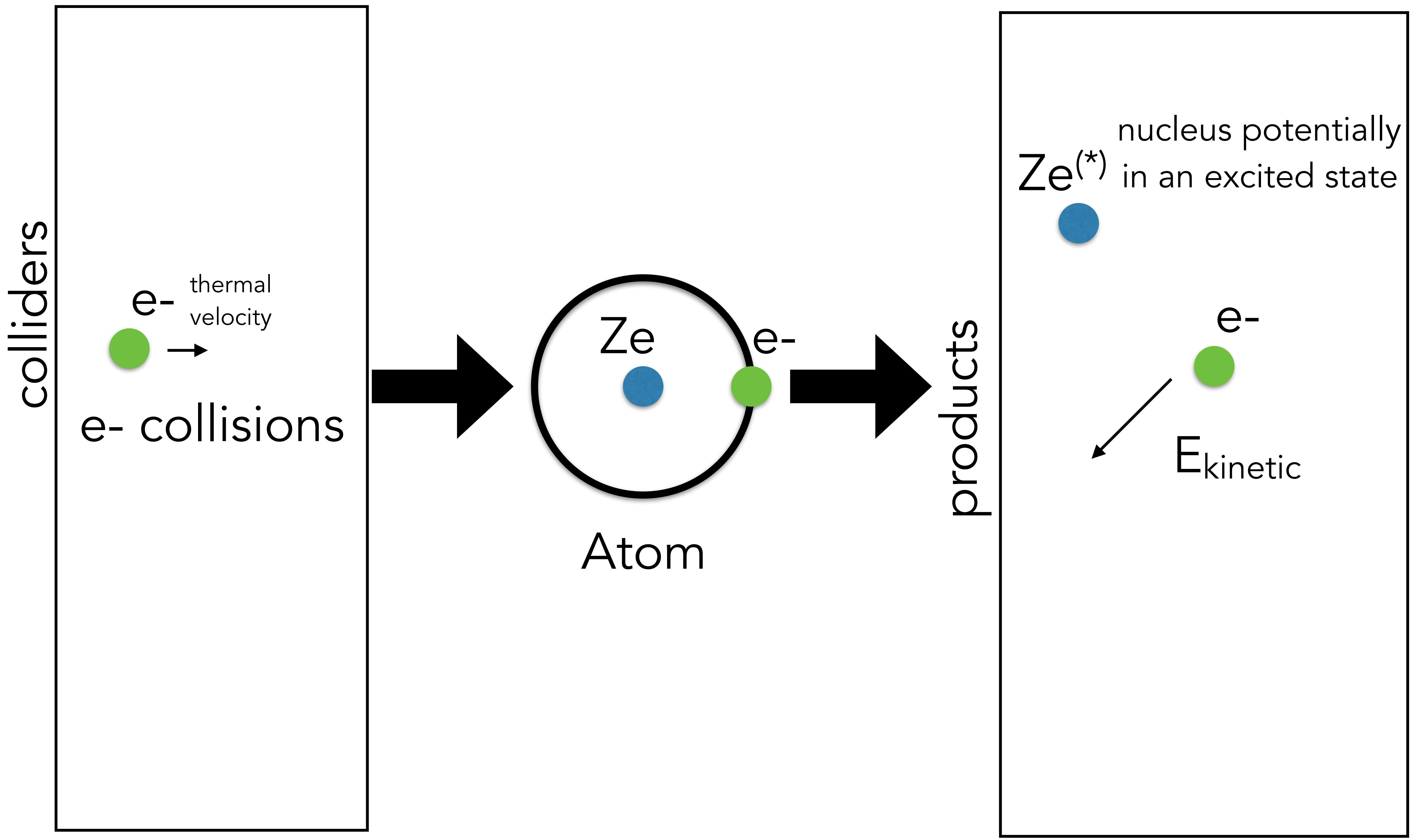
Energy of ejected photoelectron

difference between photon energy and ionization potential

For x-ray ionization E_{pe} can be big!
May go on to ionize other atoms/ions in the gas.

Secondary ionization rate depends on E_{pe} and ionization state of the gas.

Part II: Ionization Processes



Collisional Ionization

$$\zeta_{\rm ci} = k_{\rm ci} n_e n_I$$

collisional ionization rate

rate coefficient

$$k_{ci} = \int_I^\infty \sigma_{ci}(E)\ v\ f(E) dE$$

integral of cross section over Maxwellian velocity distribution

Collisional Ionization

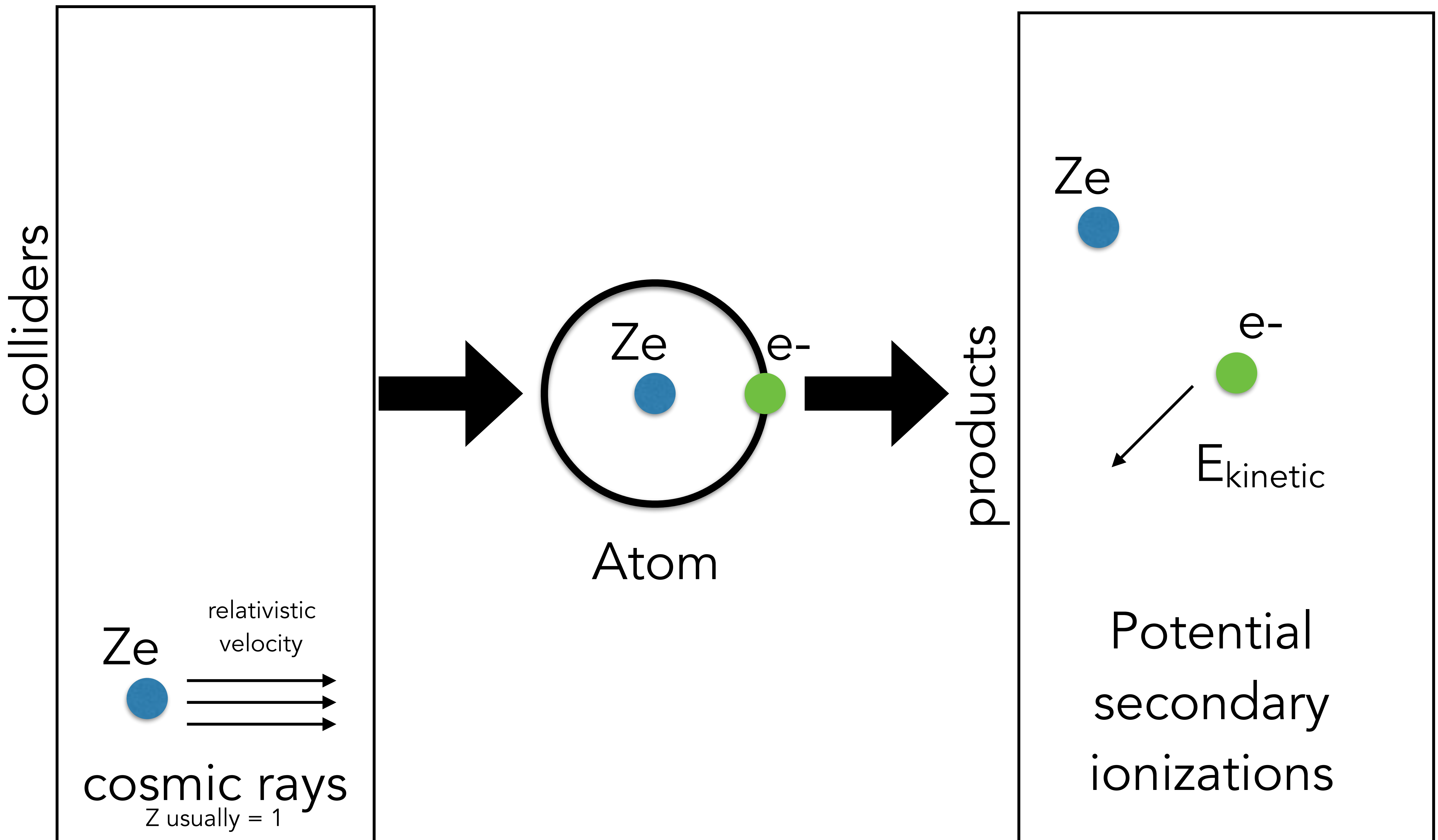
A pretty good estimate of collisional ionization cross sec when E>I:

$$\sigma_{ci}(E) = C\pi a_0^2(1 - I/E)$$

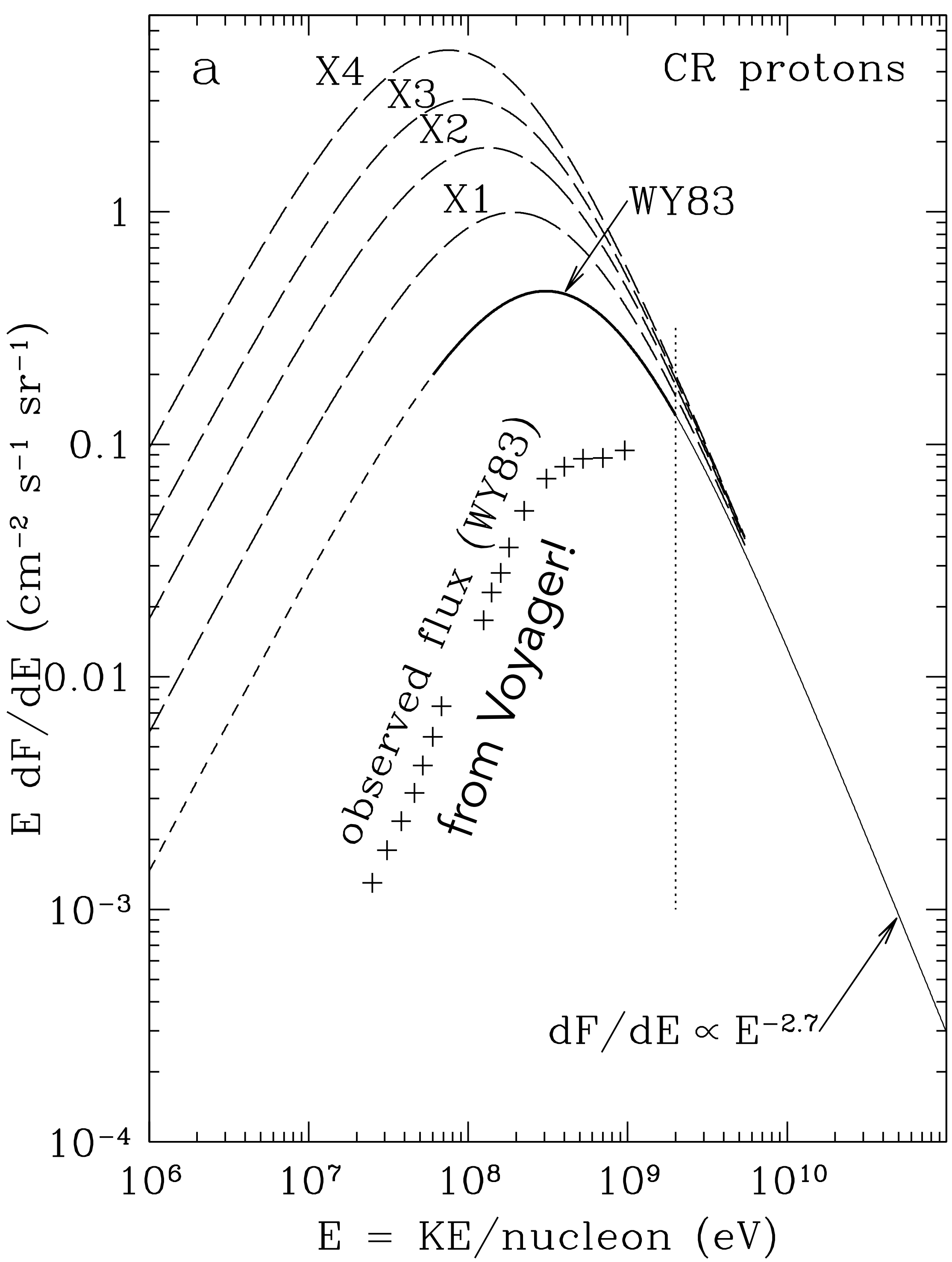
- C: constant of order unity
- πa_0^2: Bohr radius cross section
- $(1 - I/E)$: cross sec goes to zero when E=I

At higher E, cross section ~ 1/E
(can show this from the impact approx from Lecture 2)

Part II: Ionization Processes



Cosmic Ray Ionization



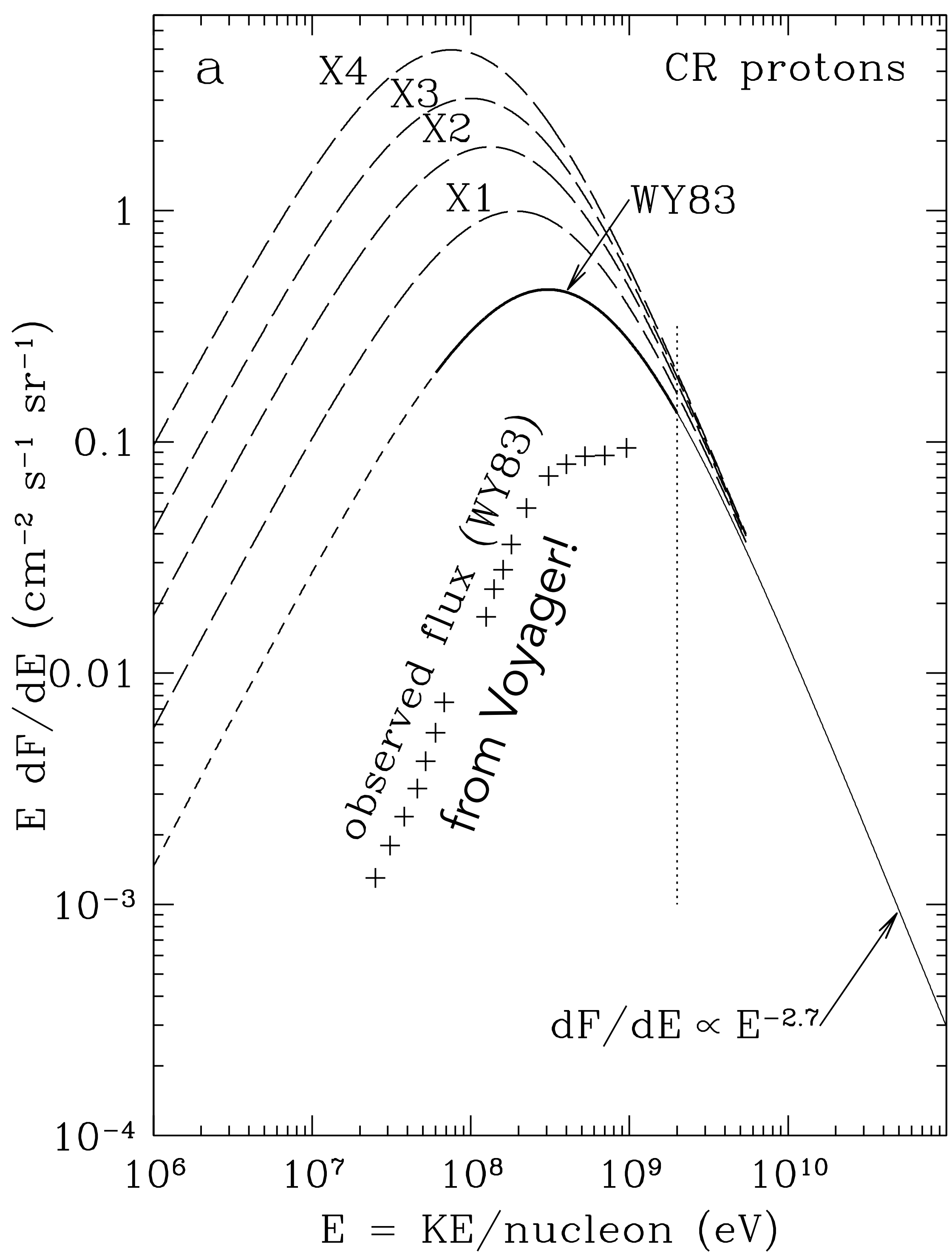
Cosmic ray energy flux is dominated by protons.

$$\zeta_{\rm CR} = 4\pi \int_{E_{\rm min}}^{\infty} \sigma_{\rm ci}(E) E \frac{dF}{dE} \cdot \frac{dE}{E}$$

Similar to before but velocity distribution is <u>not Maxwellian</u>

Big uncertainties in CR flux at low energies due to solar wind.

Cosmic Ray Ionization



CR ionization is very important in dense gas, where extinction by dust and other absorption has blocked most photons.

Will come back to this in discussing molecular clouds!

Part III: Recombination Processes

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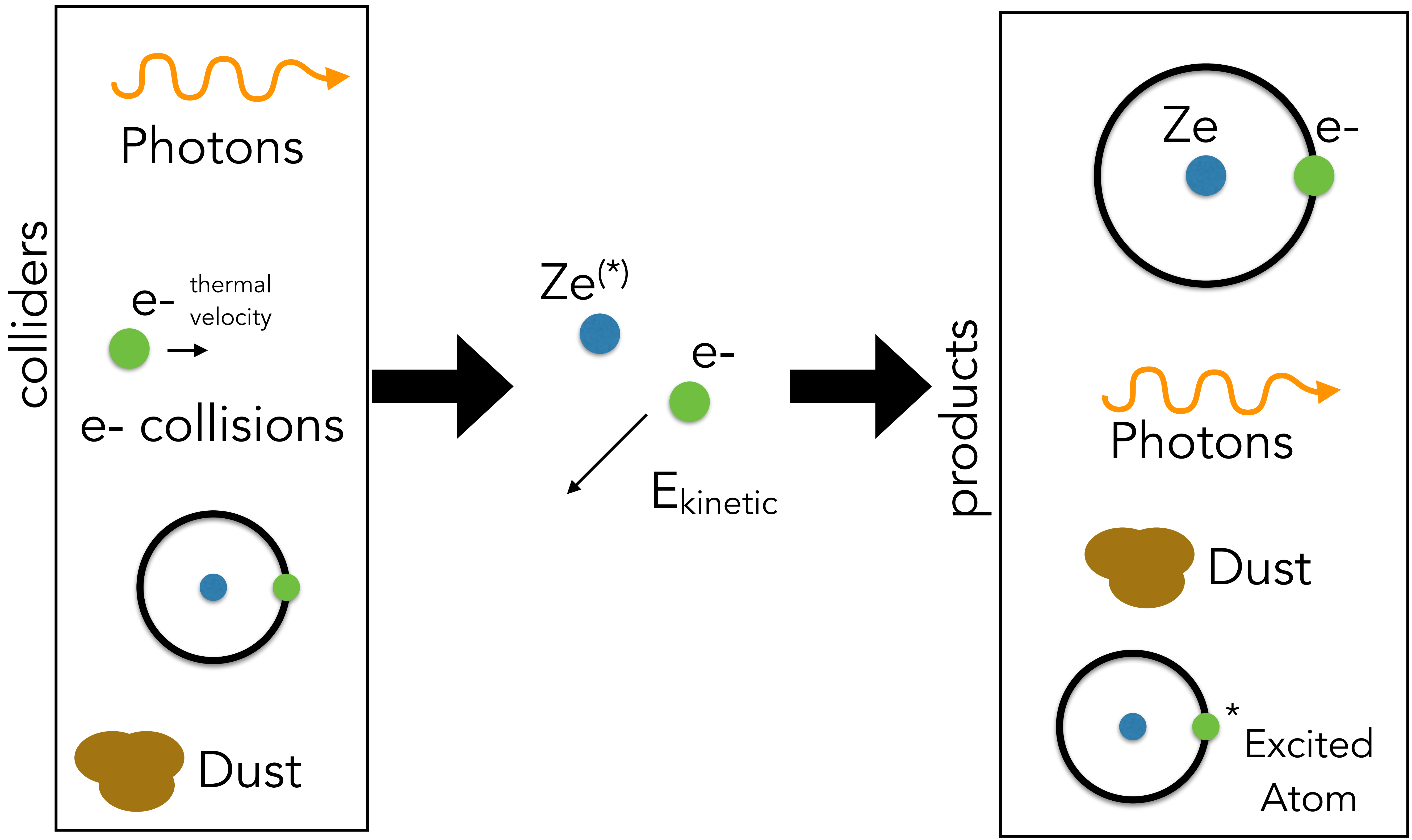
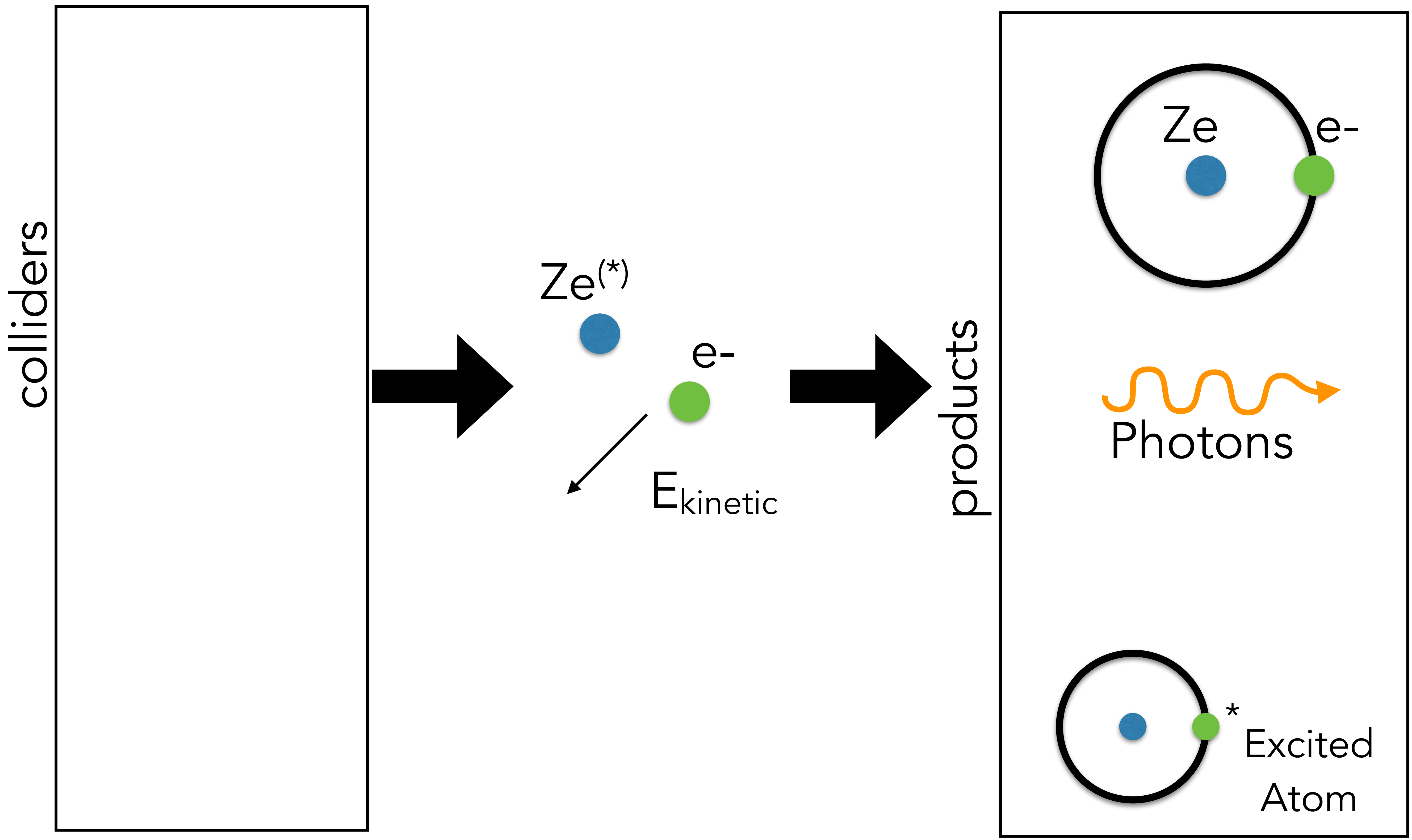
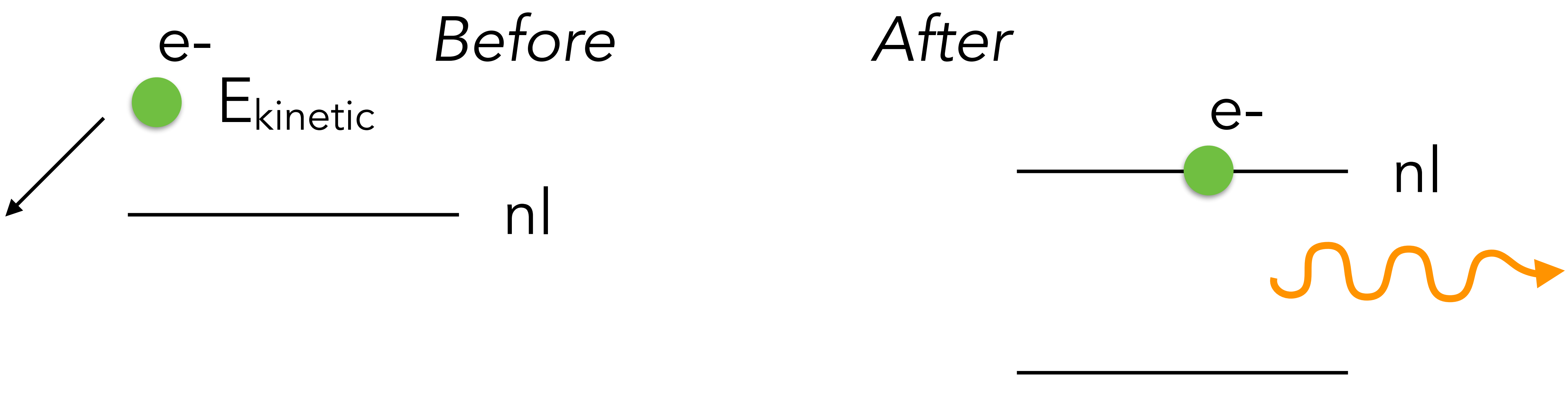


Part III: Recombination Processes



Radiative Recombination

$$X^+ + e^- \rightarrow X + h\nu$$



$E_{photon} = I_{nl} + E_{kinetic}$

I_{nl} = ionization potential from nl

Radiative Recombination

Given photoionization cross section from before, we can use detailed balance to work out radiative recombination cross section.

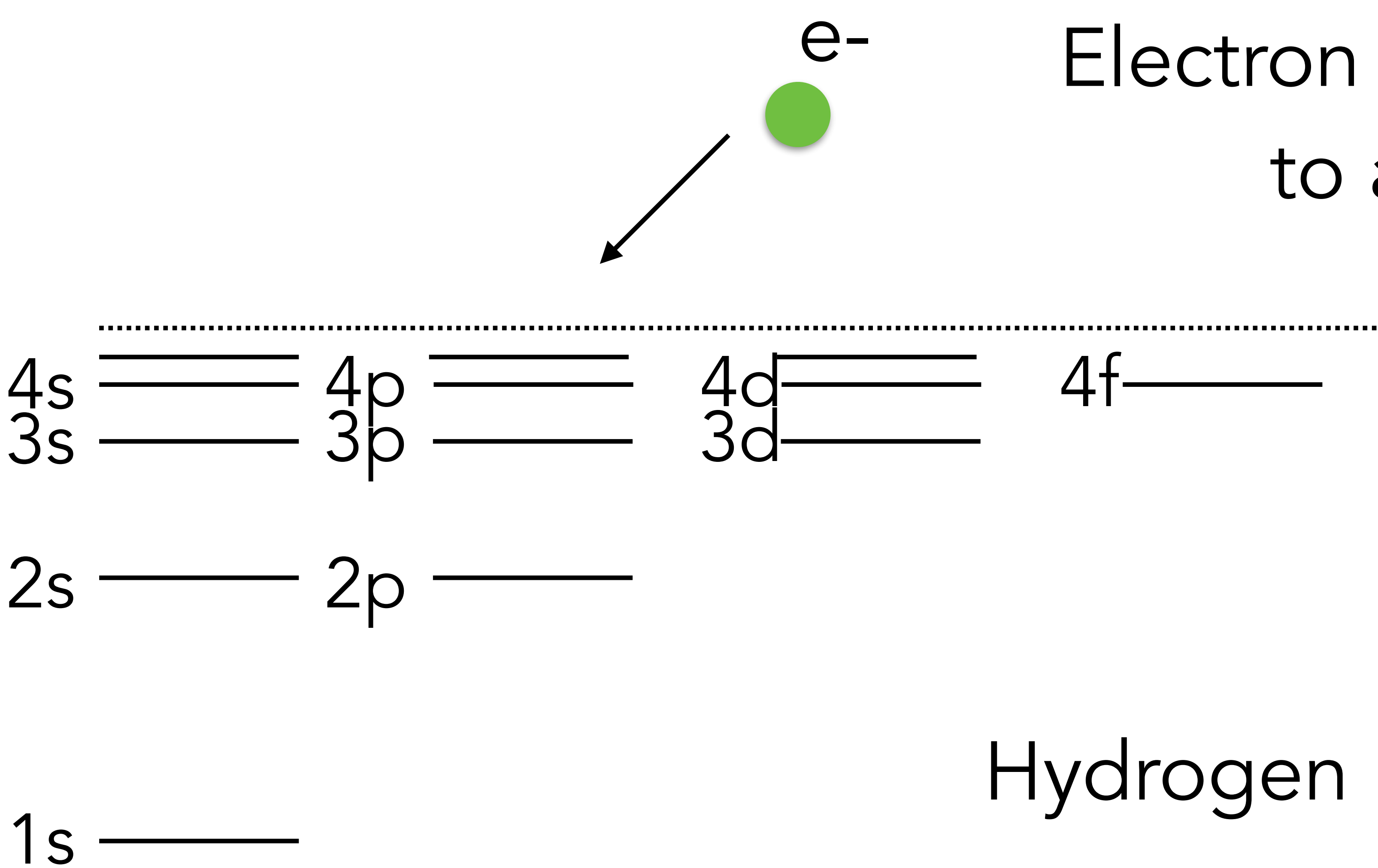
Milne Relation:

$$\sigma_{\rm rr}(E) = \frac{g_\ell}{g_u}\frac{(I_{X,u\ell}+E)^2}{E m_e c^2}\sigma_{\rm pi}(h\nu = I_{X,u\ell}+E)$$

$$X_u^+ + e^- \rightarrow X_\ell + h\nu$$

Radiative Recombination

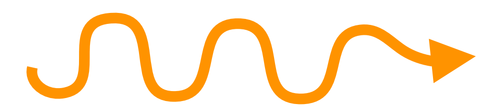
Electron can recombine to any level.



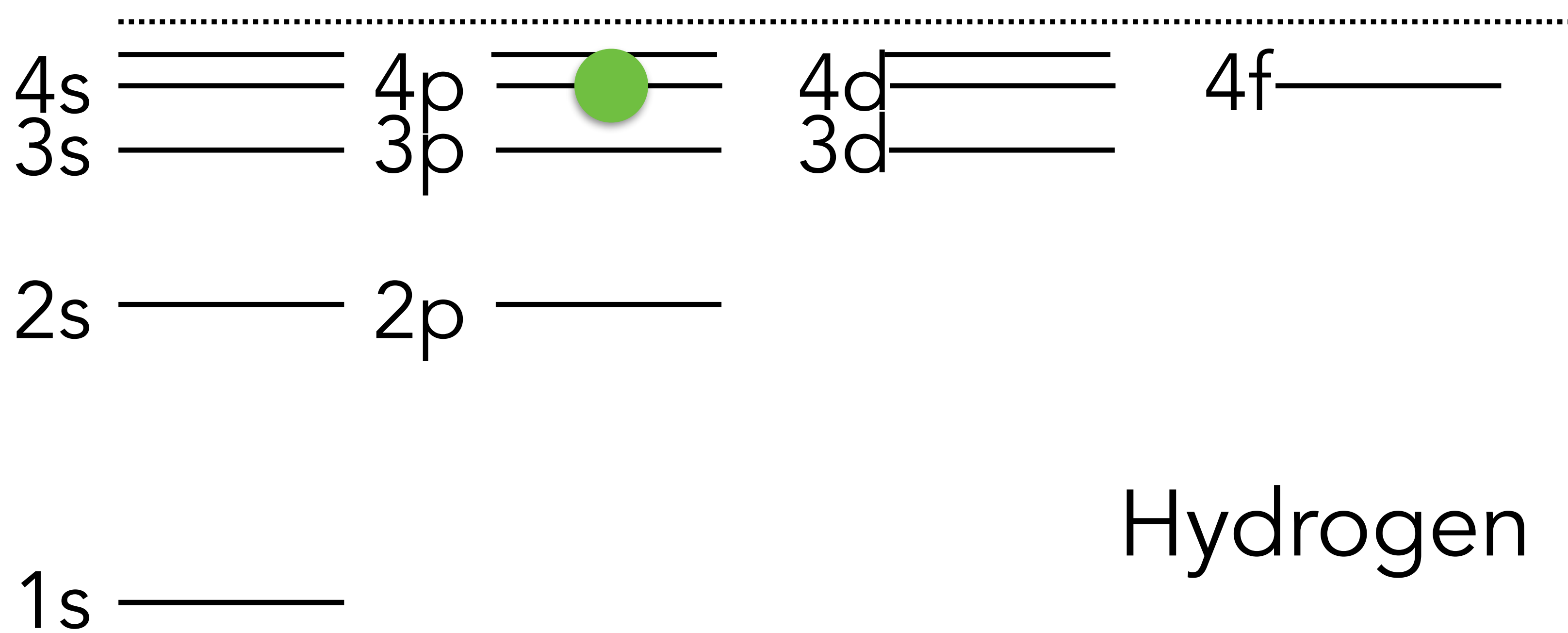
Hydrogen

Energy not to scale

Radiative Recombination



$E_{photon} = I_{nl} + E_{kinetic}$

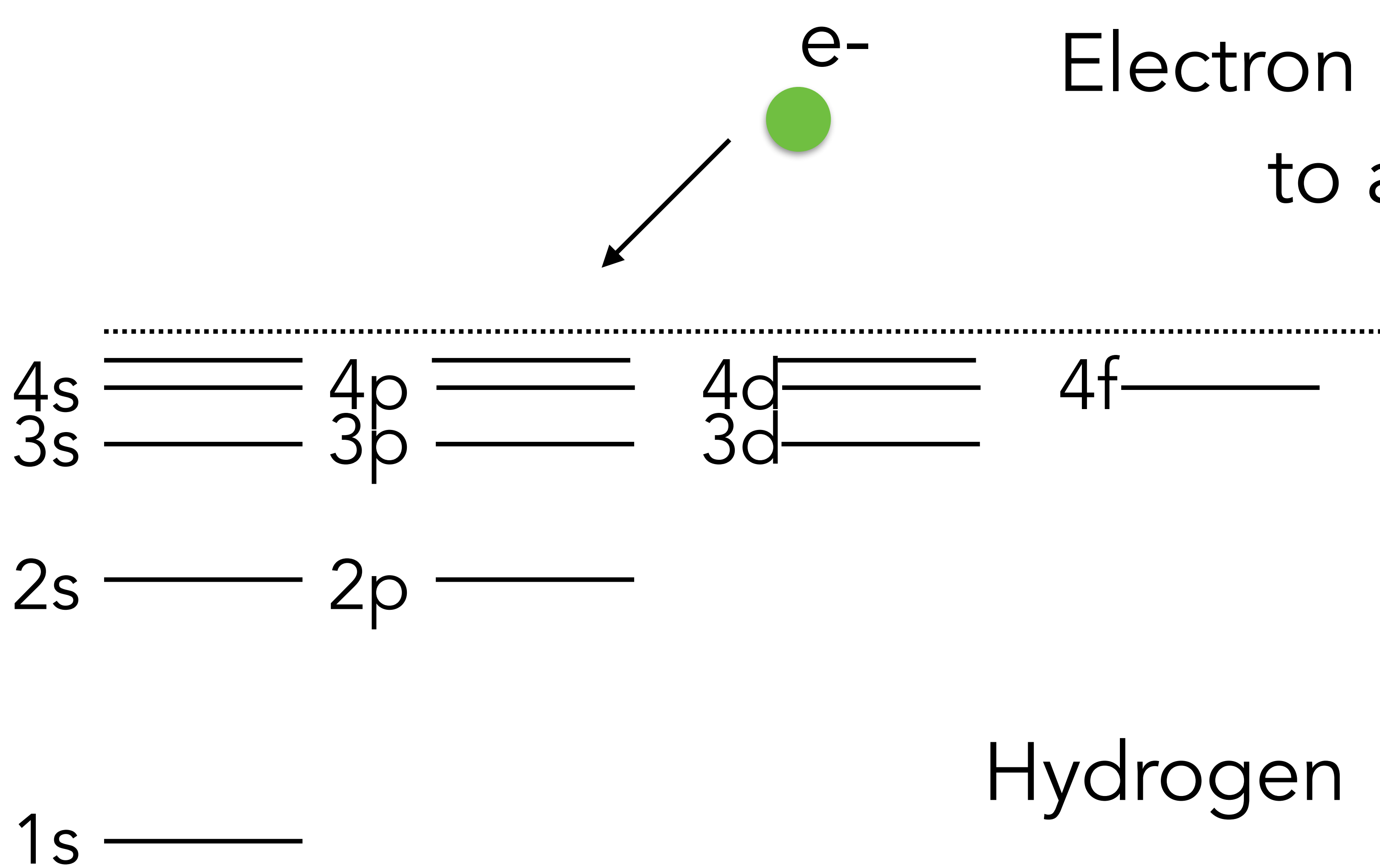


$I_{nl} = 13.6\ eV/n^2$

Hydrogen

Energy not to scale

Radiative Recombination

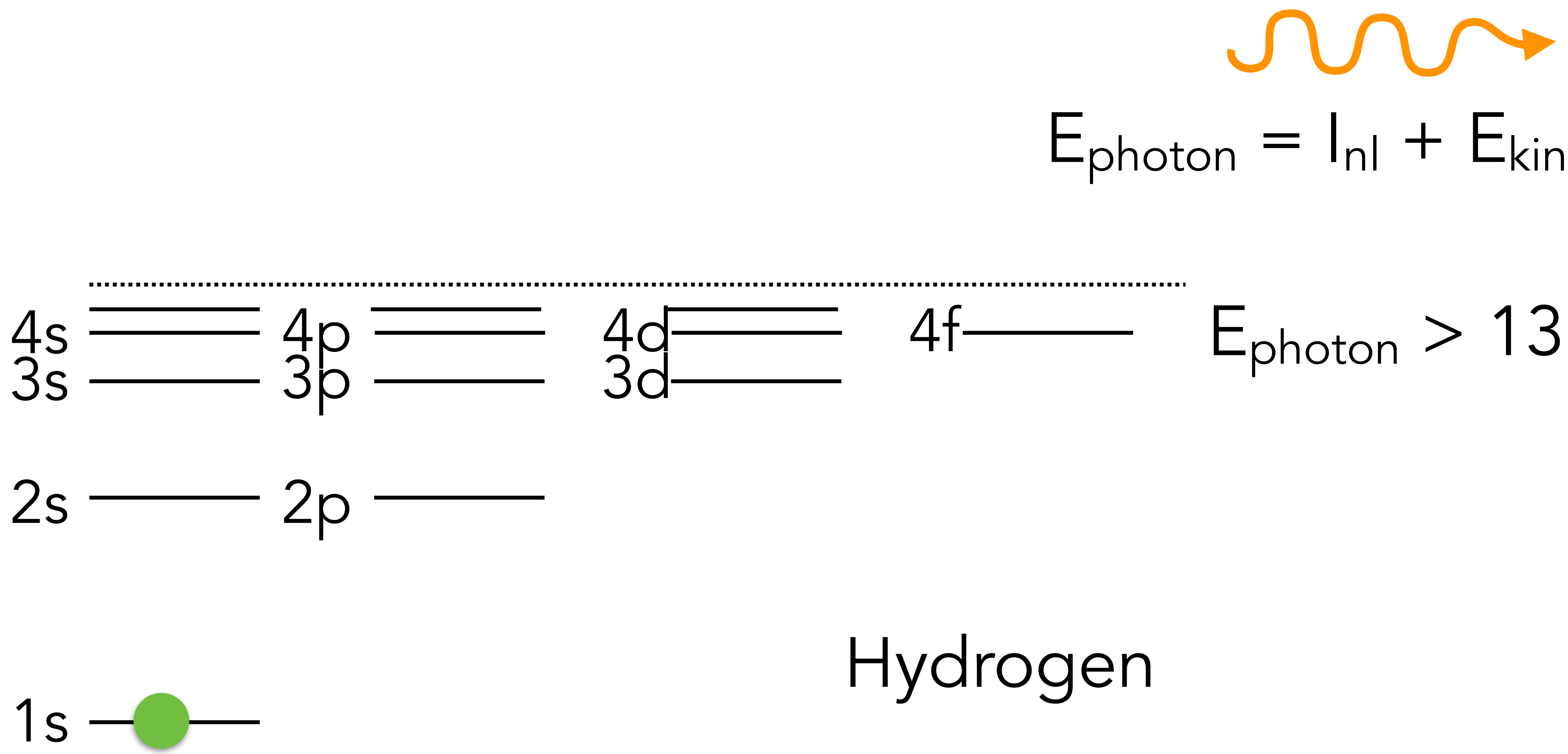


Electron can recombine to any level.

Hydrogen

Energy not to scale

Radiative Recombination



$E_{photon} = I_{nl} + E_{kinetic}$

$E_{photon} > 13.6$ eV

Hydrogen

Energy not to scale

Photon can ionize another H atom immediately if there is enough H around!

Radiative Recombination

"Case A": optically thin to ionizing radiation,
every ionizing photon from a recombination can escape

good approx for hot, collisionally ionized gas

"Case B": Optically thick to ionizing radiation,
recombinations to n=1 do not reduce ionization state of gas

good approx for "HII regions" =
photoionized nebulae around young, massive stars

Radiative Recombination

"Case A": optically thin to ionizing radiation, every ionizing photon from a recombination can escape

$$\alpha_A(T) = \sum_{n=1}^{\infty} \sum_{\ell=0}^{n-1} \alpha_{n\ell}(T)$$

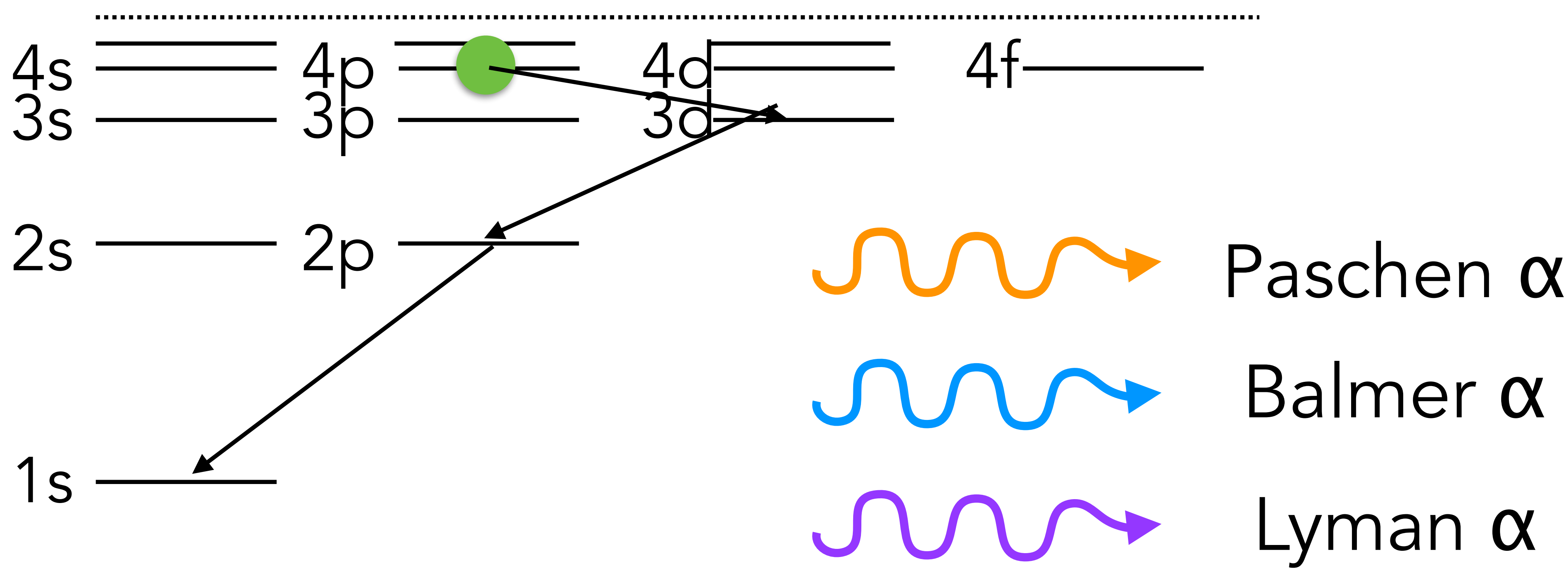
total recombination rate = sum of recombination rates to all levels

"Case B": Optically thick to ionizing radiation, recombinations to n=1 do not reduce ionization state of gas

$$\alpha_B(T) = \sum_{n=2}^{\infty} \sum_{\ell=0}^{n-1} \alpha_{n\ell}(T) = \alpha_A(T) - \alpha_{1s}(T)$$

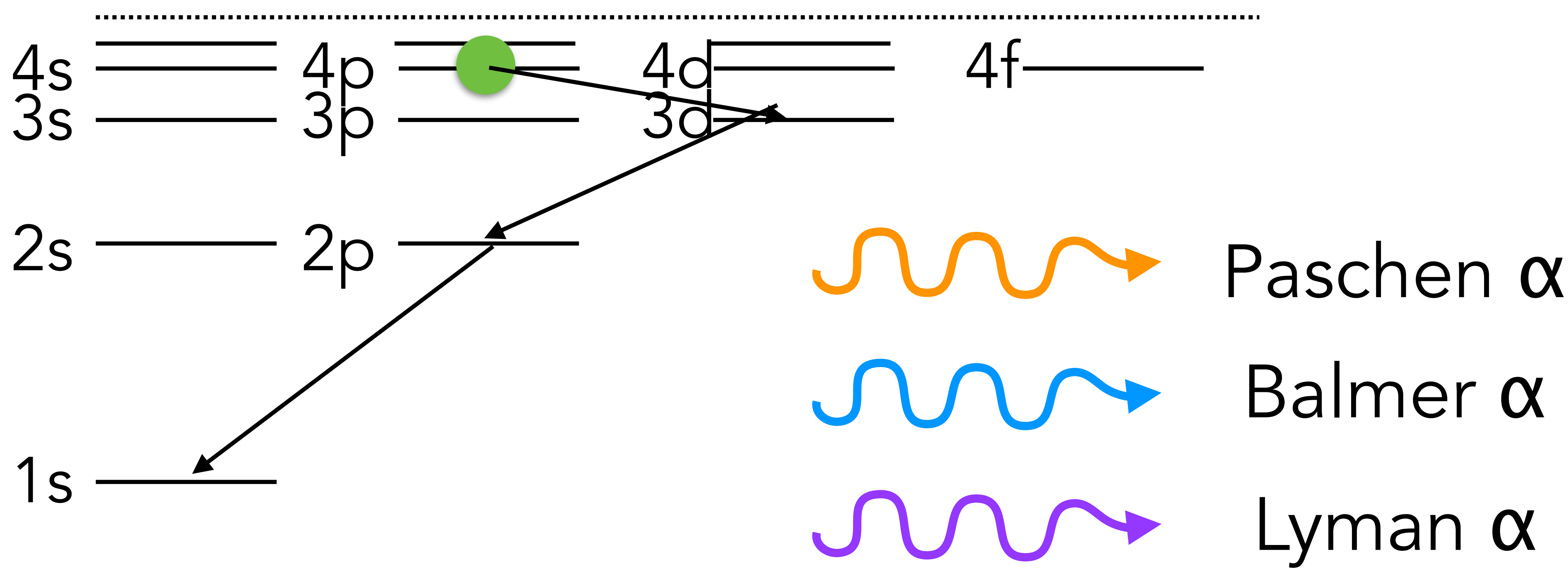
same but 1s rate is omitted

Radiative Recombination



For all but the highest n levels, collisions are much slower than radiative transitions -> recombination produces a characteristic spectrum of Hydrogen emission lines.

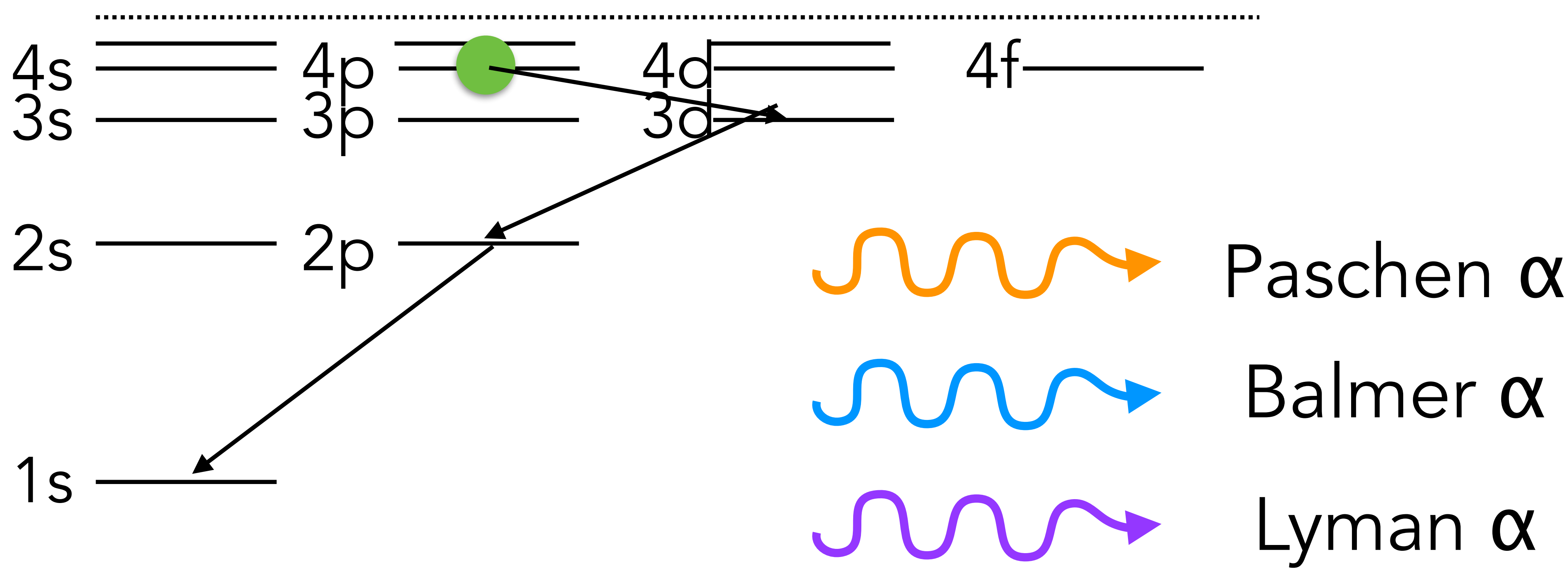
Radiative Recombination



allowed radiative decays for: $n>n'$ and $l-l'=\pm1$

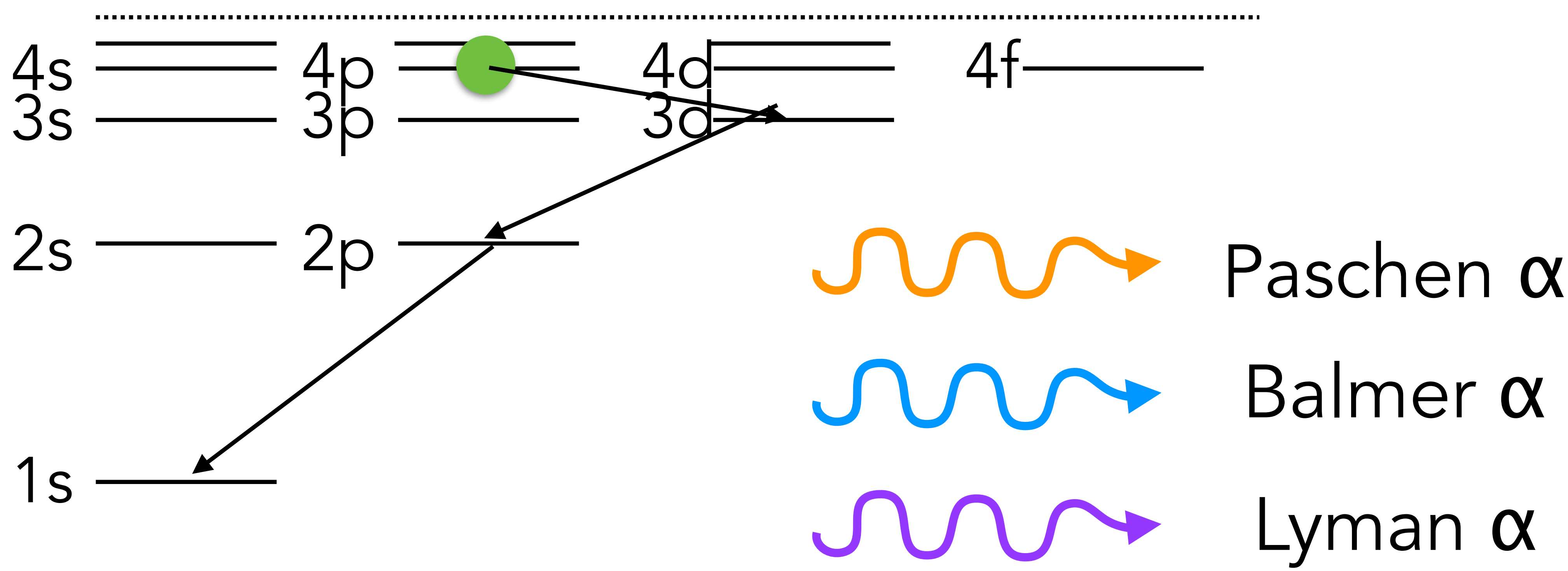
Einstein A coefficients + selection rules -> "branching ratios"

Radiative Recombination



For Case A this is straightforward.

Radiative Recombination

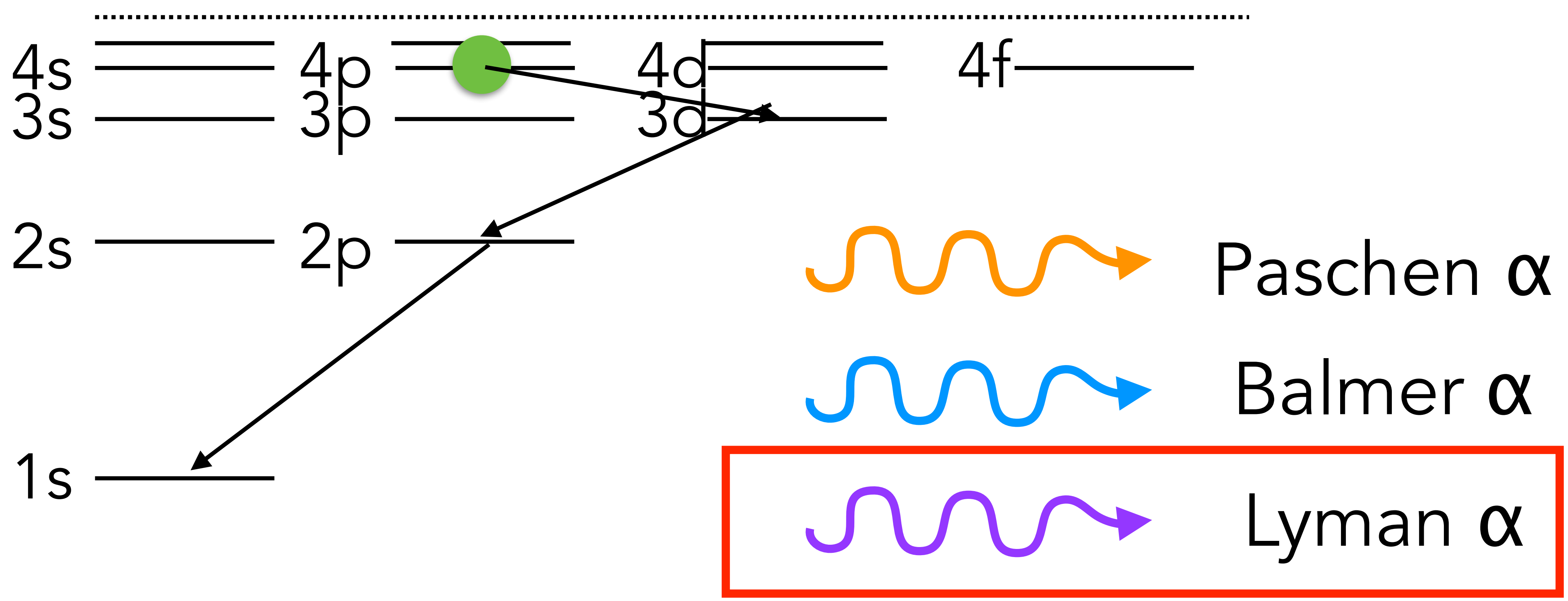


For Case B, need to recognize that cross section for Lyman transitions is big, bigger than even photoionization cross section.

for example:

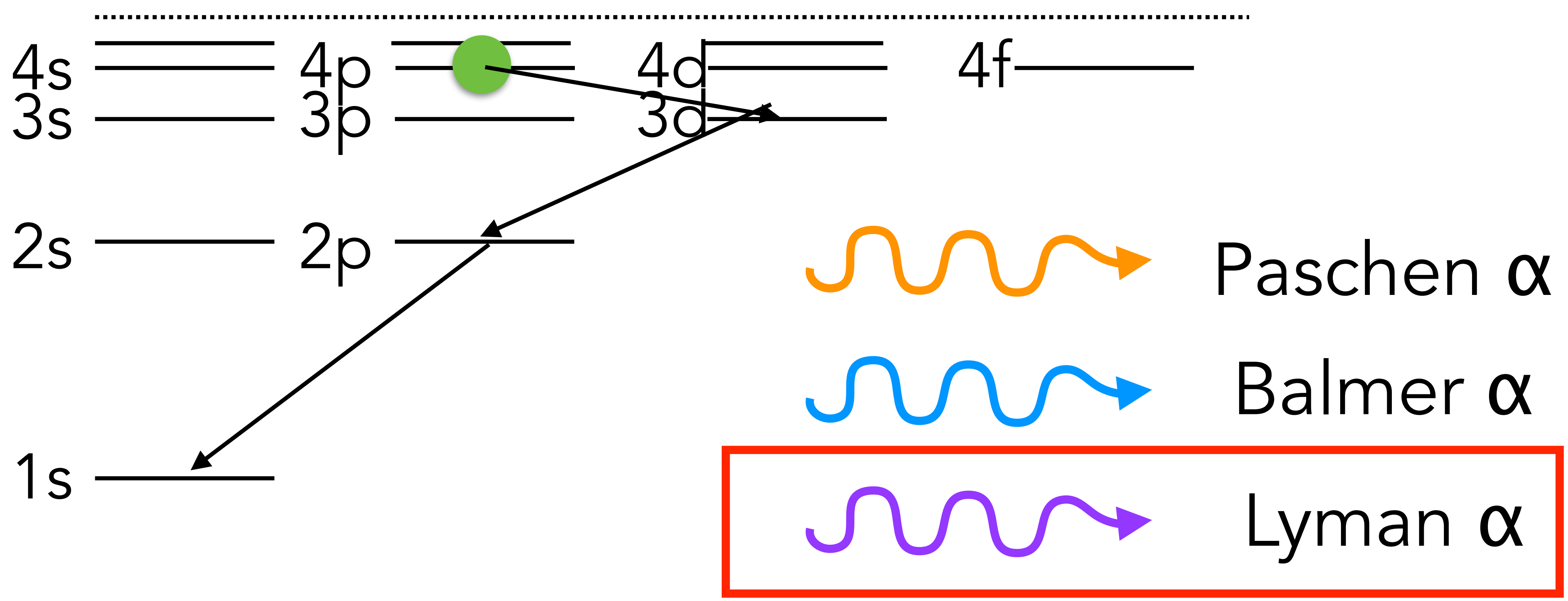
$$\tau_{Ly\alpha} = 8.0 \times 10^4 \left(\frac{15 \text{ km s}^{-1}}{b} \right) \tau_{LyC}$$

Radiative Recombination



Lyman photons will be absorbed immediately.
"resonantly scattered" with small changes in freq until a non-Lyman transition occurs

Radiative Recombination



Case B: rates for Lyman transitions -> 0
distributed instead among other transitions

Other Recombination Processes

- Dielectronic: capture of incoming electron excites one of the other bound electrons -> 2 excited e-
- Dissociative: molecular ion captures e-, dissociates
- Charge exchange: one important reaction is O^+ + H <-> O + H^+
- Neutralization by dust grains

Credit: NASA,ESA, M. Robberto (Space Telescope Science Institute/ESA)
and the Hubble Space Telescope Orion Treasury Project Team